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Comparing Monte Carlo simulations, mean particle theory estimates, and observations of H$^+$ and O$^+$ outflows at high altitudes and latitudes.

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Abstract. We carried out a comparison study between the results of Monte Carlo simulations, estimates of mean particle

10 theory, and available observations in different regions of earth magnetosphere (aurora, polar wind, central polar cap, and cusp) for H$^+$ and O$^+$ ions outflow at high latitudes and altitudes. We present altitude profiles for mean perpendicular energy W$_\perp$, mean parallel energy W$_\parallel$, and mean total energy W$_{total}$. Monte Carlo simulations are obtained by using Barghouthi model [Barghouthi, 2008], mean particle theory estimates are obtained by using Chang et al. [1986], and corresponding observations are obtained from different available publications. As a results of comparisons in different regions we have found that; 1)

15 Monte Carlo simulations and mean particle theory give similar results in auroral regions and produce no agreement in polar wind region, this is due to the strength of wave particle interaction which dominates the effects of external forces in aurora and competes with them in polar wind region, 2) using altitude dependent diffusion coefficients produce high energies, not reasonable, at middle and high altitudes, therefor it is recommended to use velocity and altitude diffusion coefficients, 3) comparison with observations in polar wind region and auroral region gives excellent agreement in aurora and good agreement

20 in polar wind, this is due to the implement of the appropriate velocity and altitude diffusion coefficient, 4) in the central polar cap and cusp we have obtained excellent agreement for both methods and observations, 5) due to the these comparisons we can claim that the wavelength of the electromagnetic wave existed in those regions (polar wind and aurora) is 8km and the altitude and velocity diffusion coefficients that have been used in Monte Carlo simulation and mean particle theory are appropriate to be used in different studies in these regions.

1 Introduction

Many studies (analytical, modelling, and observations) are developed to investigate the behaviour of ion outflow (e.g. O$^+$ and H$^+$) from polar regions of the planet Earth to outer space, and to find its characteristics such as velocity distribution, temperature, density, drift velocity and heat flux. Research and development processes are still going on to gather as much information as possible about the flow of these ions, including the energy and its components (parallel, perpendicular, and

30 total energy), which will be the topic of this research.

Different researchers have been investigating the ion outflow at high altitudes and latitudes. Chang et al. [1986] introduced the mean particle theory, which describes the perpendicular heating of ions in a dipole magnetic field. They

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Abstract. We conducted a comparative analysis of the results from Monte Carlo simulations, mean particle theory estimates,
10   and available observational data across different regions of Earth's magnetosphere, including the auroral, polar wind, central polar cap, and cusp regions, focusing on the outflow of H$^+$ and O$^+$ ions at high latitudes and altitudes. We present altitude profiles for the mean perpendicular energy W$_\perp$, mean parallel energy W$_\parallel$, and mean total energy W$_{total}$. The Monte Carlo simulations were carried out using the Barghouthi model [Barghouthi, 2008], while the mean particle theory estimates were derived from Chang et al. [1986], and the observational data were obtained from Nilsson et al. [2013] and Barghouthi et al.
15   [2016]. The results of the comparison across different regions reveal the following findings: 1) Monte Carlo simulations and mean particle theory yield similar results in the auroral regions but show discrepancies in the polar wind region. This discrepancy is attributed to the dominance of wave-particle interactions, which overshadow the effects of external forces (such as gravity, the polarization electric field, the mirror force, and centrifugal acceleration) in the auroral region and compete with them in the polar wind region. 2) The use of altitude-dependent diffusion coefficients leads to the generation of high-energy
20   particles, which are not consistent with the corresponding observational data at middle and high altitudes. Therefore, we recommend the use of velocity and altitude-dependent diffusion coefficients. 3) A comparison with observations in the auroral and polar wind regions demonstrates excellent agreement in the auroral region and good agreement in the polar wind region, attributed to the implementation of appropriate velocity and altitude-dependent diffusion coefficients. 4) In the central polar cap and cusp regions, both methods and observations show excellent agreement. 5) Based on these comparisons, we conclude
25   that the electromagnetic wave wavelength in the polar wind and auroral regions is approximately 8 km, and the velocity and altitude-dependent diffusion coefficients used in the Monte Carlo simulations and mean particle theory are suitable for application in further studies of these regions.

1 Introduction

Numerous studies, including analytical, modelling, and observational approaches, have been conducted to investigate the
30   behaviour of ion outflow (e.g., O$^+$ and H$^+$) from Earth's polar regions into outer space. These studies aim to characterize various properties of the ion outflow, such as velocity distribution, temperature, density, drift velocity, and heat flux. Ongoing

proposed that the intense broad band electric field fluctuations observed in the frequency range of (0-100 Hz) could be the cause of the transverse activation of ions through cyclonic resonance heating with left-handed polarized electromagnetic waves. Additionally, using a set of equations that described the motion of the ion in the geomagnetic field, they determined the parallel and perpendicular energies in accordance with the mean particle theory. Retterer et al. [1987] demonstrated how oxygen ions form conic distributions in the auroral zone of the Earth and used the diffusion equation to explain ion velocity distributions that obtained by using Monte Carlo method.

Barakat and Barghouthi [1994a, and b] upgraded the Monte Carlo simulation and investigated the impact of wave particle interaction (WPI) on H^+ and O^+ ion outflow in the polar wind. The electrostatic field, gravity, and geomagnetic field lines are considered in the model. These two studies considered as a parametric study because they used constant values for the quasilinear velocity diffusion rates along the simulation tube, the velocity distribution function and its velocity moments profiles were simulated and presented for both ions. Barghouthi [1997] and Barghouthi et al. [1998] processed the data obtained from plasma wave instrument (PWI) on board space dynamics explorer 1 and calculated the altitude dependence of the velocity diffusion rate, they presented the effect of altitude-dependent wave-particle interaction (WPI) on H^+ and O^+ ion outflow in the polar cap and auroral zone using Monte Carlo simulation. Additionally, Barghouthi [1997] investigated the model by comparing the energy estimates from the mean particle theory [Chang et al. 1986] with the corresponding energy results produced by Monte Carlo simulation in the auroral region. Despite the absence of supporting observations, there was strong agreement between the Monte Carlo simulations and the estimates of the mean particle theory.

Bouhram et al. [2002, 2003a, 2003b, 2004] developed a two-dimensional, Monte Carlo model for ion outflow from the dayside cusp/cleft, which is associated with transverse ion heating, they examined the cusp cleft region's transverse heating and ion outflow. They used their model to interpret the Cluster observations, i.e. saturation of transverse energization rate, in terms of finite perpendicular wavelength effects in the wave-particle interactions.

Barghouthi and Atout [2006] concentrated on the Monte Carlo simulations of toroidal H^+ and O^+ velocity distributions at high altitudes equatorward of the cusp by using a suitable form for velocity diffusion coefficient $D_\perp$. The results of the Monte Carlo simulations, toroidal H^+ and O^+ velocity distributions and H^+ and O^+ ion temperatures were compared to the toroidal H^+ and O^+ ion distributions and H^+ and O^+ ion temperatures that were observed at high altitudes equatorward of the cusp [Huddleston et al., 2000], these comparisons produced a reasonable agreement.

Barghouthi [2008] employed the Monte Carlo simulation to determine the temperatures and velocity distributions of H^+ and O^+ ions at high altitudes in the equatorward portion of the cusp by using different forms of the altitude and velocity dependent diffusion coefficient, $D_\perp(r, v_\perp)$ (RCC model, Bouhram model, and Barghouthi model), and compared between their simulation results with the corresponding observations of Huddleston et al. [2000], the simulation results of Barghouthi model were in excellent agreement with observations. In addition, Barghouthi [2008] presented much evidence, i.e., comparisons between Monte Carlo simulations results obtained by using Barghouthi model for H^+ and O^+ ions outflows with the corresponding observations obtained from different publications at different altitudes in the auroral region, that support Barghouthi model.

Waara et al. [2010] presented a case study of considerable heating of outflowing oxygen ions at high altitude (12 R_E) above the polar cap (up to 8 keV) perpendicular to the geomagnetic field. The distribution functions' shape suggests that the majority of the heating takes place locally (within 0.2–0.4 R_E of altitude). They discovered that it is unlikely that the locally observed wave fields can explain the observed ion energization because there are several events at lower altitudes. Furthermore, it is unlikely that the ions have migrated from an energizing location nearby to the observation site. This shows that at high

research efforts continue to enhance our understanding of the dynamics of these ion flows, with particular focus on the energy and its components (parallel, perpendicular, and total energy), which constitutes the primary subject of this study.

Various researchers have investigated ion outflow at high altitudes and latitudes. Chang et al. [1986] introduced the mean particle theory, which explains the perpendicular heating of ions in a dipole magnetic field. They suggested that intense broadband electric field fluctuations observed in the frequency range of 0-100 Hz could induce transverse ion activation through cyclonic resonance heating by left-handed polarized electromagnetic waves. Additionally, by employing a set of equations governing ion motion within the geomagnetic field, they derived expressions for the parallel and perpendicular energies based on the mean particle theory. Retterer et al. [1987] demonstrated how oxygen ions form conic distributions in the auroral zone and utilized the diffusion equation to explain ion velocity distributions obtained through the Monte Carlo method.

Barakat and Barghouthi [1994a, 1994b] enhanced the Monte Carlo simulation and explored the effects of wave-particle interactions (WPI) on H^+ and O^+ ion outflow in the polar wind. Their model incorporated the electrostatic field, gravity, and geomagnetic field lines. These studies are considered parametric investigations as they used constant values for the quasilinear velocity diffusion rates along the simulation tube. The velocity distribution function and its corresponding velocity moment profiles were simulated and presented for both ion species. Barghouthi [1997] and Barghouthi et al. [1998] analyzed data from the Plasma Wave Instrument (PWI) aboard the Space Dynamics Explorer 1, calculating the altitude dependence of the velocity diffusion rate. They examined the impact of altitude-dependent wave-particle interactions (WPI) on H^+ and O^+ ion outflow in both the polar cap and auroral zones using Monte Carlo simulations. Additionally, Barghouthi [1997] compared energy estimates from the mean particle theory (Chang et al., [1986]) with the corresponding energy results produced by Monte Carlo simulations in the auroral region. Despite the lack of supporting observational data, strong agreement was found between the Monte Carlo simulations and the mean particle theory estimates.

