



Ionospheric plasma flows associated with the formation of the distorted nightside end of a transpolar arc

Motoharu Nowada¹, Adrian Grocott², and Quan-Qi Shi¹

¹Shandong Key Laboratory of Optical Astronomy and Solar-Terrestrial Environment, School of Space Science and Physics, Institute of Space Sciences, Shandong University, Weihai, Shandong 264209, P.R. China

²Space and Planetary Physics Group, Department of Physics, Lancaster University, Lancaster LA1 4YW, UK

Correspondence: Motoharu Nowada (moto.nowada@sdu.edu.cn)

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Abstract. We investigate ionospheric flow patterns occurring on 28 January 2002 associated with the development of the nightside distorted end of a J-shaped transpolar arc (nightside distorted TPA). Based on the nightside ionospheric flows near to the TPA, detected by the SuperDARN (Super Dual Auroral Radar Network) radars, we discuss how the distortion of the nightside end toward the pre-midnight sector is produced. The J-shaped TPA was seen under southward interplanetary magnetic field (IMF) conditions, in the presence of a dominant dawnward IMF- B_y component. At the onset time of the nightside distorted TPA, particular equatorward plasma flows at the TPA growth point were observed in the post-midnight sector, flowing out of the polar cap and then turning toward the pre-midnight sector of the main auroral oval along the distorted nightside part of the TPA. We suggest that these plasma flows play a key role in causing the nightside distortion of the TPA. SuperDARN also found ionospheric flows typically associated with Tail Reconnection during IMF Northward Non-substorm Intervals (TRINIs) on the nightside main auroral oval, before and during the TPA interval, indicating that nightside magnetic reconnection is an integral process to the formation of the nightside distorted TPA. During the TPA growth, SuperDARN also detected anti-sunward flows across the open–closed field line boundary on the dayside that indicate the occurrence of low-latitude dayside reconnection and ongoing Dungey cycle driving. This suggests that nightside distorted TPA can grow even in Dungey-cycle-driven plasma flow patterns.

1 Introduction

Transpolar arcs (TPAs) are the bar-shaped part of theta aurora, connecting the nightside and dayside auroral ovals within the polar cap (Frank et al., 1982). Since theta auroras were discovered at the beginning of 1980s, TPAs have been the focus of much research, and various formation mechanisms have been proposed based on investigations of the ionospheric flow patterns and the relationship to the orientation of interplanetary magnetic field (IMF; see a series of reviews on polar cap arcs and TPAs by Hosokawa et al., 2020; Fear and Milan, 2012a; Mailyan et al., 2015).

The TPA formation model, based on nightside magnetic reconnection occurring under northward IMF conditions, which was proposed by Milan et al. (2005), has had a high degree of success in explaining a wide variety of TPA observations (e.g., Fear and Milan, 2012a, b; Kullen et al., 2015; Nowada et al., 2018, and references therein). In this model, nightside magnetic reconnection forms closed magnetic field lines, whose northern and southern footpoints straddle the midnight meridian. As a result, the newly closed flux has no preferential return path to the dayside (i.e., via dawn or dusk) and instead protrudes into the magnetospheric lobe and thus into the polar cap ionosphere. This protruding closed flux becomes what we call the TPA, which, in the simplest case, grows straight from the nightside to the dayside in the polar cap. In the ionosphere, azimuthal plasma flows across the midnight meridian, normally ranging between about 300 and 700 m s^{-1} but are sometimes faster than 700 m s^{-1} , are observed in the nightside auroral oval. These characteristic ionospheric flows are interpreted as evidence for nightside

reconnection (e.g., Grocott et al., 2003, 2004) and are often referred to as the flow signatures associated with Tail Reconnection during IMF Northward Non-substorm Intervals (TRINNI; e.g., Milan et al., 2005). Such flows are observed at the poleward edge of the main nightside auroral oval, which is the boundary between open and closed magnetic flux, in the vicinity of the growth point of the TPA. This indicates that the magnetotail magnetic reconnection occurs close to the region of the TPA formation (Milan et al., 2005; Fear and Milan, 2012b). However, we might also expect to find a region of much slower flow at the site of the TPA growth itself, since the TPA formation mechanism is directly related to a stagnation of the newly closed magnetic flux and associated plasma flows.

