

Espoo, 13.11.2020

Dear David,

Thank you for your thorough review of our paper. We have addressed all of your very good points which we think will significantly improve the quality of the paper. Below, we go through the points in detail.

“My first scientific request is for the authors to add one parameter to their study. They already consider densities, speeds, temperatures, and magnetic fields within the jets, but I request they add another parameter: deflection of flow away from the sun-earth line. I think it would be interesting to know if the jets maintain flow along this line as they approach the magnetopause or are deflected just like the background magnetosheath plasma flow. “

Thank you for this suggestion. We now present Figure 11, where there are two parameters, the relative flow velocity with respect to the background magnetosheath, and the relative flow deflection angle with respect to the background. This new figure shows that the jets on average flow faster than the background magnetosheath and more along the Sun-Earth line with respect of the background flow. However, in the Run with low Mach number with 30 degrees cone angle, this behaviour is different, the jets in this run flow slower and more towards the flanks with respect of the background. We discuss this new finding within the discussion. We think the jets in the LM30 run may be of different character than the ones in the other runs, however, it is clear that in the other runs the jets maintain their flow direction and are not just carried by the magnetosheath flow.

“My second scientific request is that the authors tell whether or not there is a difference in event occurrence patterns in their simulation results regarding location behind the quasi-parallel and quasi-perpendicular bow shock. I am expecting there would be more events in the dawn than the dusk magnetosheath for the spiral IMF orientations they simulate. Can the authors please add some words about that?”

Thank you for this suggestion. We added a Figure 12, showing the jet end position in Y coordinate, and the jet end position angle in the XY-plane. In this figure, we grouped the low and high Alfvén Mach numbers together, and look at the difference between the cone angles in the runs. Indeed, you are right, the jets do tend to end up at the dawn flank, especially within the runs with 30° cone angle. This same tendency is seen in the 5° cone angle runs, but it is slightly lower in magnitude, consistent with the more radial IMF orientation. We have also added discussion of this matter in the Discussion.

“My third scientific request is that the authors tell what the vertical pressure stripes are in the magnetosheath. They are so prominent in Figure 1. What is causing them?”

The large-scale flow pattern in the magnetosheath in Vlasiator is as expected: streamlines diverge from the Sun-Earth line. However, the distribution of the velocity magnitude and the density is more complex, due to kinetic processes arising at the quasi-parallel shock, which results in the large-scale structures we observe in the magnetosheath. In essence, these stripes are remnants of the ULF wave fluctuations at the bow shock. This information was added in connection to Figure 1.

Line 6, page 1. Using a statistics → using statistics Line 21, Page 5. At force → in force The stars in Figure 1 are too small. Lines 5-7 on Page 10. Can the authors state a typical dimension at this point? Line 22 on Page 10. No need for ‘RE’ Lines 25-26. Page 10. Can the authors state that the jets appear to be disintegrating or dying by diffusion here?

Thank you for these comments, we corrected them in the revision. We think based on Figure 10 that the jet “dies” through diffusion, because towards the end of its lifetime there are no steep gradients

within the jet parameters (without counting the final seconds). We added this information in connection to Figure 10.

On behalf of all the co-authors, Minna Palmroth

Espoo, 13.11.2020

Dear Referee #2,

Thank you for your thorough review of our paper. We have addressed all of your very good points which we think will significantly improve the quality of the paper. Below, we go through the points in detail.

There are two major limitations of these simulations: the kinetic treatment is truncated (the kinetic physics of electrons is disregarded) and the plasma macroscopic dynamics is reduced to two spatial dimensions, in the equatorial plane. Can the authors discuss these limitations and demonstrate that the numerical simulations based on such a truncated description of plasma dynamics still captures accurately the physics of jets as observed in the magnetosheath (see, e.g., the analysis of the jet anatomy by Karlsson et al, 2018)?

Thank you very much for this clarification.

The electron kinetic physics can be neglected for magnetosheath jets. The short answer is that the kinetic pressure of the electrons downstream of the Earth's bow shock is smaller by about a factor of 10 with respect to the ion pressure. The reason for this is that the Mach number is not large enough for the dissipation at the Earth's bow shock to necessitate heating the electrons, and thus the ions are the primary beneficiary of the heating. Since energy has to be transferred from somewhere (ion or electron thermal pressure, or magnetic pressure) to enhance the kinetic pressure (or the ram pressure) and lead to jets' formation, the primary source for such energy transfer has to be comparable in size to jets kinetic pressure. Additionally, the fact that jets have scales/dimensions between fluid and ion kinetic scales is also an indication that an ion kinetic model should be sufficient. While electrons might have some nonzero contribution in the formation of jets, the dominant effect comes from the ions. Now of course, the electron physics inside jets is important to maintain quasi-neutrality and one will certainly find electron kinetic fluctuations that are relevant to the local thermodynamical properties of the plasma, but that is a separate problem.