Bouhram et al. (2002, 2003a, 2003b, 2004) developed a two-dimensional Monte Carlo model to investigate ion outflow from the dayside cusp/cleft, focusing on the effects of transverse ion heating. They examined the mechanisms of transverse heating and ion outflow within the cusp/cleft region. Using their model, they interpreted Cluster satellite observations, specifically the saturation of the transverse energization rate, in terms of the influence of finite perpendicular wavelength effects in wave-particle interactions.

Barghouthi and Atout [2006] focused on Monte Carlo simulations of toroidal H^+ and O^+ velocity distributions at high altitudes, equatorward of the cusp, by employing an appropriate form for the velocity diffusion coefficient $D\perp$. The results of the Monte Carlo simulations, including the toroidal H^+ and O^+ velocity distributions as well as the H^+ and O^+ ion temperatures, were compared to the corresponding toroidal H^+ and O^+ ion distributions and ion temperatures observed at high altitudes, equatorward of the cusp [Huddleston et al., 2000]. These comparisons yielded reasonable agreement.

Barghouthi [2008] utilized Monte Carlo simulations to determine the temperatures and velocity distributions of H^+ and O^+ ions at high altitudes in the equatorward portion of the cusp, using various forms of altitude- and velocity-dependent diffusion

altitudes, additional, fundamentally distinct ion energization pathways exist. One explanation is that the ions' magnetic moment is not conserved, which would lead to slower outflow velocities and a longer ion energization period.

Waara et al. [2011] they provide the average values of coefficient which can be used to describing the diffusion in ion velocity at various altitude which can be consider a useful way to study ion outflow behaviour and their energies. The average energies of O^+ can be explained by the observed average wave in high altitudes range (8 - 15 R_E) in cusp and mantle regions according to their test particle calculations. They expected the relation between electric and magnetic field spectral density according to their results and the diffusion confection of O^+ increases with altitude.

Barghouthi et al. [2012] compared the simulation results of ion outflow (density, drift velocity, perpendicular and parallel temperatures, and ion velocity distributions at different altitudes) in two different regions, polar wind and auroral region, based on the Barghouthi model. They found that wave particle interactions have a greater impact in the auroral zone than they do in the polar wind region, and that they have a greater impact on the energizing of O^+ ions than H^+ ions.

Barghouthi et al. [2016] updated the Monte Carlo model by taking into account the effects of gravity, ambipolar electric field, centrifugal acceleration, mirror force and wave particle interactions to study O^+ and H^+ ions outflow above the polar cap, they changed various parameters like (centrifugal acceleration, velocity diffusion coefficients, and boundary conditions at lower-altitude), they compared their results with the observations obtained by different instruments on board Cluster spacecraft, their results were in good agreement with observed data.

In this study, our main objective is to compare between the simulation results (perpendicular energies $W_\perp$, parallel energies $W_\parallel$, and the total energies W_{total}) of O^+ and H^+ ions obtained by using Monte Carlo model (Barghouthi model) and mean particle theory, and available observations in different regions of earth magnetosphere (polar wind, auroral region, cusp, and central polar cap).

2 Formulations

2.1 Monte Carlo simulation

The motion of the plasma's constituents and interactions between plasma species can be precisely defined via Monte Carlo simulation. When working with plasma, it is practical to characterize each species using a different velocity distribution function, f_s ($\mathbf{v}_s$, $\mathbf{r}_s$, t). The velocity distribution function is defined such that f_s ($\mathbf{v}_s$, $\mathbf{r}_s$, $t)d\mathbf{v}_s d\mathbf{r}_s$ represents the number of particles of species s which at time t have velocity between $\mathbf{v}_s$ and $\mathbf{v}_s + d\mathbf{v}_s$ and positions between $\mathbf{r}_s$ and $\mathbf{r}_s + d\mathbf{r}_s$. The net result of collisions and interactions, and the movement of species in phase space under the influence of external forces define the evolution of the species velocity distribution function throughout time [Schunk, 1977]. The well-known Boltzmann equation provides a mathematical representation of this evolution:

$$\frac{\partial f_s}{\partial t} + \mathbf{v}_s . \nabla f_s + \left(\frac{e_s}{m_s}\right)\left[\boldsymbol{E} + \frac{\mathbf{v_s} \times \mathbf{B}}{c}\right] . \nabla_{v_s} f_s = \frac{\delta f_s}{\delta t} \qquad (1)$$

In this equation, the left-hand side represents the evolution of the velocity distribution function f_s ($\mathbf{v}_s$, $\mathbf{r}_s$, t) under the effects of external forces, and the right-hand side represents the Boltzmann collision integral, here it represents the rate at which f_s changed as a result of wave particle interactions in the region of study. In the above equation, E is the polarization electric field, B is the geomagnetic field, c is the speed of light, ∇ is the coordinate space gradient, and ∇_{v_s} is the velocity space

65   coefficients $D_\perp(r, v_\perp)$, including the RCC model, Bouhram model, and Barghouthi model. The simulation results were compared with the corresponding observations from Huddleston et al. [2000], and the results from the Barghouthi model showed excellent agreement with the observations. Furthermore, Barghouthi [2008] provided substantial evidence, including comparisons between Monte Carlo simulation results for H^+ and O^+ ion outflows obtained using the Barghouthi model and corresponding observational data from multiple sources at various altitudes in the auroral region, further supporting the validity

70   of the Barghouthi model.

Waara et al. [2010] presented a case study of significant heating of outflowing oxygen ions at high altitudes (12 R_E) above the polar cap, with energies reaching up to 8 keV perpendicular to the geomagnetic field. The shape of the distribution functions suggests that the majority of the heating occurs locally, within 0.2–0.4 R_E of altitude. They found that the locally observed wave fields are unlikely to account for the observed ion energization. Additionally, it is improbable that the ions originated

75   from a nearby energizing location and migrated to the observation site. These findings indicate the existence of additional, fundamentally distinct ion energization mechanisms at high altitudes. One possible explanation is that the ions' magnetic moment is not conserved, which would result in slower outflow velocities and an extended period of ion energization.

Waara et al. [2011, 2012] conducted a statistical analysis of ion heating and associated wave activity. They provided average values for coefficients that describe diffusion in ion velocity space at various altitudes, offering a useful framework for studying

80   ion outflow behaviour and energy characteristics. Their test particle calculations suggest that the average energies of O^+ ions correlate with the observed wave activity at high altitudes (8-15 R_E) in the cusp and mantle regions. They also found that the electric-to-magnetic field spectral density ratios closely match the expected values for Alfven waves. According to their findings, the diffusion coefficient for O^+ ions increases with altitude.

Barghouthi et al. [2012] compared simulation results of ion outflow (including ion density, drift velocity,

85   perpendicular and parallel temperatures, and ion velocity distributions at various altitudes) in two distinct regions, the polar wind and auroral regions, based on the Barghouthi model. They found that wave-particle interactions had a more significant effect in the auroral zone compared to the polar wind region and that these interactions had a greater influence on the energization of O^+ ions than on H^+ ions.

Barghouthi et al. [2016] updated the Monte Carlo model to include the effects of gravity, ambipolar electric field,

90   centrifugal acceleration, mirror force, and wave-particle interactions in the study of O^+ and H^+ ion outflow above the polar cap. They modified several parameters, such as centrifugal acceleration, velocity diffusion coefficients, and boundary conditions at lower altitudes. The results were compared with observational data obtained from various instruments aboard the Cluster spacecraft, and the simulation results were found to be in good agreement with the observed data.

The primary objective of this study is to compare the simulation results (perpendicular energies $W_\perp$, parallel energies

95    $W_\parallel$, and total energies W_{total}) of O^+ and H^+ ions obtained using the Monte Carlo model (Barghouthi model) and mean particle theory with available observational data from various regions of Earth's magnetosphere, including the polar wind, auroral

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gradient, and e_s and m_s are the charge and mass of species s, respectively. The suitable expression for $\left(\frac{\delta f_s}{\delta t}\right)$ in case of wave particle interactions is given by Retterer et al. [1987], they considered the effects of (WPI) as particles diffusion in the velocity space.

$$\frac{\delta f}{\delta t}\bigg|_{WPI} = \left(\frac{1}{v_\perp}\right)\frac{\partial}{\partial v_\perp}\left[D_\perp \, v_\perp \, \frac{\partial f}{\partial v_\perp}\right] \tag{2}$$

110 where $D_\perp$ is provided by Retterer et al. [1987] and represents the quasi-linear velocity diffusion rate perpendicular to the geomagnetic field,

$$D_\perp = (\eta \, q^2/4m^2)|E_x(\omega = \Omega)|^2 \tag{3}$$

where $|E_x(\omega)|^2$ is the measured spectral density of the electromagnetic turbulence, η is the proportion of the measured spectral density by plasma wave instrument (PWI) on board dynamic explorer 1 (DE-1) spacecraft that corresponds to the left-hand 115 polarized wave, q is the ion's charge, m is the ion's mass, Ω is the ion's gyrofrequency, and ω is the angular frequency of the electromagnetic turbulence.

The expression of the velocity diffusion rate $D_\perp$ as given in Eq. (3) is independent of velocity and depends on position (altitude) via changes in the ion gyrofrequency, Ω, along the geomagnetic field lines. By examining experimental data of electric field spectral density obtained by plasma wave instrument (PWI) onboard the DE-1 satellite (i.e. for high solar activity conditions), 120 Barghouthi [1997] and Barghouti et al. [1998] calculated the altitude dependence of ($D_\perp$). They came up with the following expressions for the velocity diffusion coefficient $D_\perp$ in the polar wind region [Barghouthi et al., 1998] as follows:

$$D_\perp(r) = \begin{cases} 5.77 \times 10^3 \left(\frac{r}{R_E}\right)^{7.95} cm^2 s^{-3}, \text{for } H^+ \\ 9.55 \times 10^2 \left(\frac{r}{R_E}\right)^{13.3} cm^2 s^{-3}, \text{for } O^+ \end{cases} \tag{4}$$

In the auroral region, $D_\perp(r)$ is given by Barghouthi [1997] as follows:

$$D_\perp(r) = \begin{cases} 4.45 \times 10^7 \left(\frac{r}{R_E}\right)^{7.95} cm^2 s^{-3}, \text{for } H^+ \\ 6.94 \times 10^5 \left(\frac{r}{R_E}\right)^{13.3} cm^2 s^{-3}, \text{for } O^+ \end{cases} \tag{5}$$

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In central polar cap (CPC) and cusp regions, $D_\perp(r)$ is given by Nilsson et al. [2013] as follows:

For central polar cape region

$$D_\perp(r) = \begin{cases} 20 \left(\frac{r}{R_E}\right)^{9.77} cm^2 s^{-3}, \text{for } H^+ \\ 0.5 \times 10^5 \left(\frac{r}{R_E}\right)^{5.5} cm^2 s^{-3}, \text{for } O^+ \end{cases} \tag{6}$$

and for cusp region

region, cusp, and central polar cap. This comparison seeks to evaluate the importance of including altitude-dependent velocity diffusion coefficients $D_\perp(r)$ or altitude-and velocity-dependent velocity diffusion coefficients $D_\perp(r, v_\perp)$, while also considering the constraints posed by wavelength limitations.

