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A new method to identify flux ropes in space plasmas

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Abstract

Flux ropes are frequently observed in the space plasmas, such as magnetosphere, magnetosheath, and solar wind etc., and play an important role in the reconnection process and mass and flux transportation. One usually used bipolar signature and strong core field to identify the flux ropes. We propose here one new method to identify flux ropes based on the correlations between the variables of the data from in-situ spacecraft observations and the target-function-to-be-correlated (TFC) from the ideal flux rope model. Through comparing the correlation coefficients of different variables at different time and scales, and performing weighted average technique, this method can derive the scales and locations of the flux ropes. We discuss the limitation of our method and also compare it with other methods.

1. Introduction

Flux ropes, as one universal structure in the space plasma, are formed as a helical magnetic structure with magnetic field lines wrapping and rotating around a central axis (e.g., *Hughes and Sibeck*, 1987; *Slavin et al.*, 2003; *Zong et al.*, 2004; *Zhang et al.*, 2010). It is generally believed that flux ropes can be generated by magnetic reconnection in the eruptive energy processes, such as rapid variations of the

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reconnection rate at a single X-line (e.g. *Nakamura and Scholer*, 2000; *Fu et al.*, 2013), multiple X-line reconnection (e.g. *Lee et al.*, 1985; *Deng et al.*, 2004). Flux ropes play important roles in dissipating magnetic energy and controlling the microscale dynamics of magnetic reconnection (e.g., *Drake et al.*, 2006; *Daughton et al.*, 2007; *Fu et al.*, 2017). These structures are frequently observed and widely studied recently in the magnetosphere, magnetosheath and solar wind (e.g. *Hu and Sonnerup*, 2001; *Slavin et al.*, 2003; *Zong et al.*, 2004; *Zhang et al.*, 2010; *Huang et al.*, 2012, 2014a, 2014b, 2015, 2016a, 2016b; *Rong et al.*, 2013). Many works have tried to model flux rope from *in-situ* measurements based on the force-free constant-alpha flux rope (e.g., *Lepping et al.*, 1990), non-force-free model (e.g., *Hidalgo et al.*, 2002), or the Grad-Shafranov equilibrium (e.g., *Hu and Sonnerup*, 2002).

Flux ropes embedded in current sheet are characterized by the bipolar signature of the normal component of magnetic field, strong core field in the axis direction, and enhancement in magnetic field strength. Therefore, one used negative-positive (positive-negative) bipolar signature of the south-north magnetic field component in the earthward (tailward) flow with an enhancement in the cross-tail component and strength of magnetic field to identify flux rope in the magnetotail (e.g., *Slavin et al.*, 2003; *Huang et al.*, 2012). At the magnetopause, the bipolar variation is usually along the sun-earth direction, and the core field is typically along the dawn-dusk direction (e.g., *Zhang et al.*, 2010). However, flux ropes in the magnetosheath, which has been reported recently by MMS (*Huang et al.*, 2016b), can move in any directions due to the large fluctuations of the shocked solar wind. This leads to difficultly in identifying the flux ropes there.

Several attempts are tried to survey flux ropes in the Earth's magnetotail by eyes based on their signatures, such as bipolar variation of north-south magnetic field (e.g., *Richardson et al*., 1987; *Slavin et al*., 2003). Also, some methods are proposed to automatically in some degrees survey the flux rope or flux transfer events (FTEs) via

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bipolar field deflections (e.g., *Kawano and Russell*, 1996; *Vogt et al*., 2010; *Jackman et al*., 2014; *Smith et al*., 2016). *Karimabadi et al*. (2009) have applied data mining technique (MineTool) to search FTEs using magnetic field and plasma data. Recently, *Smith et al*. (2017) developed a method to automatically detect cylindrically symmetric force-free flux ropes in the magnetotail only using magnetic field data. That method first locates the significant deflections in the north-south magnetic field component with peaks in the dawn-dusk component or total field. Then, the candidates are using Minimum Variance Analysis (MVA) to determine a local coordinate system. Finally, the candidates are fitted by a fore-free model to determine whether they belong to flux ropes or not.