Regarding the 2D spatial dimensionality: The largest caveat of this is the position of the magnetopause, as we already state in the manuscript. The bow shock – magnetosheath interactions are rather accurately reproduced by 2D kinetic simulations as can be seen from a large number of previous papers (e.g., Karimabadi et al., 2014 PoP <http://dx.doi.org/10.1063/1.4882875>; Blanco-Cano et al al, 2006 JGR doi:10.1029/2005JA011421). The first Referee suggested a new plot looking at the final position of the jet within the dawn and dusk in the 30-degree cone angle runs. Of course the final position of the jet within the magnetosheath will lack the third dimension in a 2D simulation. However, it is important to investigate the dawn/dusk asymmetry of the final position, as the first Referee suggests.

We have added a summary of the caveats in the Discussion as requested by the Reviewer.

The usefulness of the comparison between a statistical analysis of simulations and MMS observations of jets (collected over almost four years) is not clear to me. The authors state that MMS data serve to “verify the simulation behavior”, thus, they aim to validate the hybrid Vlasov simulations with MMS data. However, there is no indication on the upstream solar wind conditions for the 6142 jets detected by MMS1. On the other hand, the statistical ensemble of 924 jets observed in hybrid Vlasov simulations is generated by only four sets of solar wind parameters (magnetic field, density, bulk velocity, cone angle and Mach number), as shown in Table 1. Therefore I am wondering how could one relate, compare and validate simulations of the magnetosheath jets as

resulting from these four sets of solar wind parameters with the observations of magnetosheath jets corresponding to 6142 MMS events, and therefore for a much larger set of solar wind conditions? The authors do show that MMS observations for one jet (figure 2) correspond to a solar wind velocity “around 410 kms⁻¹, density about 3.8 particles per cubic centimetre, the magnetic field vector about [-4, 3, -2] nT, and MA is 7” relatively close to the initial conditions assumed for the simulations reported in case LM30. However, even for this case there are significant differences in terms of jet characteristics, as outlined by the authors: “The velocity and the magnetic field show some discrepancies, such as slower flows and more variable magnetic fields at MMS, reflecting the differences in the solar wind conditions as well as the different relative positions within the magnetosheath. [. . .] Similarly, the differences in the density and temperature can be understood in terms of the differences in the solar wind parameters”. Although in Figure 5 is suggested a normalization with respect to solar wind parameters, it is not clear where the solar wind parameters are derived from in case of MMS data.

Thank you for this important clarification. The statistical comparison can never be made in an exact manner for the following reason: With an infinite amount of computational resources we could run the 6142 events, detect jets from all these runs, and then make the comparison. However, even in this case, there would be discrepancies, because in the observational data set one gets one event per solar wind condition, while in the simulation one set of solar wind conditions will produce a large amount of jets. The comparison would be accurate only if the observations would gather jets simultaneously from all around the magnetosheath, which is as impossible with current resources as it is to run 6142 kinetic simulation runs.

We would like to emphasise that the main comparison is being done through variables that have subtracted the background magnetosheath measurements (ΔX). Due to the differences in solar wind conditions, we are also using normalization to the solar wind values as units (Figures 5,6). This indirectly solves the discussed issue. The plasma behind the bow shock regardless of the phenomenon we study is of solar wind origin. While variation in the observations will certainly depend on the solar wind conditions, by subtracting the background conditions and normalizing to the solar wind we are minimizing this effect.

We would also like to mention that the larger number of observational jets is very important to statistically verify the Vlasiator results. This is because then we avoid cherry-picking the data observed by MMS, while we are not in control of its orbit or position.

The solar wind parameters used to normalise MMS data are from the OMNI database.

We added this information in the paper.

The supplementary video material raises additional questions. The magnetosheath can be roughly identified in this movie as the spatial region between (an imaginary) magnetopause (as in Figure 3 of the manuscript – the Shue model) and the red line (the bow-shock). (The authors point out the magnetopause cannot be easily identified in their simulations). It looks like the dynamical pressure irregularities do not follow the general streaming or the magnetosheath flow as they seem to be predominantly convected in the duskward direction, which is puzzling. From the general gasdynamics of a supersonic flow around an obstacle one would expect a stagnation point in the nose region and azimuthal flows, duskward and dawnward, at the two flanks.