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2 Formulations

2.1 Monte Carlo simulation

In the study of space plasma, it is advantageous to describe each constituent species utilizing distinct velocity distribution function, denoted as $f_s(\mathbf{v}_s, \mathbf{r}_s, t)$. The velocity distribution function is defined such that the expression $f_s(\mathbf{v}_s, \mathbf{r}_s, t) d\mathbf{v}_s d\mathbf{r}_s$

105   quantifies the number of particles of species s that, at time t, possess velocities within the interval $\mathbf{v}_s$ and $\mathbf{v}_s + d\mathbf{v}_s$ and positions between $\mathbf{r}_s$ and $\mathbf{r}_s + d\mathbf{r}_s$. The temporal evolution of the species' velocity distribution function is governed by the cumulative effects of collisions and interactions, as well as the dynamical movements of the species in phase space under the influence of external forces [Schunk, 1977]. This evolution can be mathematically represented by the well-established Boltzmann equation:

$$\frac{\partial f_s}{\partial t} + \mathbf{v}_s . \nabla f_s + \left(\frac{e_s}{m_s}\right)\left[\boldsymbol{E} + \frac{\mathbf{v}_s \times \mathbf{B}}{c}\right] . \nabla_{v_s} f_s = \frac{\delta f_s}{\delta t} \tag{1}$$

110   In this equation, the left-hand side represents the evolution of the velocity distribution function $f_s(\mathbf{v}_s, \mathbf{r}_s, t)$ under the effects of external forces, and the right-hand side represents the Boltzmann collision integral, here it represents the rate at which f_s changed as a result of wave particle interactions in the region of study. In the above equation, E is the polarization electric field, B is the geomagnetic field, c is the speed of light, ∇ is the coordinate space gradient, and ∇_{v_s} is the velocity space gradient, and e_s and m_s are the charge and mass of species s, respectively. The suitable expression for $\left(\frac{\delta f_s}{\delta t}\right)$ in case of wave

115   particle interactions is given by Retterer et al. [1987], they considered the effects of (WPI) as particles diffusion in the velocity space.

$$\left.\frac{\delta f}{\delta t}\right|_{WPI} = \left(\frac{1}{v_\perp}\right)\frac{\partial}{\partial v_\perp}\left[D_\perp \, v_\perp \, \frac{\partial f}{\partial v_\perp}\right] \tag{2}$$

where $D_\perp$ is provided by Retterer et al. [1987] and represents the quasi-linear velocity diffusion rate perpendicular to the geomagnetic field,

120   $$D_\perp = (\eta \, q^2/4m^2)|E_x(\omega = \Omega)|^2 \tag{3}$$

where $|E_x(\omega)|^2$ is the measured spectral density of the electromagnetic turbulence, η is the proportion of the measured spectral density by plasma wave instrument (PWI) on board dynamic explorer 1 (DE-1) spacecraft that corresponds to the left-hand polarized wave, q is the ion's charge, m is the ion's mass, Ω is the ion's gyrofrequency, and ω is the angular frequency of the electromagnetic turbulence.

125   The expression of the velocity diffusion rate $D_\perp$ as given in Eq. (3) is independent of velocity and depends on position (altitude) via changes in the ion gyrofrequency, Ω, along the geomagnetic field lines. By examining experimental data of electric field spectral density obtained by plasma wave instrument (PWI) onboard the DE-1 satellite (i.e. for high solar activity conditions),

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$$130 \quad D_\perp (r) = \begin{cases} 1.01 \times 10^6 \left(\frac{r}{R_E}\right)^{5.61} cm^2 s^{-3}, \text{for } H^+ \\ 2.5 \times 10^4 \left(\frac{r}{R_E}\right)^{6.4} cm^2 s^{-3}, \text{for } O^+ \end{cases} \qquad (7)$$

The diffusion coefficient was given a new form by Barghouthi [2008], who discovered that it is a velocity-dependent in addition to altitude-dependent.

$$D_\perp (r, v_\perp) = D_\perp (r) \begin{cases} 1 & \text{for } \left(\frac{k_\perp v_\perp}{\Omega_i}\right) < 1 \\ \left(\frac{k_\perp v_\perp}{\Omega_i}\right)^{-3} & \text{for } \left(\frac{k_\perp v_\perp}{\Omega_i}\right) \geq 1 \end{cases} \qquad (8)$$

135 Where $D_\perp (r, v_\perp)$ is the quasi-linear velocity diffusion rate perpendicular to the geomagnetic field lines (altitude and velocity dependent), Ω_i is the ion gyrofrequency and $k_\perp$ is perpendicular wave number and related to the characteristic perpendicular wavelength of the electromagnetic turbulence $\lambda_\perp$.

By solving Boltzmann equation, Eq. (1), using Monte Carlo technique the velocity distribution functions were obtained for each species (in this study, O^+ and H^+ ions) and its velocity moments, i.e. density n_s, drift velocity u_s, and parallel $T_{s\parallel}$ and perpendicular $T_{s\perp}$ temperatures. The moments considered here are defined as follows [Barghouthi, 1997]:

$$140 \quad n_s = \int f_s d\mathbf{v}_s \qquad (9)$$

$$u_s = \frac{1}{n_s} \int v_{s\parallel} f_s d\mathbf{v}_s \qquad (10)$$

$$T_{s\parallel} = \frac{m_s}{n_s k} \int \left(v_{s\parallel} - u_s\right)^2 f_s d\mathbf{v}_s \qquad (11)$$

$$T_{s\perp} = \frac{m_s}{2n_s k} \int (v_{s\perp})^2 f_s d\mathbf{v}_s \qquad (12)$$

These Monte Carlo results will be used to calculate the mean parallel energy, mean perpendicular energy, and total mean 145 energy as given in the following expressions [Barghouthi, 1997], respectively:

$$W_{s\parallel} = \frac{1}{2} m u_s^2 + \frac{1}{2} k T_{s\parallel} \qquad (13)$$

$$W_{s\perp} = k T_{s\perp} \qquad (14)$$

$$W_s = W_{s\parallel} + W_{s\perp} \qquad (15)$$

Where u_s, $T_{s\parallel}$ and $T_{s\perp}$ are given by equations (10), (11) and (12), respectively and $W_{s\parallel}$ and $W_{s\perp}$ are the mean parallel and 150 perpendicular energies, respectively; W_s is the total mean energy; and s denotes the type of the ion (O^+ or H^+), k is Boltzmann constant.

Barghouthi [1997] and Barghouti et al. [1998] calculated the altitude dependence of $(D_\perp)$. They came up with the following expressions for the velocity diffusion coefficient $D_\perp$ in the polar wind region [Barghouthi et al., 1998] as follows:

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$$D_\perp(r) = \begin{cases} 5.77 \times 10^3 \left(\frac{r}{R_E}\right)^{7.95} cm^2 s^{-3}, \text{for } H^+ \\ 9.55 \times 10^2 \left(\frac{r}{R_E}\right)^{13.3} cm^2 s^{-3}, \text{for } O^+ \end{cases} \tag{4}$$

In the auroral region, $D_\perp(r)$ is given by Barghouthi [1997] as follows:

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In central polar cap (CPC) and cusp regions, $D_\perp(r)$ is given by Nilsson et al. [2013] as follows:

135 For central polar cape region

$$D_\perp(r) = \begin{cases} 20 \left(\frac{r}{R_E}\right)^{9.77} cm^2 s^{-3}, \text{for } H^+ \\ 0.5 \times 10^5 \left(\frac{r}{R_E}\right)^{5.5} cm^2 s^{-3}, \text{for } O^+ \end{cases} \tag{6}$$

and for cusp region

$$D_\perp(r) = \begin{cases} 1.01 \times 10^6 \left(\frac{r}{R_E}\right)^{5.61} cm^2 s^{-3}, \text{for } H^+ \\ 2.5 \times 10^4 \left(\frac{r}{R_E}\right)^{6.4} cm^2 s^{-3}, \text{for } O^+ \end{cases} \tag{7}$$

The diffusion coefficient was given a new form by Barghouthi [2008], who discovered that it is a velocity-dependent in

140 addition to altitude-dependent.

$$D_\perp(r, v_\perp) = D_\perp(r) \begin{cases} 1 & \text{for } \left(\frac{k_\perp v_\perp}{\Omega_i}\right) < 1 \\ \left(\frac{k_\perp v_\perp}{\Omega_i}\right)^{-3} & \text{for } \left(\frac{k_\perp v_\perp}{\Omega_i}\right) \geq 1 \end{cases} \tag{8}$$

Where $D_\perp(r, v_\perp)$ is the quasi-linear velocity diffusion rate perpendicular to the geomagnetic field lines (altitude and velocity dependent), Ω_i is the ion gyrofrequency and $k_\perp$ is perpendicular wave number and related to the characteristic perpendicular wavelength of the electromagnetic turbulence $\lambda_\perp$. Equation (8) indicates the diffusion coefficient that is dependent on altitude

145 and velocity. However, as ions are heated and move to higher altitudes, their gyroradius may approach the perpendicular wavelength of the electromagnetic turbulence, and when the ratio $(k_\perp v_\perp/\Omega_i)$ exceeds 1, the heating rate becomes self-limited. Bouhram et al. [2004] derived an alternative form for the altitude and velocity dependent diffusion coefficient and interpreted their results in terms of finite wavelength effects.