For some flux ropes with short duration, the plasma data have not enough high time resolution or even worse are not available. Thus the identification of flux ropes relies heavily on the magnetic field data. All aforementioned automatical methods are a bit complex, or require plasma data. Therefore, to identify flux rope only using the magnetic field data from single spacecraft, we propose a new and simple method based on the correlation coefficients between the signal and the ideal model of flux rope to identify flux ropes in space plasmas. The paper will be presented as follows: an introduction of the method in section 2; the test of the method on artificial data from the model in section 3; the applications of the method on the Cluster and MMS data in section 4; summary is given in section 5.

2. Approach

In this section, we simply introduce our method.

Firstly, we derive target-function-to-be-correlated (TFC) from the ideal model of flux rope. Considering the variable and complicated observed flux ropes, we use the ideal non-force-free model of flux rope proposed by *Elphic and Russell* (1983), named as Elphic and Russell (E-R) Model because most of flux ropes with nonnegligible perpendicular current are not consistent with force-free model (e.g., *Hidalgo et al.*, 2002; *Zong et al.*, 2004; *Zhang et al.*, 2010; *Borg et al.*, 2012; *Huang et al.*, 2012,

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2016b). This model is constructed with an intense core field inside of flux rope, which is shown in Figure 1. The equation of this model in the cylindrical coordinate (Y is defined as the axis orientation of flux rope) can be modified as below:

$$\begin{cases} B_y = B(r)\cos(\alpha(r)) \\ B_\varphi = B(r)\sin(\alpha(r)) \\ B(r) = B_0 \exp(-r^2/b^2) \end{cases} \quad (1)$$

Where α (r) = π/2(1-exp(-r^2/a^2)), B_y is the core field component, B_0, a, and b are the constant, r is the radial distance to the flux rope center.

Figure 1 shows sketched diagram of the cylindrical flux rope from E-R model. For convenience, the rectangular coordinate is used in our analyses (shown in Figure 1). Y is the axis orientation of the flux rope, and the X-Z plane is the cross-section perpendicular to the axis orientation. X can be treated as sun-earth orientation, Y is the dawn-dusk orientation, and Z is similar to the south-north orientation in the magnetotail. If the one spacecraft cross the flux rope following the red path in Figure 1, B_z component will be characterized as bipolar signature, and B_y component has strong peak.

Figure 2 shows the observations when one virtual spacecraft cross the ideal flux rope (see spacecraft path in Figure 1). Here we assume the scale of flux rope as one unit, and 1 unite/s of moving speed of the spacecraft, thus set a = 0.735 s and b = 0.735 s, B_0 =10 nT, and use the B_z as the bipolar variation component, B_y as the core field component, B_t as the total magnetic field. The center of the flux rope is located at 2.5 s. One can see the B_z bipolar signature, and the peak of core field and total magnetic field inside the flux rope.

Considering the previous observations, in which the B_z component during the crossing of the flux rope usually does not reach zero like that shown in Figure 2a, we cut out one part of the ideal flux rope as the TFC which is shown in Figure 3. The TFC is similar to the sinusoidal function when one performs Fast Fourier Transform (FFT)

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analysis. We only used two components (B_y and B_z) and magnetic strength (B_t) as the TFC since only B_z and B_y components and B_t have very obvious typical feature usually from in-situ measurements (i.e., B_z has bipolar signature, B_y is strong core field, and B_t has peak inside flux rope), and B_x component has not common feature from observation viewpoint (e.g., *Slavin et al.*, 2003; *Huang et al.*, 2014a).

Secondly, we calculate the correlation coefficients between the signal and the TFC at different time and different scales. Before calculating the correlation coefficients, the amplitude of the TFC will be estimated from the signal. For example, the maximum value of B_t during the time interval is used as the amplitude of B_t in the TFC. The sliding time window is used in the calculation of the correlation coefficients. The calculated results of correlation coefficients are similar to the power spectral densities by FFT that displays the power spectral density at different time and different frequency. The higher values of the correlation coefficients, the more suitable for the description of the model on the signal.