TPAs are sometimes seen during southward IMF intervals. However, in most of those cases, there had been prolonged northward IMF intervals before the TPA occurrences (e.g., Craven et al., 1991; Newell and Meng, 1995; Pulkkinen et al., 2020, and references therein). Certainly, in particular southward IMF TPA cases discussed by Craven et al. (1991) and Newell and Meng (1995), the magnetospheric and ionospheric dynamics triggered by the change in IMF orientation from northward to southward seemed not to have played an essential role in the TPA formation processes. However, using a combination of auroral imager observations and magnetohydrodynamic (MHD) global simulations of magnetotail dynamics, Pulkkinen et al. (2020) recently suggested that fast plasma flows triggered by strong magnetotail reconnection in the distant magnetotail may be a source of TPAs under southward IMF conditions.

In contrast to these straightforward TPAs (hereafter referred to as regular TPA) frequently occurring under northward IMF conditions, bending or hooked-shaped arcs have also been reported. These arcs grow from the dawnside or duskside main auroral oval to the dayside. They are observed when the IMF orientation is southward or when it turns from long-term southward (northward) to northward (southward) in the presence of a dominant IMF- B_y component (Kullen et al., 2002, 2015; Carter et al., 2015). Their formation can be explained by the magnetic reconnection at the low-latitude dayside magnetopause (Kullen et al., 2015; Carter et al., 2015). Carter et al. (2015) proposed the detailed formation process of bending arcs, which are formed by the entry of the solar wind (magnetosheath) particles along open field lines generated by low-latitude dayside reconnection. Their growth toward pre- or post-noon is caused by dawn–dusk asymmetric ionospheric plasma convection caused by the presence of IMF- B_y penetration (e.g., Cowley and Lockwood, 1992). TRINNI flows were not found during the bending arc development, suggesting that nightside magnetic reconnection is not related with the formation of bending arcs. (Kullen et al., 2015; Carter et al., 2015).

Nightside distorted TPAs are duskside (dawnside) TPAs, with their nightside ends distorted toward post- (pre-) midnight, and were first identified based on a statistical study by

Fear and Milan (2012b). Nowada et al. (2020) proposed a possible formation scenario of the nightside distorted TPAs. According to their scenario, the essential source of nightside distorted TPAs is upward (flowing out of the ionosphere to the magnetotail) field-aligned currents (FACs), which are generated by plasma flow shear between fast plasma flows triggered by magnetotail magnetic reconnection and slower background magnetospheric flows. They also postulated that the TPA growth to the dayside is attributed to the retreat of the magnetotail reconnection points to further down the magnetotail. During the development of nightside distorted TPAs, as the reconnection site goes further tailward, the magnetotail becomes more deformed, and associated field lines are also twisted more strongly (Tsyganenko et al., 2015; Tsyganenko and Fairfield, 2004), caused by the IMF- B_y penetration (Gosling et al., 1990; Cowley, 1981, 1994). Nowada et al. (2020) concluded that, owing to the magnetotail deformation and field line twisting, the TPA does not grow straightforwardly from the nightside main auroral oval to the opposite dayside oval but develops with a distortion of its nightside end toward dawn or dusk. However, in situ observational evidence on the TPA deformation is yet to be detected in either the magnetosphere or ionosphere.