By solving Boltzmann equation, Eq. (1), using Monte Carlo technique the velocity distribution functions were obtained for

150 each species (in this study, O^+ and H^+ ions) and its velocity moments, i.e. density n_s, drift velocity u_s, and parallel $T_{s\parallel}$ and perpendicular $T_{s\perp}$ temperatures. The moments considered here are defined as follows [Barghouthi, 1997]:

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2.2 Barghouthi model

Barghouthi model was developed to study the behaviour of ions (H^+ and O^+) outflow at high altitudes and high-latitudes, the simulation results of this model provide an excellent agreement to observations in different regions, auroral region [Barghouthi, 2008] and polar wind region [Barghouthi et al., 2011]. This model simulates different effects (gravity, polarization electric field, diverging geomagnetic field, and wave particle interactions which acts on (H^+ and O^+) ion outflow at high altitudes and high latitudes. This model emphasises the significant role of wave particle interactions that is responsible for ion heating, this effect depends on the velocity diffusion coefficient $D_\perp (r, v_\perp)$, they developed a form for this coefficient as a function of position (r/R_E) along geomagnetic field lines of the Earth and injected ion perpendicular velocity $(v_\perp)$. The different forms of velocity diffusion coefficient, $D_\perp (r, v_\perp)$ have been used in the Monte Carlo simulation to obtain O^+ and H^+ ions temperatures and velocity distribution functions at high altitudes and latitudes. In the Monte Carlo simulation, the appropriate velocity diffusion coefficient related to the specific region of study, $D_\perp (r, v_\perp)$, were used to determine the temperatures and velocity distributions of H^+ and O^+ ions at high altitudes and latitudes.

2.3 Mean particle theory

Chang et al. [1986] provided a theory for estimation of the values of the mean perpendicular, mean parallel, and total mean energies as a function of geocentric distance by including the average rate of heating for each ion in a set of equations that describe the motion of the ion along geomagnetic field lines, as follows:

$$W_{i\parallel} = \frac{9m_i}{2^{1/3}} \left[\frac{r D_\perp (r)}{(3\alpha+1)(6\alpha+11)} \right]^{2/3} \tag{16}$$

$$W_{i\perp} = \frac{(6\alpha+2)m_i}{2^{1/3}} \left[\frac{r D_\perp (r)}{(3\alpha+1)(6\alpha+11)} \right]^{2/3} \tag{17}$$

$$W_i = W_{i\parallel} + W_{i\perp} = (3\alpha + 11/2)^{1/3} m_i \left[\frac{r D_\perp (r)}{(3\alpha+1)} \right]^{2/3} \tag{18}$$

Where $W_\parallel$ and $W_\perp$ are the mean parallel and perpendicular energies, respectively, W_i is the total mean energy, i denotes the type of the ion (H^+ or O^+), and α is a fitting parameter. In that theory the mean energy ratio $W_\perp / W_\parallel$ asymptotically approaches a constant value.

3 Comparisons

Here, we will do comparison between Monte Carlo simulations and mean particle theory estimates in auroral, and polar wind regions, sub-section 3.1, and comparison between Monte Carlo simulations, mean particle theory estimates, and corresponding observations, sub-section 3.2.

3.1 Comparison between Monte Carlo simulations and estimates of mean particle theory.

By using Monte Carlo technique (i.e. Barghouthi model) and mean particle theory, we have obtained different altitude profiles for mean parallel, mean perpendicular, and total mean energies in different regions of earth magnetosphere, auroral region Fig. 1, and polar wind region Fig. 2, for H^+ (left panels) and O^+ (right panels) ions outflow. In both methods we have

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$$n_s = \int f_s d\mathbf{v}_s \tag{9}$$

$$u_s = \frac{1}{n_s} \int v_{s\parallel} f_s d\mathbf{v}_s \tag{10}$$

$$T_{s\parallel} = \frac{m_s}{n_s k} \int (v_{s\parallel} - u_s)^2 f_s d\mathbf{v}_s \tag{11}$$

$$T_{s\perp} = \frac{m_s}{2n_s k} \int (v_{s\perp})^2 f_s d\mathbf{v}_s \tag{12}$$

These Monte Carlo results will be used to calculate the mean parallel energy, mean perpendicular energy, and total mean energy as given in the following expressions [Barghouthi, 1997], respectively:

$$W_{s\parallel} = \frac{1}{2} m u_s^2 + \frac{1}{2} k T_{s\parallel} \tag{13}$$

$$W_{s\perp} = k T_{s\perp} \tag{14}$$

$$W_s = W_{s\parallel} + W_{s\perp} \tag{15}$$

Where u_s, $T_{s\parallel}$ and $T_{s\perp}$ are given by equations (10), (11) and (12), respectively and $W_{s\parallel}$ and $W_{s\perp}$ are the mean parallel and perpendicular energies, respectively; W_s is the total mean energy; and s denotes the type of the ion (O^+ or H^+), k is Boltzmann constant.

2.2 Barghouthi model

The Barghouthi model was developed to investigate the behavior of ion outflows, specifically for H^+ and O^+ ions, at high altitudes and high latitudes. The simulation results generated by this model show excellent agreement with observational data from various regions, including the auroral region [Barghouthi, 2008] and the polar wind region [Barghouthi et al., 2011]. This model accounts for multiple influential factors, including gravity, the polarization electric field, the diverging geomagnetic field, and wave-particle interactions, all of which affect H^+ and O^+ ion outflows at elevated altitudes and latitudes. Notably, the model highlights the significant contribution of wave-particle interactions, which are responsible for ion heating. The impact of these wave-particle interactions is characterized by the velocity diffusion coefficient $D_\perp (r, v_\perp)$, which has been formulated as a function of position (r/R_E) along Earth's geomagnetic field lines and the perpendicular velocity of the injected ions $(v_\perp)$. Various functional forms of the velocity diffusion coefficient $D_\perp (r, v_\perp)$ have been employed in Monte Carlo simulations to derive the density, drift velocity, parallel and perpendicular temperatures, heat fluxes, and velocity distribution functions for H^+ and O^+ ions at high altitudes and latitudes. In the Monte Carlo simulations of this study, the appropriate velocity diffusion coefficient corresponding to the specific study region, $D_\perp (r, v_\perp)$, was utilized to ascertain the temperatures and velocity distributions of H^+ and O^+ ions at these elevated altitudes and latitudes.

2.3 Mean particle theory

Chang et al. [1986] presented a theoretical framework for estimating the mean perpendicular, mean parallel, and total mean energies as functions of geocentric distance. This framework incorporates the average heating rate for each ion into a set of equations that describes the motion of ions along geomagnetic field lines, as outlined below:

$$W_{i\parallel} = \frac{9m_i}{2^{1/3}} \left[\frac{rD_\perp(r)}{(3\alpha+1)(6\alpha+11)} \right]^{2/3} \tag{16}$$

$$W_{i\perp} = \frac{(6\alpha+2)m_i}{2^{1/3}} \left[\frac{rD_\perp(r)}{(3\alpha+1)(6\alpha+11)} \right]^{2/3} \tag{17}$$

$$W_i = W_{i\parallel} + W_{i\perp} = (3\alpha + 11/2)^{1/3} m_i \left[\frac{rD_\perp(r)}{(3\alpha+1)} \right]^{2/3} \tag{18}$$

185   Where $W_\parallel$ and $W_\perp$ are the mean parallel and perpendicular energies, respectively, W_i is the total mean energy, i denotes the type of the ion (H^+ or O^+), and α is a fitting parameter. In that theory the mean energy ratio $W_\perp / W_\parallel$ asymptotically approaches a constant value.

3 Comparisons

In this section, we will compare Monte Carlo simulations with mean particle theory estimates in the auroral and polar wind
190   regions (sub-section 3.1). Additionally, we will present a comparison among Monte Carlo simulations, mean particle theory estimates, and relevant observational data in sub-section 3.2.

3.1 Comparison between Monte Carlo simulations and estimates of mean particle theory.

By employing the Monte Carlo technique (specifically the Barghouthi model) alongside mean particle theory, we have derived varying altitude profiles for the mean parallel, mean perpendicular, and total mean energies across different regions of Earth's
195   magnetosphere—specifically, the auroral region (Figure 1) and the polar wind region (Figure 2)—for H^+ (left panels) and O^+ (right panels) ions outflow. In both methods, we utilized an appropriate $D_\perp(r)$ as the velocity diffusion coefficient for each region. Overall, our findings in the auroral region (Figure 1) indicate an excellent agreement between