This is an excellent question. The large-scale flow pattern in the magnetosheath in Vlasiator is as expected: streamlines diverge from the Sun-Earth line. The structures follow the field lines as they pile up and drape around the magnetopause. However, the distribution of the velocity magnitude and the density is more complex, due to kinetic processes arising at the quasi-parallel shock, which results in the large-scale structures we observe in the magnetosheath. In essence, these stripes are

remnants of the ULF wave fluctuations at the bow shock, and they can only become evident in a simulation that reproduces bow shock – magnetosheath interactions, i.e., they are not visible in a fluid description. We added this information to the revision.

Also, one notes that some jets (identified by red dots) apparently occur spontaneously at significant distances from the bow-shock, in the magnetosheath, where they spend some time and then disappear (“die”) (e.g. the sequence between $t=551$ and $t=557$ seconds, roughly at $X=12$, $Y=1$). I am not sure this behavior really captures the dynamics of magnetosheath jets which are believed to be structures born at the bow-shock (e.g., Karlsson et al., 2018). When watching the movie in the supplementary material I am wondering if some of the structures considered in this study are not in fact local magnetosheath fluctuations (some of them related perhaps to the inherent numerical simulation noise and/or truncated physics) and not jets, as they are defined in the literature.

We did in fact discuss before submitting, whether we should restrict to events which are born at the bow shock, or just take all events which fulfil the criterion. We decided to take all events fulfilling the criterion because in the observations one cannot do this restriction: No observational jet statistics can restrict to cases originating in the bow shock because they have no idea whether the event they observe in fact did originate there, or whether the magnetosheath just altered the flow conditions such that the local parameters fulfil the jet criteria. We do suspect that this occurs in many statistical data sets, and this is also why we limit to smaller jets to avoid large regions near flanks (page 5, lines 21-23 of the submitted version). We actually have a follow-up paper to be submitted in 2020, where we restrict to cases originating at the bow shock. We added this information to the revision.

Further, we would like to add that the origin of jets is still controversial (see the review by Plaschke et al., 2018 <https://doi.org/10.1007/s11214-018-0516-3>). Some of the generation mechanisms do not involve bow shock processes, as e.g., some authors have proposed that magnetosheath reconnection could produce jets (Retino et al, 2007 <https://doi.org/10.1038/nphys574>).

The statistical analysis of simulations results is interesting. However, given the complexity of this type of numerical simulations, the authors should give more evidence on how the results depend on the numerical setup and noise. I assume such tests were extensively performed and the authors should discuss if and how the results depend on the spatial and velocity resolution, as well as the time step assumed for the simulations. This is important for understanding the results of the statistical analysis.

The spatial resolution is better in Vlasiator than in any of the existing ion-kinetic simulation studying the jets currently. The same goes for the velocity space resolution. Timestep is determined by the CFL conditions, which it fulfils, and therefore we do not think that the timestep is an issue. We would like to emphasise that Vlasiator is a supercomputer simulation, requiring about 3-5 million core-hours per run, and around 10 T of disk space to save the results. Unfortunately, we cannot afford large-scale numerical tests with Vlasiator. However, we have carefully analysed the minimum requirements for spatial and velocity space resolution with a simpler setup of Vlasiator solvers (Pfau-Kempf et al., 2018, doi:10.3389/fphy.2018.00044), and we have carefully chosen all parameters so that the physics of jets is properly described. Indeed, we can miss the smallest jets due to our resolution, but such small jets would also be missed from spacecraft observations. We added this information in the revision.

Further, methods based on the Vlasov approach modelling the distribution function do not include noise. That is their main advantage compared to the complementary particle-in-cell methods, where the distribution functions are constructed from the statistics of particles. We state this in the first paragraph of 2.1.

The jet size distribution, as revealed by Figures 7 and 8, indicates that the structures studied with the hybrid Vlasov simulations are rather small scale and short lived, with most probable sizes of the order of hundred of kilometers and most probable lifetimes of the order of seconds. While lifetimes were not yet investigated with in-situ data, scale sizes derived from observations (e.g. Plaschke et al, 2016) are orders of magnitude larger, as noted by the authors themselves. This difference leads me back to the remark I made in the previous paragraph: are the structures studied in this report magnetosheath jets or local non-homogeneities downstream the bow-shock, and satisfying the selection criteria based on the dynamic pressure excess (Eq. 1) ?