Thirdly, we compare the correlation coefficients of the bipolar variation component B_z, core field component B_y, and total magnetic field B_t, and find out the high correlations (larger than the given threshold) at the same time and the same scale. This is due to that the bipolar signature in B_z, the enhancements of core field B_y and magnetic strength B_t should appear simultaneously with the same duration when one spacecraft cross the flux ropes.

Fourthly, we infer the location and the scale of the flux ropes based on the weighed average (it will be shown later), and the amplitude from minimum to maximum values of the bipolar variation.

3. Model test

One test is performed on the artificial data from E-R model plus the random noise. Figure 4 presents the test results. The test artificial data is shown in Figure 4a where

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the noise is 10% of the amplitude of the flux rope. A series of the calculations are carried on B_z, B_y and B_t to obtain the correlation coefficients. One should point out that the absolute values of the correlation coefficients of B_z and B_y are given in Figure 4b and 4c respectively, because the bipolar structure can be positive-negative or negative-positive variation and the core field can be positive or negative. It can be seen that the correlation coefficients are largest at the scale of 0.6 ~ 1.5 s during the crossing of the flux rope (around *time* ~ 3.5 s).

We set the threshold as 0.9 to represent the results in Figure 5 where only the correlation coefficients with > 0.9 are displayed with black shadows. All correlation coefficients of the three variables have peaks at the *time* ~ 3.5 s with the time scale ~ 1 s. We use the weighted average technique (shown below) to identify the flux rope and estimate its time scale.

$$\tau = \sum coef_i \times \tau_i / \sum coef_i \tag{2}$$

where $coef_i$ is the correlation coefficient at time scale τ_i.

Figure 5e shows the estimated results. The crossing of the flux rope is marked with "1" and the duration is its scale, the center of the flux rope is at the center of the line. In this test, the scale is estimated as 1.039 s, the location is 3.496 s. The amplitude is estimate as 4.43 nT from minimum to maximum values of the bipolar variation. Aforementioned sets, one can estimate the error of the scale as 3.9%, i.e., (1.039-1.0)/1.0 = 3.9%. Therefore, our method can successfully identify the flux rope, and estimate its scale, location and amplitude.

4. Application

In this section, we apply our new method to the spacecraft measurements in the magnetosheath and the magnetotail.

4.1 Flux rope in the magnetosheath

Flux ropes are successfully identified in the magnetosheath using the unprecedented

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high resolution data from Magnetospheric Multiscale (MMS) (*Burch et al.*, 2016) mission (*Huang et al.*, 2016b). Their observations demonstrate that highly dynamical, strong wave activities and electron-scale physics occur in the magnetosheath ion-scale flux ropes. Figure 6 gives the observations of ~14 s from MMS2 on 25 Oct 2015 and the test results of our method. Similar to the model test, we use the same variables to present the components of the bipolar variation, core field and total magnetic field after transformed to minimum variable analysis (MVA) analysis (*Huang et al.*, 2016b). The threshold of the correlation coefficients is also set as 0.9 in Figure 6. We can see that the correlation coefficients of the three variables (Figure 6b-6d) only have high values at the same time around *time* = 5.5 s, implying that one flux rope is identified by this method. Based on the weighted average method in equation (2), the time scale of the flux rope is 1.11 s, and its central location is at 5.38 s. The amplitude is estimate as 115 nT. All these results are consistent with previous findings from multi-spacecraft data in *Huang et al.* (2016b).

4.2 Flux rope in the magnetotail

Flux ropes are frequently observed in the magnetotail, and play an important role during magnetic reconnection and magnetotail dynamic (e.g., *Slavin et al.*, 2003; *Zong et al.*, 2004; *Chen et al.*, 2008; *Huang et al.*, 2012, 2016a; *Fu et al.*, 2015, 2016). Chen et al. (2008) have identified several flux ropes filled with energetic electrons during magnetic reconnection on 10 Jan 2001 by using the Cluster data. Figure 7 shows the magnetic field in GSM coordinates from the Cluster mission (*Escoubet et al.*, 1997) in the magnetotail and the application results of our method. There are several bipolar variations in B_z during this time interval (Figure 7a). Figures 7b-7d present the correlation coefficients (larger than 0.9 of the threshold) of the three variables. Here we try to identify small-scale flux ropes, so that we perform the method only at short time scale. There are full of high correlation coefficients (grey shadows) in Figures 6b-6d. After compare with the correlation coefficients at the same time and same scale, our method resolves three possible flux ropes in Figure 7e. The results are summarized in Table 1. The three structures are close to ideal flux

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rope with bipolar signature in B_z, and peaks in core field B_y and total magnetic field B_t. All three flux ropes identified by our method have been reported in *Chen et al.* (2007).