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used suitable $D_\perp(r)$ for velocity diffusion coefficient in each region. In general, we have found in the auroral region (figure 1)
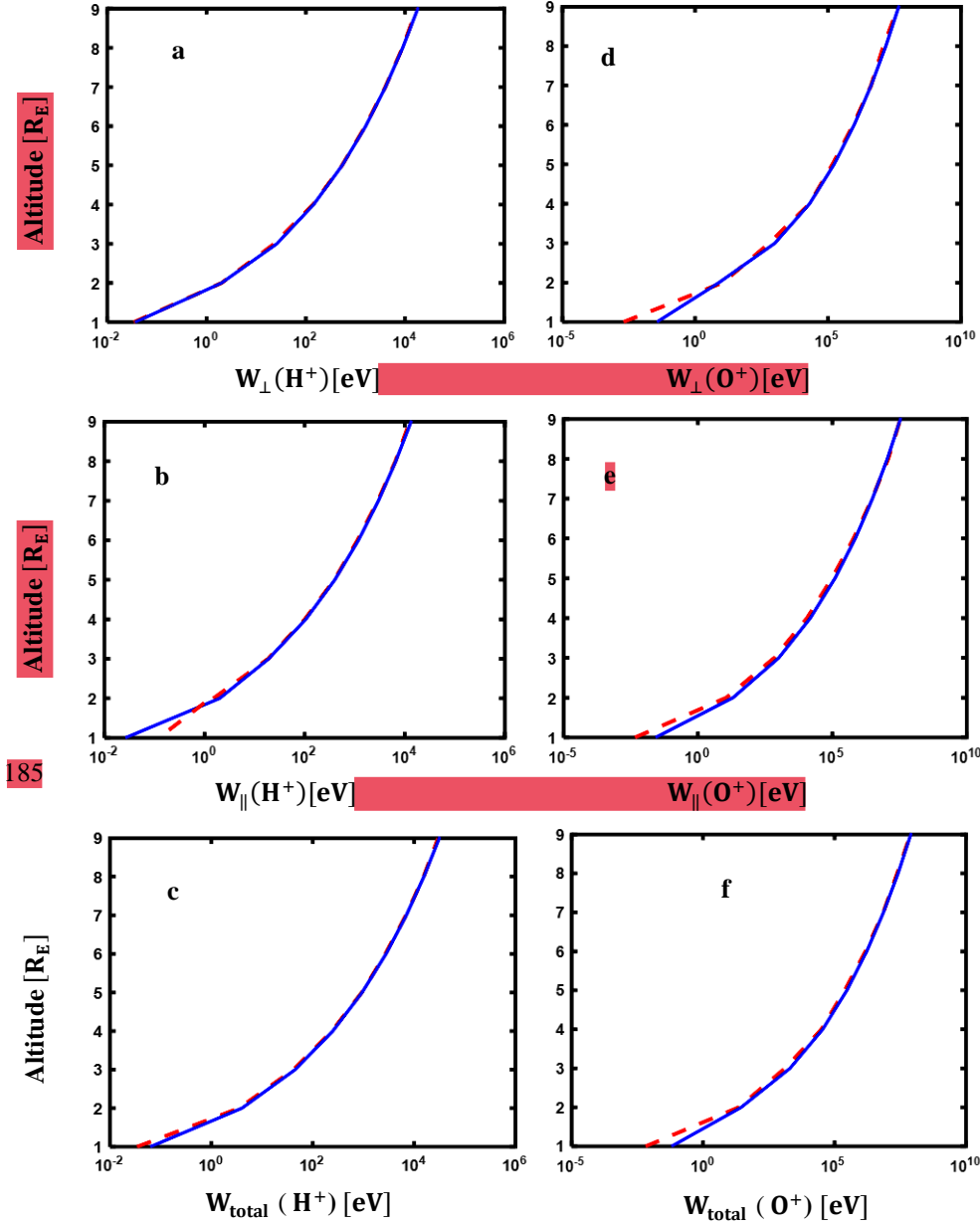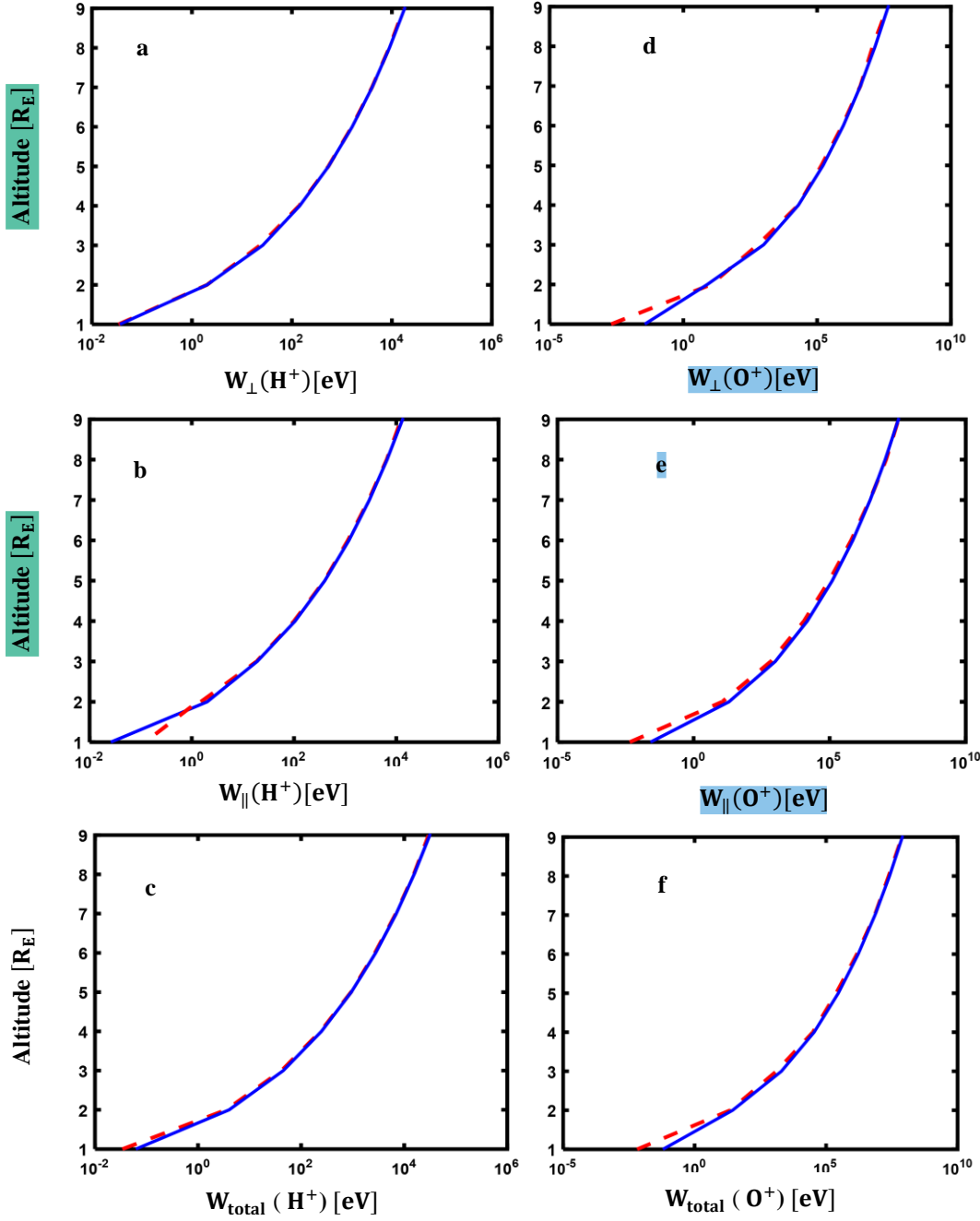
Figure 1: Altitude profiles of the estimates of the mean particle theory (blue solid lines) for auroral conditions with the Monte Carlo calculations (red dashed lines). Left panels, a, b, and c are for H$^+$ ions. Right panels, d, e, and f are for O$^+$ ions. The mean perpendicular energy W$_\perp$ considered here (panels a and d), mean parallel energy W$_\parallel$ (panels b and e) and mean total energy W$_{total}$ (panels c and f).

Figure 1: Altitude profiles comparing the mean particle theory estimates (blue solid lines) under auroral conditions with Monte Carlo calculations (red dashed lines). The left panels (a, b, and c) represent H^+ ions, while the right panels (d, e, and f) illustrate O^+ ions. Panels a and d show the mean perpendicular energy $W_\perp$, panels b and e present the mean parallel energy $W_\parallel$, and panels c and f depict the mean total energy W_{total}.
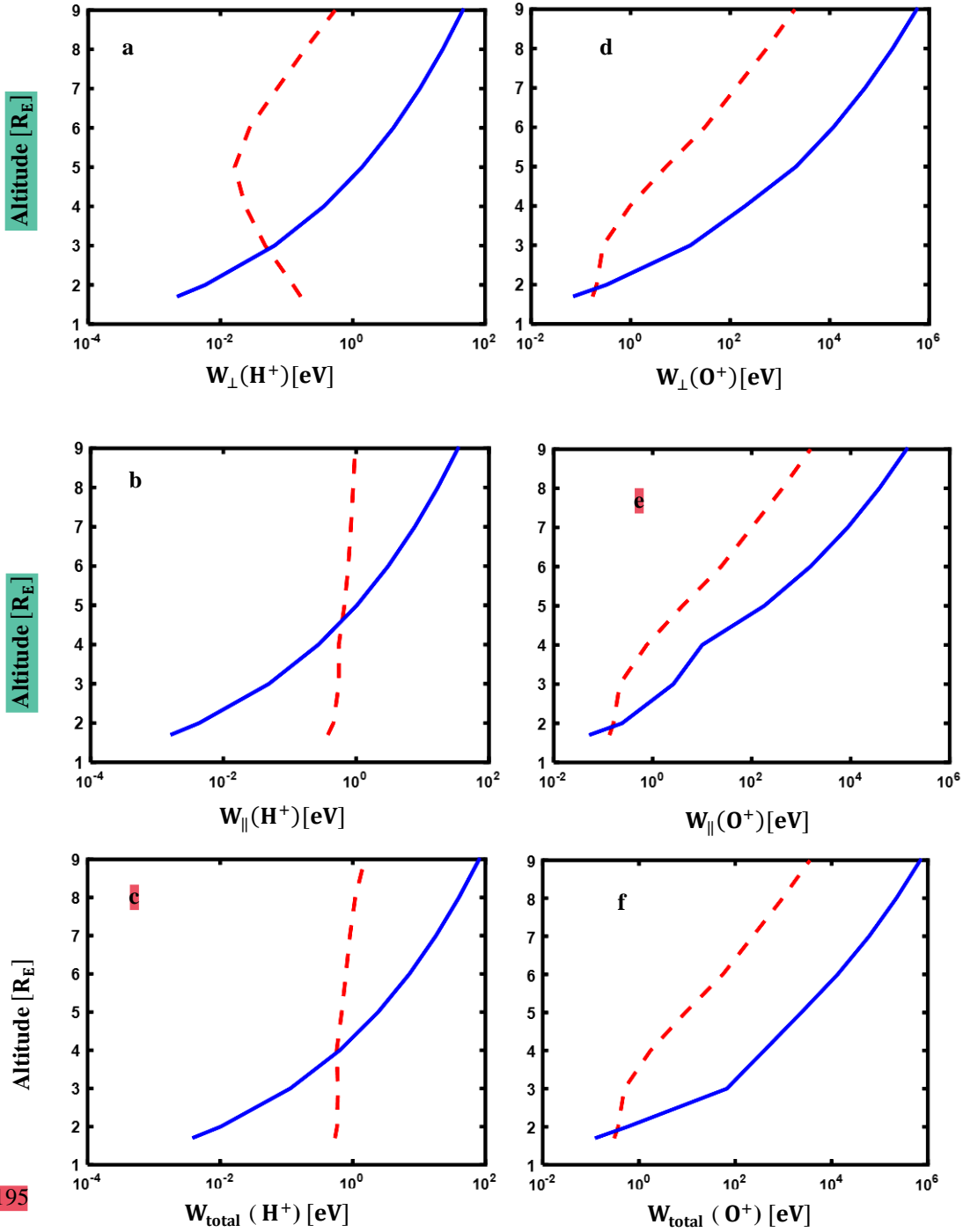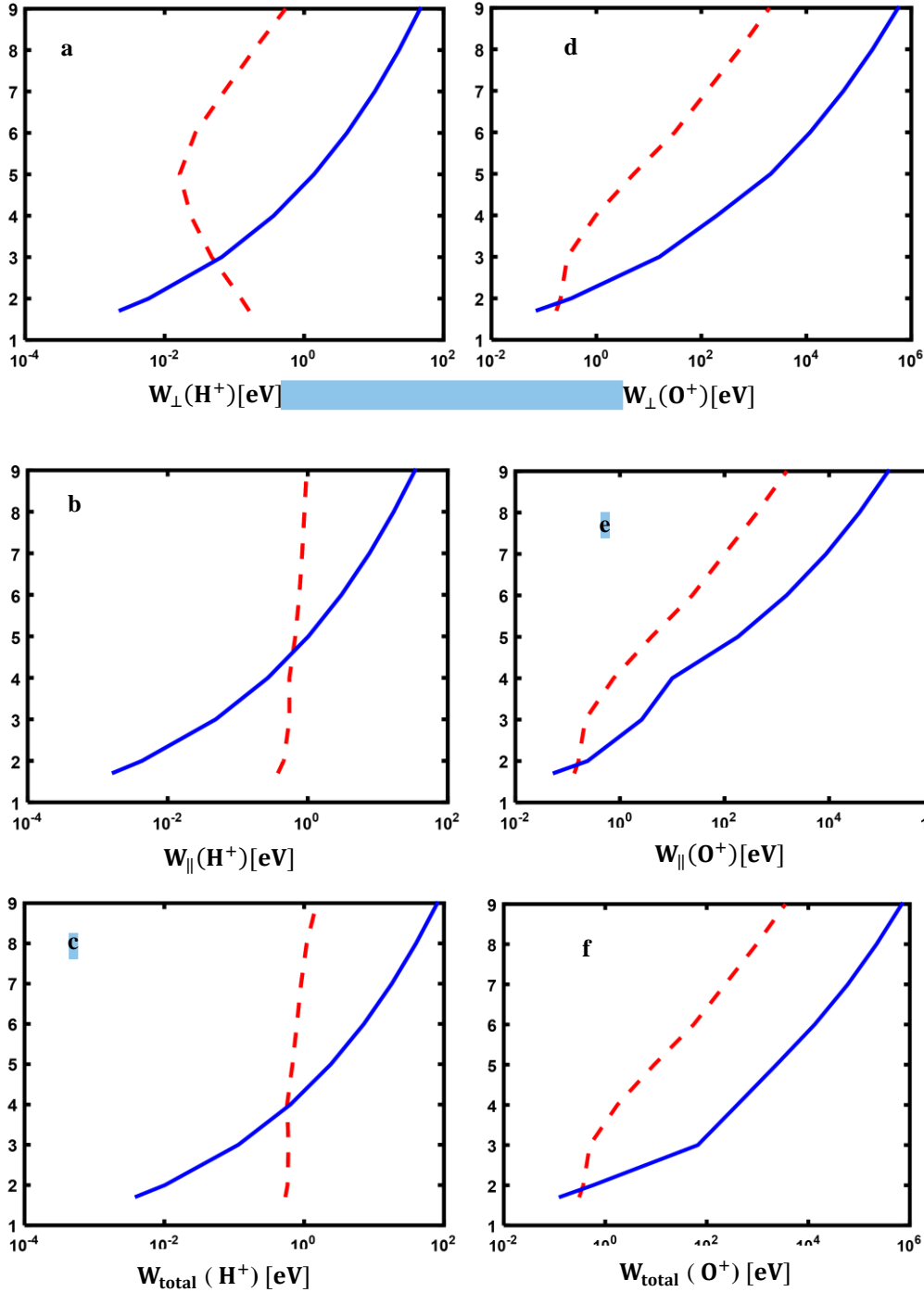
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Figure 2: Comparison of the estimates of the mean particle theory (blue solid lines) for polar wind conditions with the Monte Carlo calculations (red dashed lines). (right panels, d, e and f) for O^+ ions and (left panels, a, b and c) for H^+ ions and. The mean perpendicular energy $W_\perp$ represented by (panels a and d), mean parallel energy $W_\parallel$ (panels b and e) and total energy (panels c and f) with electromagnetic turbulence wavelengths ($\lambda_\perp \rightarrow \infty$), and altitude dependent diffusion coefficients.