The answer is shortly: Yes. The previous literature of magnetosheath jets based on observational statistics has no way of confirming that the features they observe are in fact born at the bow shock. The only way to characterise the jets is to define a set of criteria and follow them; this is what we have done, and as our results are so closely in agreement with the observations, we can say that both jets born at the bow shock (which only Vlasov can confirm) and the local homogeneities fulfilling the criteria are in accordance with observations. Both types are called magnetosheath jets in observational literature, while only models can state something of their origin.

We would also like to point out that the size of the jets is a function of the criteria used. In our previous paper (Palmroth et al., 2018, reference in the manuscript), we studied the Vlasov jet with the Plaschke criterion along with the Archer and Horbury criterion we have in the present paper. The Plaschke criterion produces larger jets than the Archer and Horbury criterion, and hence our results on the size and dimension are not directly comparable to Plaschke et al 2016 statistics. Regarding Plaschke criterion, a new paper by them (<https://agupubs.onlinelibrary.wiley.com/doi/abs/10.1029/2020JA027962>) states in its keypoints that “Most magnetosheath jets are an order of magnitude smaller than previously reported”, and in Discussion they state “First, jets are, in general, much smaller than reported in most, if not all, literature on this subject so far, where scale sizes on the order of 1RE are typically stated. Median scale sizes of $D_{perp} = 0.12RE$ and $D_{par} = 0.15RE$ ”. These results agree with our results. Finally, we stress here that there is also a bias in observations towards larger structures: smaller jets won't be detected from spacecraft measurements.

We shall add the above information to the revision.

Finally, I have a question for the authors. This manuscript, and many others dedicated to magnetosheath jets, presents statistics of various parameters of jets in order to describe their properties. Thus, it is implicitly recognized that these jets are plasma structures to which we can assign a shape, a size, a bulk velocity, a temperature, a lifetime. My question for the authors is which is, in their opinion, the physical process that grants the jets their individual identity, allowing them to have an individual dynamics, apparently different/independent than/from the background magnetosheath.

Thank you for this enjoyable, a bit more philosophical question, if we may use this characteristic for the Reviewer's words. As we can see from Figures 8 and 9, the jets constitute a region of plasma that has very different properties than the surrounding magnetosheath. It is perhaps like the injection of a bubble of cold air into hotter air, or a low-pressure weather system. Of course it will mix with the surroundings eventually, but it can have a clear identity. We added this information to the Discussion.

I see an issue with the availability of simulations and MMS data on which the findings of this manuscript are based. I am wondering if there is compliance with Annales Geophysicae policy on data availability. Indeed, the authors indicate Acknowledgments that “Data presented in this paper can be accessed by following the data policy on the Vlasov web site.” On the Vlasov website I found the following statement regarding simulations data availability: “All data requests and other

support questions should be addressed to the PI. The PI-team decides about the time and place in which the peer-reviewed data becomes public.” I think the authors should demonstrate the simulations and experimental data are available from a public data repository as requested by Annales Geophysicae data policy (https://www.annales-geophysicae.net/about/data_policy.html), I quote : “It is particularly important that data and other information underpinning the research findings are "findable, accessible, interoperable, and reusable" (FAIR) not only for humans but also for machines. Therefore, Copernicus Publications requests depositing data that correspond to journal articles in reliable (public) data repositories, assigning digital object identifiers, and properly citing data sets as individual contributions. [...] A data citation in a publication resembles a bibliographic citation and needs to be included in the publication’s reference list. [...] In addition, data sets, model code, video supplements, video abstracts, International Geo Sample Numbers, and other digital assets should be linked to the article through DOIs in the assets tab.”

The Vlasiator simulation itself is open source and freely executable by anyone wishing to reproduce these results. One needs to download the source code from the Git repository, follow the guidelines given in the manuscript to make the runs, and then do the postprocessing as explained in the manuscript. The request of the data from the PI works such that a person indicates which data they are interested of, and we will then be able to give that exact data. Our analysis software is open source, and hence the given data can again be post-processed by the requesting side openly and freely. Therefore, we do fulfil the FAIR principles. The reason we do not give all run data openly on a web service is because the size of our run data base so far is well above 200 T. There is no such system currently that would make it possible for us to give these data (and metadata), and to make this kind of a system from scratch would require a huge amount of coding, which is outside of the scientific work required by the funding agencies which fund our work.

Minor remark: Section 2.2. consists on a single paragraph. I suggest the authors expand it and include a description of the MMS data used in their study or integrate the paragraph in the previous section.

We would like to keep this paragraph as it is although it is not an elegant solution. We think the merit in keeping it as is the more structured paper as well as ease in finding information more quickly.

On behalf of all the co-authors,
Minna Palmroth