We should point out that our method only can identifies the flux rope and derives its duration. If the plasma velocity data is available, then we can estimate the actual spatial scale of the flux rope. If multi-spacecraft data are available for the time interval of interest, one can derive the size, the orientation, and the motion of the flux rope using by the multi-spacecraft method such as *Sonnerup et al.* (2004), *Shi et al.* (2005, 2006) and *Zhou et al.* (2006a, 2006b). However, the separation of the Cluster was much lager than the size of the flux ropes on 01 October 2001, implying that one cannot use multi-spacecraft method here.

5. Summary and Discussion

In summary, we developed a new method to identify flux ropes in the space plasmas. This method is based on the correlation coefficients between the signal and the TFC from E-R model. If the correlation coefficients of three variables (B_z, B_y and B_t) of the signal have high values of correlation coefficients at the same time and same scale, one can deduce the existence of one flux rope and estimate its location and its time scale (i.e., the duration). The tests on the artificial data and the in-situ realistic spacecraft data show that our method can successfully search out the flux ropes and obtain their locations and time scales.

Bipolar variation in B_z component and the enhancement in core field and magnetic field strength are the typical signatures for most of flux rope. But it doesn't mean that all observations from any crossing of the spacecraft would have those signatures, which depends on the spacecraft trajectory. However, one only can select or identify the flux rope showing the typical signatures, and miss other flux rope not having the typical signatures. Some special field structures may induce the similar signatures along some special trajectories. But we thought this opportunity is too few in the

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magnetotail. However, one can use the plasma measurements to rule out this possibility.

Aforementioned attempts are used to identify flux ropes in the Earth's magnetotail by eyes or half-automatically based on the bipolar variation of (e.g., *Richardson et al.*, 1987; *Slavin et al.*, 2003; *Kawano and Russell*, 1996; *Vogt et al.*, 2010; *Jackman et al.*, 2014; *Smith et al.*, 2016). *Karimabadi et al.* (2009) used data mining technique (MineTool) to search flux ropes using both magnetic field and plasma data. That method is too complex to apply in the data analysis. *Smith et al.* (2017) proposed one method to automatically detect force-free flux ropes based on magnetic field data from single spacecraft. In present study, we used the TFC derived from non-force-free flux rope model to calculate the correlation coefficients with the signal, and then compare the large correlation coefficients of different variables to identify the flux rope. Our method is flexible and easy to apply in the in-situ spacecraft data compared with other methods. We will quantitatively model the flux ropes identified by our method and derive more information of the flux ropes.

We should point out that there are several limitations in our method.

1. Our method can only detect the nearly ideal cylindrical flux rope since we used non-force-free E-R model to describe the TFC, which limits the application of this method. The non-force-free model proposed by E-R is just one possible solution of all the flux rope that satisfies $J \times B \neq 0$. Actually one can use the other flux rope models to replace E-R model, and extend our method to identify the flux ropes.

2. If the flux ropes are not well regular, there are large time deviations between B_z, B_y and B_t which will lead to miss of some flux ropes when we apply the method.

3. The threshold value of correlation coefficients can affect the results. When the threshold value is too small that the method finds out some possible structures which do not belong to flux ropes, or too large that the method will miss some flux ropes.

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4. The correlation coefficients at small scale (especially in B_y and B_t) could be very large, which may affect our results. The method may find some possible structures related to such fluctuations. We will improve this method and apply it to detect the flux ropes in the turbulent magnetosheath in the future.

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Table 1. The location, sale and amplitude of the flux ropes identified by the method. The amplitude is defined as the values of the bipolar variation from minimum to maximum.