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Figure 2: Comparison of the estimates of the mean particle theory (blue solid lines) for polar wind conditions with the Monte Carlo calculations (red dashed lines). (right panels, d, e and f) for O^+ ions and (left panels, a, b and c) for H^+ ions and. The mean perpendicular energy $W_\perp$ represented by (panels a and d), mean parallel energy $W_\parallel$ (panels b and e) and total energy (panels c and f).

200 and for $\lambda_\perp \to \infty$, an excellent agreement between Monte Carlo simulations and estimates of mean particle theory for H$^+$ and O$^+$ ions. It is worthwhile to mention that the heating, the energization, of the ions in the mean particle theory is due to the effect of wave particle interaction, only, that mainly depends on the value of the velocity diffusion coefficient. However, the energization process in the Monte Carlo model is a result of competition between wave particle interaction and external forces (gravity, polarization electric field and divergence geomagnetic field). This close agreement is due to the dominant effect of
205 wave particle interactions that overcome the effects of external forces because of the large values of $D_\perp(r)$, in other words, wave particle interactions dominate the energization process for both ions in this auroral region.

For polar wind region and for $\lambda_\perp \to \infty$, i.e. altitude dependent velocity diffusion coefficient $D_\perp(r)$. The values of the velocity diffusion coefficient are less than those in the auroral region, see eqs. (4) and (5). According to Fig. 2 there is a poor agreement between Monte Carlo simulations and mean particle theory estimates, this is due to the contribution of the
210 external forces in Barghouthi model that competes with the effect of wave particle interaction, however in the mean particle theory the external forces are not considered and the heating is due to wave particle interaction, only. We report here that mean particle theory is not suitable to be used in this region, and Monte Carlo simulations are more appropriate to be used as shown in Barghouthi et al. [2011] when they compared their Monte Carlo results with observations. Also, we have found that, when we use the diffusion coefficient that depends on altitude, its value becomes very large as altitude increases, and therefore the
215 values of the energies obtained from eqs. (16), (17) and (18) turn to be very high, but when the diffusion coefficient becomes velocity and altitude dependent according to eq. (8) (i.e. $k_\perp v_\perp/\Omega \geq 1$), the produced energies turn to be reasonable as shown in Fig.3 (blue solid lines and red dashed lines).

3.2 Comparison between Monte Carlo simulations, estimates of the mean particle theory, and available observations

In this sub-section, we present the Monte Carlo simulation results and the estimates of the mean particle theory that
220 have corresponding observations. Barghouthi [2008] compared the Monte Carlo simulation results that obtained from Barghouthi model with the corresponding observations for H$^+$ and O$^+$ ions outflows in the auroral region at different altitudes in the simulation tube, he obtained an excellent agreement, particularly, when the typical perpendicular wavelength of the electromagnetic turbulence was 8 km. Also, he observed that there is a broad agreement between the simulation results of the polar wind for this wavelength and the corresponding observations. For these reasons, we chose to have the results of
225 the comparison with corresponding observations when $\lambda_\perp = 8\ km$, i.e. when the velocity diffusion coefficient is altitude and velocity dependent. We will compare the outcomes of our Monte Carlo simulations, the estimates of mean particle theory, and available observations obtained from different published articles. Observations of O$^+$ ions at various altitudes were obtained for parallel velocity, perpendicular temperature and parallel temperature for both polar wind and auroral regions from [Nilsson et al., 2013] and observations for parallel velocity and perpendicular temperature for O$^+$ ions in central polar cap and cusp
230 regions from Barghouthi et al. [2016].

Figure 3 presents the results of comparisons in auroral (left panels) and polar wind (right panels) regions. The comparison results between Monte Carlo method (red dashed lines), and estimates of the mean particle theory (blue solid lines) at $\lambda_\perp = 8\ km$ with observations (minimum (doted lines), avarege (doted dashed lines) and maximam (black dashed lines)) for O$^+$ ions. It is very clear, in four panels, that the Monte Carlo simulation results and the estimates of the mean particle
235 theory are very close to each other and they have excellent agreement at low altitudes, however at high altitude they have similar qualitative behaviour and good agreement, we claim that this acceptable and reasonable agreement is due the implementation of the appropriate altitude and velocity dependent diffusion coefficient in both methods. According to

Monte Carlo simulations and estimates of mean particle theory for H^+ and O^+ ions. It is important to highlight that in the mean particle theory, ion heating and energization occur solely due to wave-particle interactions, which primarily depend on the value of the velocity diffusion coefficient. In contrast, the energization process in the Monte Carlo model results from the competition between wave-particle interactions and external forces, such as gravity, the polarization electric field, and the divergence of the geomagnetic field. The close agreement between the two methods can be attributed to the dominant influence of wave-particle interactions, which prevail over the effects of external forces due to the high values of $D_\perp(r)$. In other words, wave-particle interactions are the primary drivers of the energization process for both H^+ and O^+ ions in this auroral region. For polar wind region and for $\lambda_\perp \to \infty$, this means the velocity diffusion coefficient, $D_\perp(r)$, is altitude dependent. The values of the velocity diffusion coefficient are less than those in the auroral region, see eqs. (4) and (5). According to Fig. 2 there is no agreement between Monte Carlo simulations and mean particle theory estimates, this is due to the contribution of the external forces in Barghouthi model that competes with the effect of wave particle interaction, however in mean particle theory the external forces are not considered and the heating is due to wave particle interaction. We report here that mean particle theory is not suitable to be used in this region, and Monte Carlo simulations are more appropriate to be used as shown in Barghouthi et al. [2011] when they compared their Monte Carlo results with observations. Also, we have found that, when we use the diffusion coefficient that depends on altitude only, its value becomes very large as altitude increases. Therefore the values of the particle energies obtained from eqs. (16), (17) and (18) turn to be very high, but when the diffusion coefficient becomes velocity and altitude dependent according to eq. (8) (i.e. $k_\perp v_\perp / \Omega \geq 1$), the produced particle energies turn to be reasonable as shown in Fig.3 (blue solid lines and red dashed lines).

3.2 Comparison between Monte Carlo simulations, estimates of the mean particle theory, and available observations

In this sub-section, we present, only, the simulation results and the estimates of the mean particle theory that have corresponding observations. Barghouthi [2008] compared the Monte Carlo simulation results that obtained from Barghouthi model with the corresponding observations for H^+ and O^+ ions outflows in the auroral region at different altitudes in the simulation tube, he obtained an excellent agreement, particularly, when the typical perpendicular wavelength of the electromagnetic turbulence was 8 km. Also, he observed that there is a broad agreement between the simulation results of the polar wind for this wavelength and the corresponding observations. For these reasons, we chose to have the results of the comparison with corresponding observations when $\lambda_\perp = 8 \, km$, i.e. when the velocity diffusion coefficient is altitude and velocity dependent. We will compare the outcomes of our Monte Carlo simulations, the estimates of mean particle theory, and available observations obtained from different published articles. Observations of O^+ ions at various altitudes were obtained for parallel velocity, perpendicular temperature and parallel temperature for both polar wind and auroral regions from [Nilsson et al., 2013] and observations for parallel velocity and perpendicular temperature for O^+ ions in central polar cap and cusp regions from Barghouthi et al. [2016].
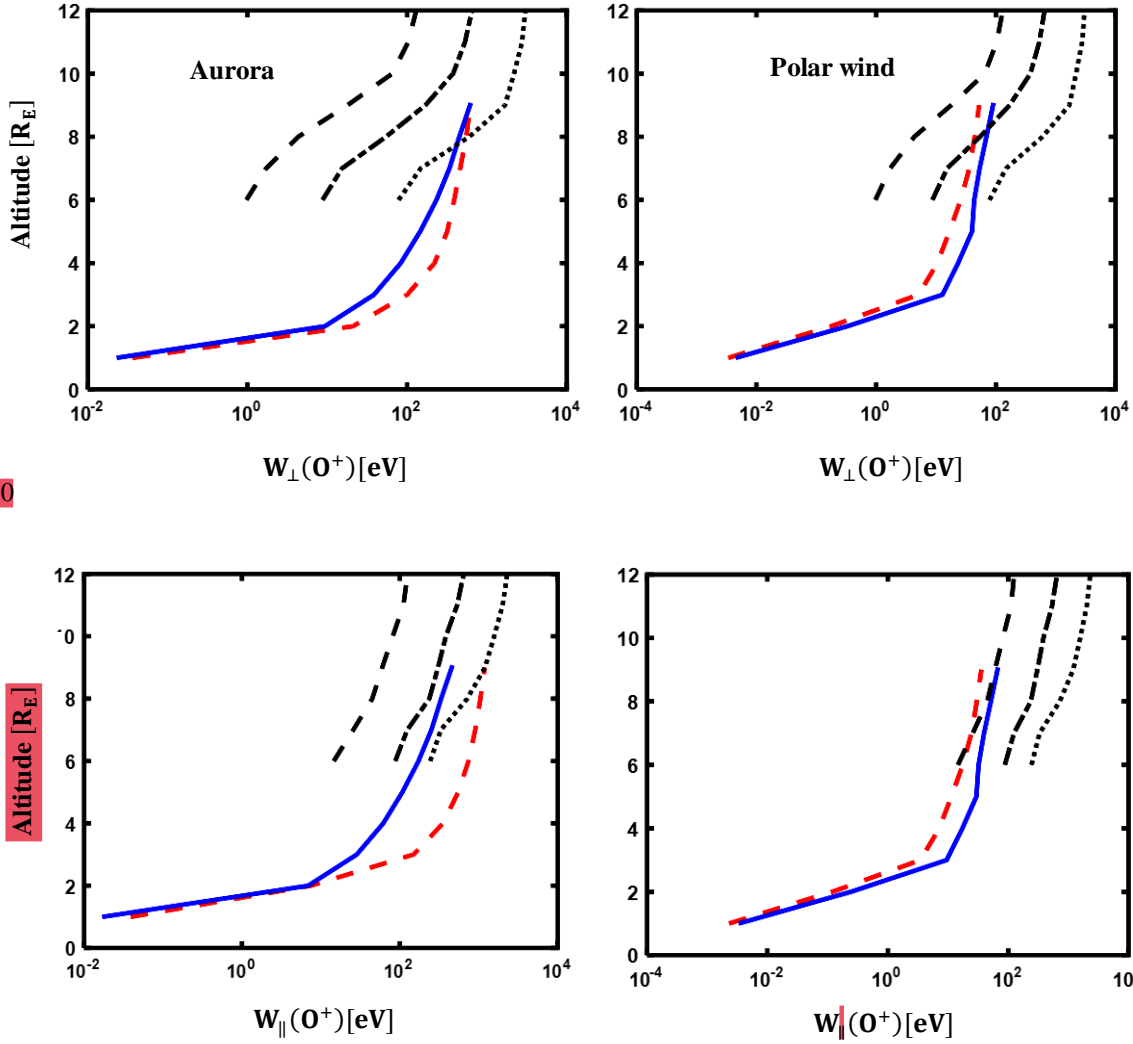
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Figure 3: Comparison of the estimates of the mean particle theory (blue solid lines) and Monte Carlo calculations (red dashed lines) for auroral conditions (left panels) and polar wind conditions (right panels), in addition to observations, minimum (dott lines), avarege (dotted dashed lines) and maximam (black dashed lines) for O^+ ions. The mean perpendicular energy $W_\perp(O^+)$ considered here (top) and mean parallel energy $W_\parallel(O^+)$ (bottom), we have considered the wavelength of the electromagnetic turbulence to be $\lambda_\perp = 8km$.

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Figure 3 displays the comparison results for the auroral (left panels) and polar wind (right panels) regions. It contrasts the Monte Carlo method (red dashed lines) and the estimates from mean particle theory (blue solid lines) at $\lambda_\perp = 8$ km, alongside observational data represented by maximum (dotted lines), average (dashed-dotted lines), and minimum (black dashed lines) values for O^+ ions. Across all four panels, it is evident that the results from the Monte Carlo simulations and the mean particle theory estimates are closely aligned, demonstrating excellent agreement at lower altitudes. At higher altitudes, while both methods display similar qualitative behaviour and maintain good agreement, we attribute this acceptable and reasonable correlation to the application of a suitable altitude- and velocity-dependent diffusion coefficient in both approaches.
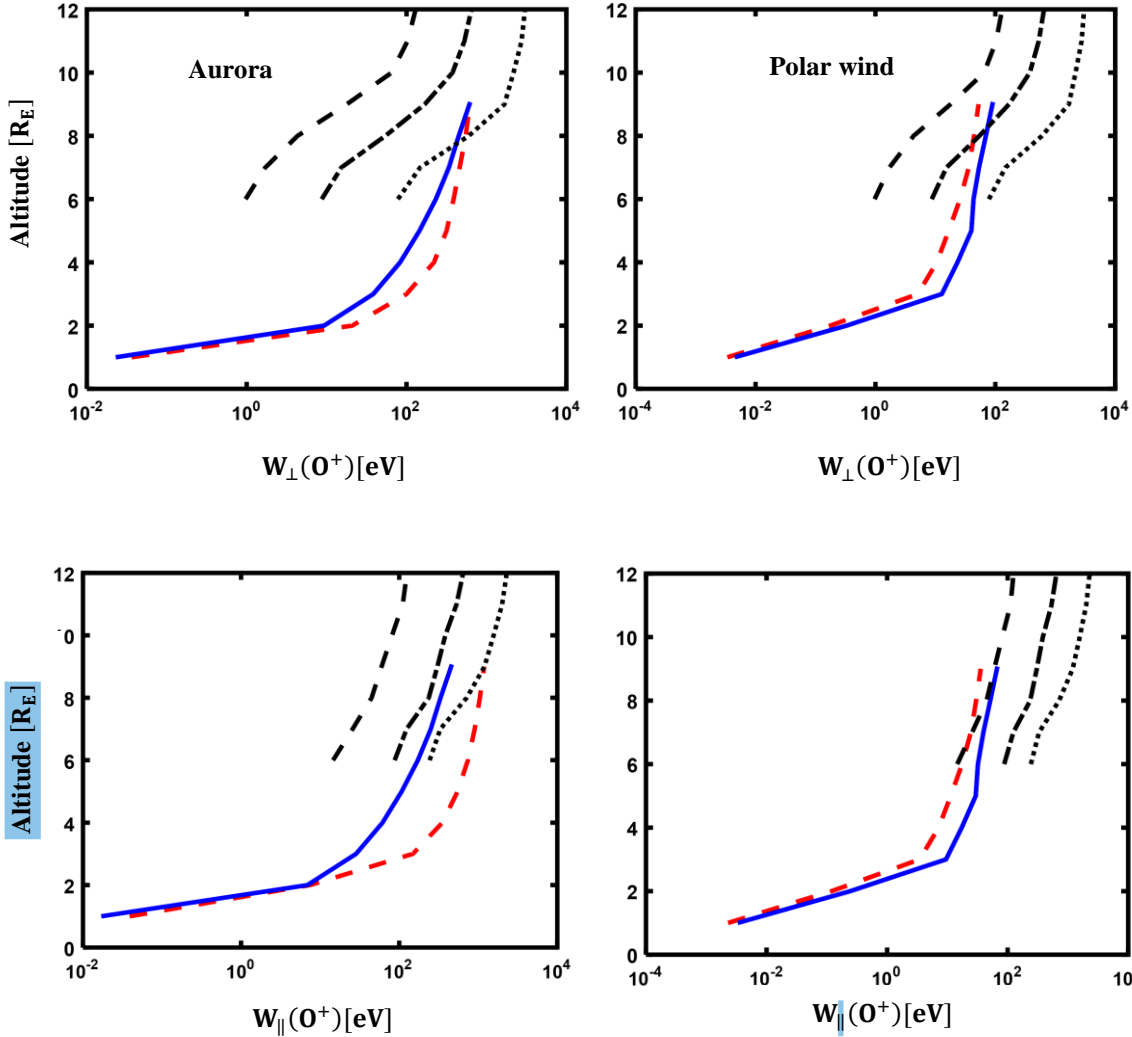