# of flux rope	1	2	3
Location [s]	37.91	113.79	127.93
Scale [s]	1.99	2.84	2.05
Amplitude [nT]	9.96	20.49	12.59

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Figure captions

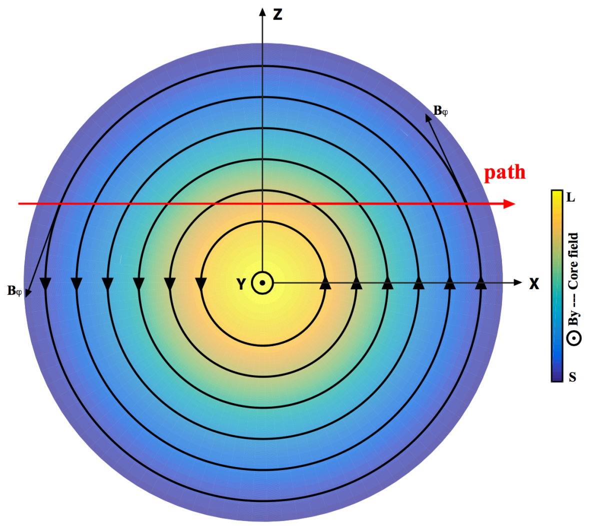

Figure 1. Sketched diagram of the cylindrical flux rope. The flux rope is right-hand handedness structure. The black circled lines are the magnetic field lines. The red arrow is the projection of spacecraft path. The rectangular coordinate is used in our analyses. Y is the axis orientation of the flux rope, and the X-Z plane is the cross-section perpendicular to the axis orientation. The core field is out-of-plane, and the color represents the relative strength of core field (yellow: large, blue: small).

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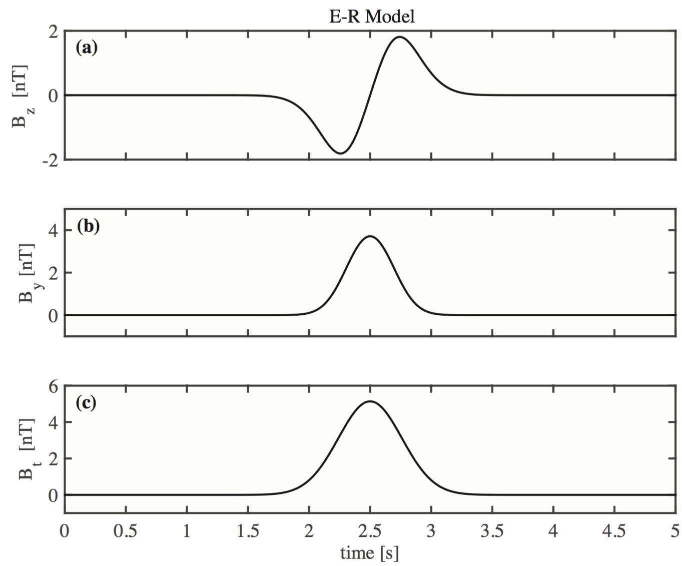

Figure 2. The three variables B_z (a), B_y (b), and B_t (c) of the ideal cylindrical flux rope described by E-R model.

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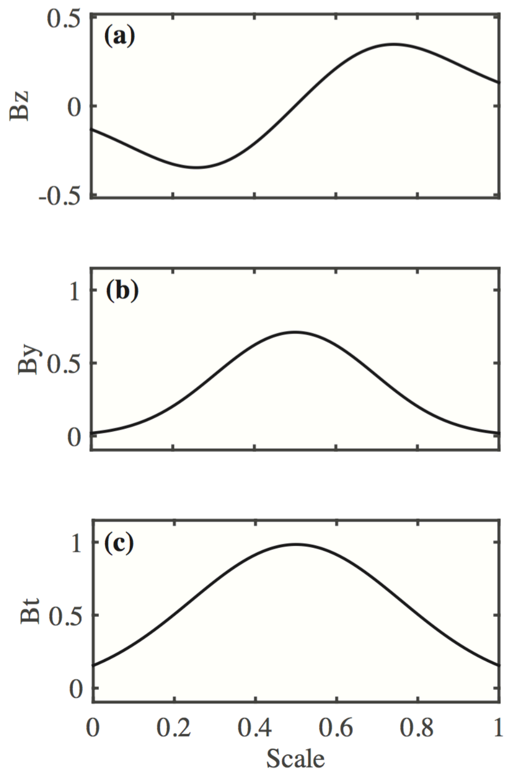

Figure 3. The target-function-to-be-correlated (TFC) derived from E-R model. The amplitudes and scale are dimensionless.