Figure 3: Comparison of mean particle theory estimates (blue solid lines) and Monte Carlo calculations (red dashed lines) for auroral conditions (left panels) and polar wind conditions (right panels), along with observational data represented by maximum (dotted
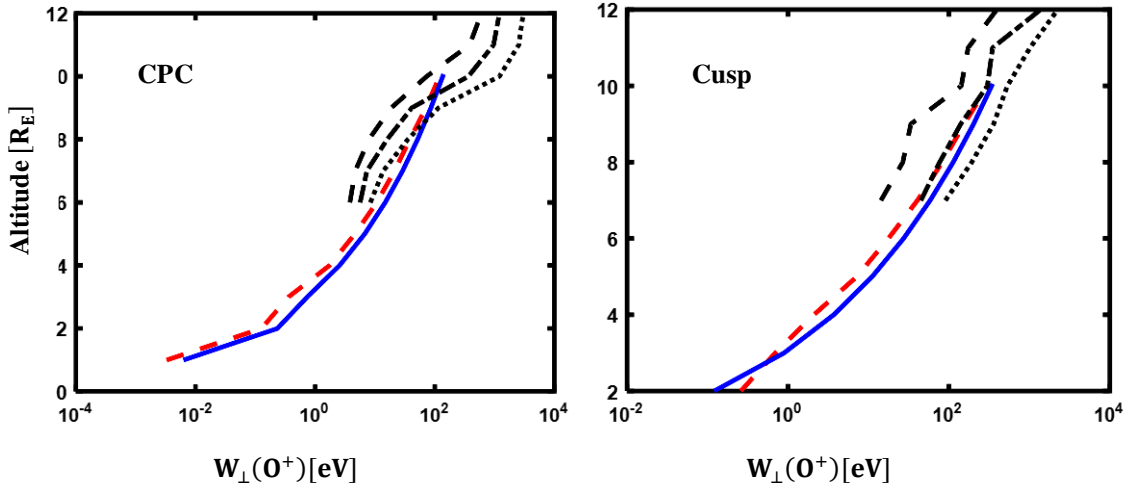
Figure 4: Comparison of the estimates of the mean particle theory (blue solid lines) with the Monte Carlo calculations (red dashed line), for central polar cap conditions (left panel) and cusp conditions (right panel) in addition to the observations (minimum (doted lines) , avarege (doted dashed lines), and maximam (black dashed lines)) for O^+ ions. Mean perpendicular energy $W_\perp(O^+)$ is considered with electromagnetic turbulence wavelength ($\lambda_\perp \to \infty$).

comparison with observations, it is obvious that the simulation results and the estimates of the mean particle theory are in the range of observations. This is evidence that $D_\perp(r, v_\perp)$ is suitable to be used in Monte Carlo simulation and in mean particle theory. To be specific, the results of perpendicular energy $W_\perp(O^+)$ from both studies, Monte Carlo and mean particle theory are in good agreement with minimum values of the observations in the auroral region, however, they are in good agreement with average values of the observations. For mean parallel energy $W_\|(O^+)$ (Fig. 3, bottom panels), simulation results and mean particle theory estimates are close to the average and minimum values of the observations in the auroral region and they are in excellent agreement with maximum values of the observations in the polar wind region.

In central polar cap (Fig. 4, left panel) and cusp (Fig. 4, right panel) regions, we have used Nilsson et al. [2013] diffusion coefficient that is altitude dependent only, i.e. when $\lambda_\perp \to \infty$, eqs. (6) and (7), we have obtained excellent agreement, at all altitudes, between Monte Carlo simulations and the estimates of the mean particle theory. Also, we have a very good agreement with observations in both regions, all results are very close with minimum and average observations in central polar cap region and they are very close to the average values of the observations in the cusp region.

As a result of these comparisons, it is important to choose the appropriate form of the velocity diffusion coefficient in every region of study in earth magnetosphere when considering the energization, heating, process of ions.

4 Conclusions

We have compared the energy components, mean perpendicular energy $W_\perp$, mean parallel energy $W_\|$, and total mean energy W_{total} for H^+ and O^+ ions by using Monte Carlo method and the estimates of the mean particle theory in different

lines), average (dashed-dotted lines), and minimum (black dashed lines) values for O^+ ions. The top panel shows the mean perpendicular energy, while the bottom panel presents the mean parallel energy. In this analysis, we have considered the wavelength of electromagnetic turbulence to be $\lambda_\perp = 8$ km.
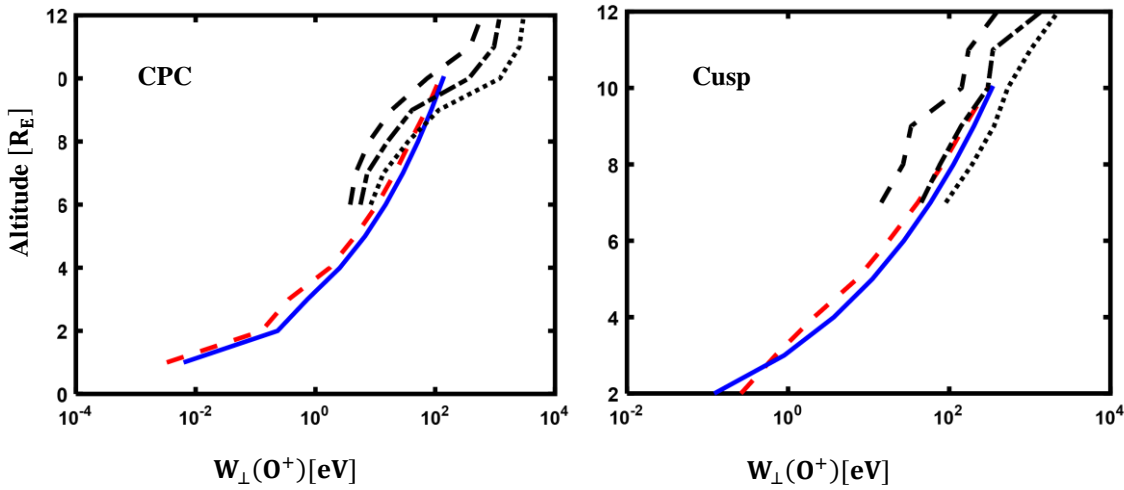
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Figure 4: Comparison of the estimates of the mean particle theory (blue solid lines) with the Monte Carlo calculations (red dashed line), for central polar cap conditions (left panel) and cusp conditions (right panel) in addition to the observations (maximum (doted lines) , avarege (doted dashed lines), and minimum (black dashed lines)) for O^+ ions. In this analysis, we have considered the wavelength of electromagnetic turbulence $\lambda_\perp \to \infty$.

Based on the comparison with observations, it is evident that the simulation results and estimates from the mean particle theory fall within the observational range. This indicates that $D_\perp (r, v_\perp)$ is appropriate for use in both the Monte Carlo simulation and the mean particle theory. Specifically, the results for perpendicular energy $W_\perp(O^+)$ from both the Monte Carlo method and mean particle theory align well with the maximum observed values in the auroral region, and they also exhibit good agreement with the average observed values in the polar wind region. Regarding the mean parallel energy $W_\parallel(O^+)$ (as shown in Fig. 3, bottom panels), the simulation results and estimates from the mean particle theory are closely aligned with both the average and maximum observed values in the auroral region, and they demonstrate excellent agreement with the minimum observed values in the polar wind region.

In the central polar cap (Fig. 4, left panel) and cusp (Fig. 4, right panel) regions, we utilized the altitude-dependent diffusion coefficient from Nilsson et al. [2013], specifically when $\lambda_\perp \to \infty$. This approach resulted in excellent agreement at all altitudes between the Monte Carlo simulations and the estimates of the mean particle theory. Furthermore, we observed a strong correspondence with observational data in both regions; the results closely match the maximum and average values in the central polar cap region and align well with the average values from observations in the cusp region.

As a result of these comparisons, it is crucial to select the appropriate form of the velocity diffusion coefficient for each region of interest within the Earth's magnetosphere when analysing ion energization and heating processes.

regions of earth magnetosphere (polar wind, auroral, central polar cap, and cusp regions). For altitude dependent diffusion coefficient, we have found excellent agreement between the results of both methods in the auroral regions and poor agreement in polar wind region, also we have found that the energy turned to be very high, not reasonable, at middle and high altitudes. Because of that we introduce the velocity and altitude dependent diffusion coefficient that produce reasonable energies at low, middle, and high altitudes in both regions. When we have compared the Monte Carlo simulations and estimates of the mean particle theory with corresponding observations, we have found good agreement between simulations, estimates and observations in all regions, i.e. polar wind, aurora, central polar cap, and cusp. Simulation results and the estimates of mean particle theory are in the same range of observations. Here, we report that the diffusion coefficients that have been used in Monte Carlo simulation and in the mean particle theory and produced acceptable agreement when compared to the corresponding observations are appropriate to be used in these regions.

For future work, we need to search for more observations in different earth magnetosphere regions in order to have more comparisons and to confirm which diffusion coefficient is more appropriate and which method gives more accurate results when compared to corresponding observations.

5 Code and data availability

The source code, data, and input files necessary to reproduce the results are available from the authors upon request (barghouthi@staff.alquds.edu).

6 Author contribution

First author (Imad Barghouthi) suggested the problem, provided the model (Barghouthi model), discussed the simulation results, and wrote the manuscript.

Second author (May Halaika) ran the model (Barghouthi model), obtained the simulation results, and plotted the figures.

7 Competing interests

The authors (Barghouthi and Halaika) declare that they do not have any competing interests.

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4 Conclusions

We conducted a comparative analysis of energy components— mean perpendicular energy $W_\perp$, mean parallel energy $W_\parallel$, and total mean energy W_{total}, for H^+ and O^+ ions using both the Monte Carlo method and estimates from mean particle theory across various regions of Earth's magnetosphere, including the polar wind, auroral, central polar cap, and cusp regions. Utilizing altitude-dependent diffusion coefficients, we found excellent agreement between the two methods in the auroral regions; however, there was a lack of concordance in the polar wind region. Moreover, we observed that energy values at middle and high altitudes were unrealistically elevated compared to observational data. To mitigate this issue, we implemented velocity- and altitude-dependent diffusion coefficients, which yielded reasonable energy values at low, middle, and high altitudes across both regions. When we compared the results of the Monte Carlo simulations and the mean particle theory estimates with available observational data, we noted good agreement throughout all studied regions—polar wind, auroral, central polar cap, and cusp. Both the simulation results and the mean particle theory estimates aligned well within the observed data range. We conclude that the velocity diffusion coefficients used in both methods, which produced acceptable consistency with the observations, are appropriate for application in these regions.

For future research, it is essential to conduct further observations in various regions of Earth's magnetosphere. This will facilitate additional comparisons and help identify the most suitable diffusion coefficient, as well as ascertain which method yields the most accurate results in relation to the corresponding observational data.

5 Code and data availability

The source code, data, and input files necessary to reproduce the results are available from the authors upon request (barghouthi@staff.alquds.edu).

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First author (Imad Barghouthi) suggested the problem, provided the model (Barghouthi model), discussed the simulation results, and wrote the manuscript. Second author (May Halaika) ran the model (Barghouthi model), obtained the simulation results, and plotted the figures.

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