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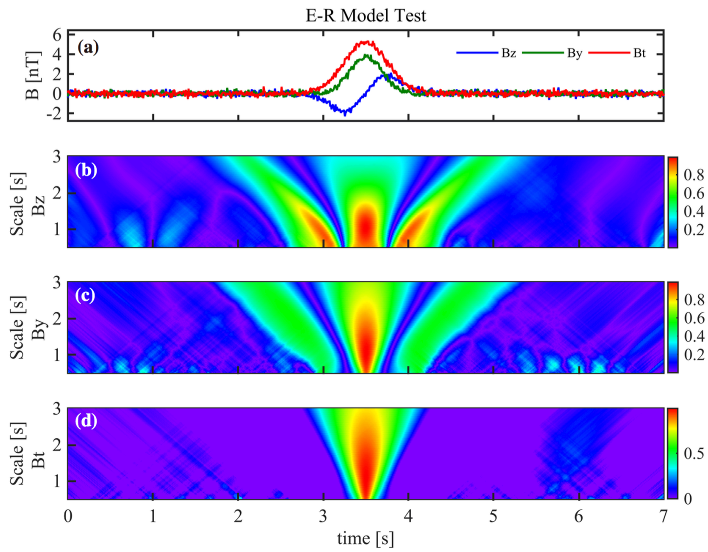

Figure 4. The test results on E-R model. (a) three variables B_z, B_y, and B_t from E-R model with 10% random noise; (b-d) the correlation coefficients between the variables of B_z, B_y, and B_t and the TFC shown in Figure 3, respectively.

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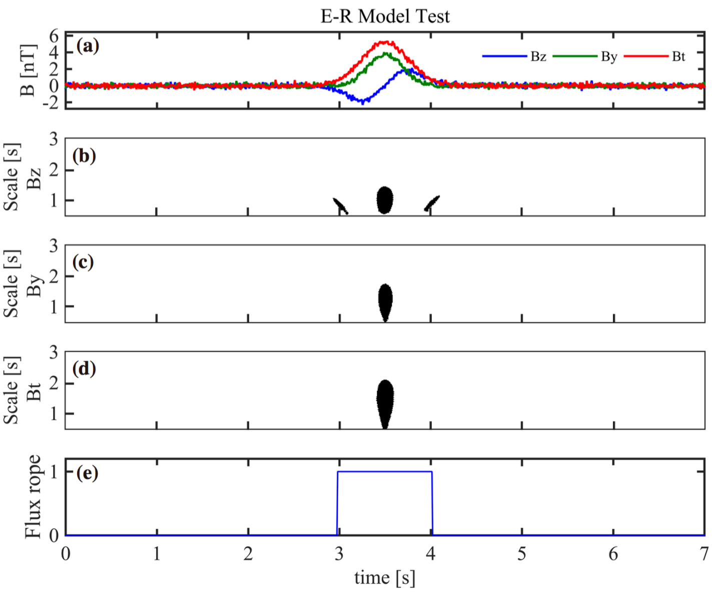

Figure 5. The test results on E-R model with a threshold 0.9. (a) three variables B_z, B_y, and B_t from E-R model with 10% random noise; (b-d) the correlation coefficients (≥0.9) between the variables of B_z, B_y, and B_t and the TFC, respectively; (e) the index when the virtual spacecraft cross the flux rope (if the spacecraft cross the flux rope, the index is 1; if not, the index is 0). The duration of the index presents the time scale of the flux rope.

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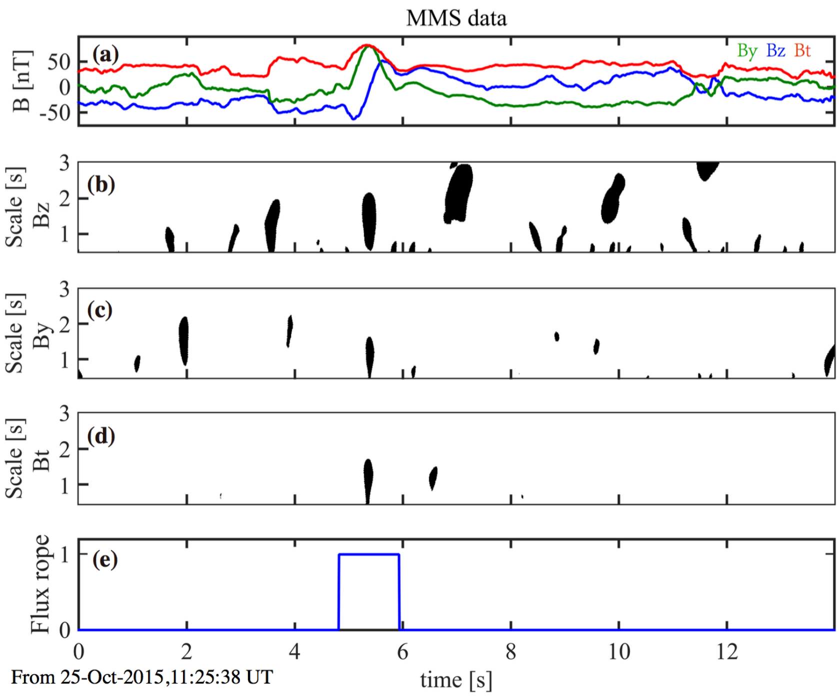

Figure 6. Testing the method on MMS data in the magnetosheath. The same format as in Figure 5.

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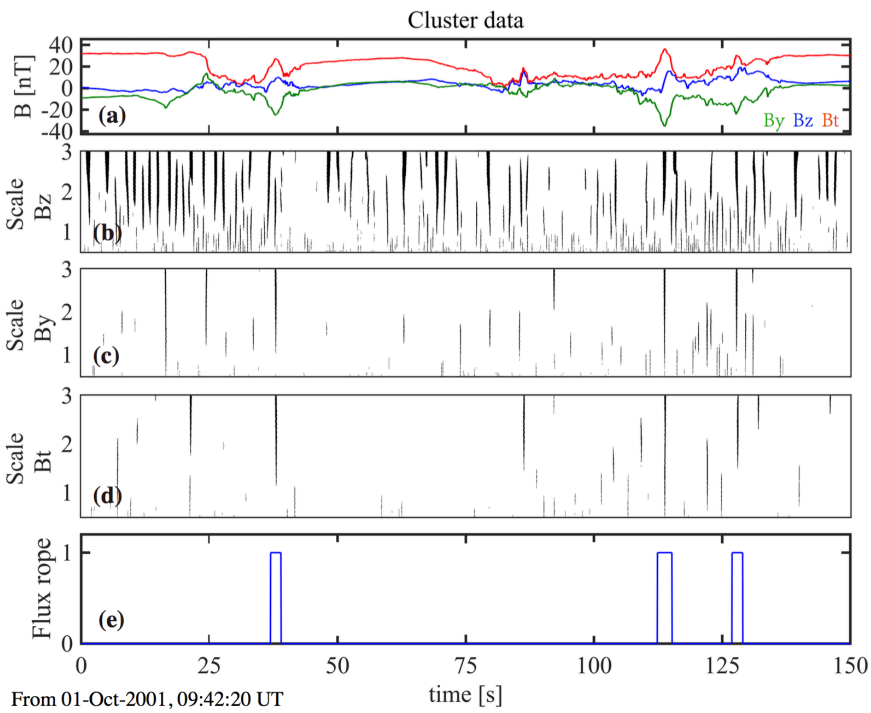

Figure 7. Testing the method on Cluster data in the magnetotail. The same format as in Figure 5.