The nightside distorted part of the TPA is frequently aligned with the main auroral oval but is a distinct feature at the auroral oval's poleward edge. This might be related to other observations where a part of the nightside auroral oval appears as bifurcated branches, equatorward and poleward, with a gap (or a weak emission region) between them. This separated auroral feature is identified as a double auroral oval (e.g., Elphinstone et al., 1995a, b). Ohtani et al. (2012) investigated the detailed electric current structures and formation mechanism of a double auroral oval seen in the dusk–midnight sector. Such double auroral ovals are frequently seen under geomagnetically active conditions and during the latter (recovery) phase of intense polar substorms. The equatorward branch of the double auroral oval is embedded in upward field-aligned currents (FACs), flowing out of the ionosphere to the magnetotail, with downward FACs from the magnetotail to the ionosphere dominantly collocated in the double oval poleward branch. Each branch of the bifurcated auroral oval therefore connects to different region of the magnetotail. Ohtani et al. (2012) concluded that the equatorward branch connected the field lines at geosynchronous altitudes and outside, corresponding to the ring current and the near-Earth part of the tail current, while the poleward double oval branch mapped to a wide area farther downtail, where accelerated auroral particles precipitating to the poleward branch are generated.

In this paper, we report a significant finding in relation to the ionospheric plasma flows that may explain the generation of the nightside distorted end of the TPA. This study is achieved using ionospheric flow patterns measured by the SuperDARN (Super Dual Auroral Radar Network) HF (high frequency) radars and the auroral imager data obtained from

the Imager for Magnetopause-to-Aurora Global Exploration (IMAGE) satellite.

This paper consists of six sections. The introduction and the instrumentation used in this study are given in Sect. 1 and 2. The introduction of the nightside distorted TPA is shown in Sect. 3. In Sect. 4, the observational results of solar wind conditions and global ionospheric plasma flows associated with the formation of nightside distorted TPA are reported. Finally, in Sects. 5 and 6, we present discussions and our conclusions of this study.

2 Instrumentation and data processing

2.1 Auroral images

Nightside distorted TPAs were identified using auroral observations by the Wideband Imaging Camera (WIC), which is part of the Far Ultraviolet (FUV) instrument (Mende et al., 2000a, b, c), on board the Imager for Magnetopause-to-Aurora Global Exploration (IMAGE), launched in March 2000. IMAGE FUV-WIC imaged the aurora in a broad wavelength range from 140 to 190 nm with 2 min cadence. The IMAGE FUV-WIC data include non-auroral optical signals due to sunlight (dayglow) and the instrumental optical noise. In this study, we have removed non-auroral data as much as possible, using the methods described in Nowada et al. (2020).

2.2 Ionospheric convection maps

Ionospheric plasma flow data were obtained by the SuperDARN HF radars (Greenwald et al., 1995; Chisham et al., 2007). These high-latitude radar arrays in both the Northern and Southern hemispheres make line-of-sight measurements of ionospheric flow velocity. For this study, data from all radars in the Northern Hemisphere were combined using the map potential technique (Ruohoniemi and Baker, 1998), which fits an eighth order spherical harmonic expansion of the ionospheric electric potential to the measured flows to provide large-scale maps of the ionospheric convection pattern. This is achieved by first median averaging the line-of-sight data onto an equal area magnetic latitude and longitude grid, within cell size $\sim 110 \times 110$ km, to remove anomalous data. A lower-latitude boundary to the convection is then estimated from the distribution of measured velocities as the lowest latitude at which a threshold of at least three measurements of $\sim 100 \text{ m s}^{-1}$ is met. The background statistical model of Ruohoniemi and Greenwald (1996) is then used to provide a set of vectors that supplement the observations to provide enough measurements for the spherical harmonic fit to converge.

As recently discussed by Walach et al. (2022), the map potential solutions are sensitive to a number of factors that govern the resulting convection maps. One factor is the latitudinal extent of the radar coverage. For example, more re-

cent additions to the radar network at mid-latitudes can improve estimates of the flow in these regions. However, in this study, only the auroral zone SuperDARN radars had been built. This is unlikely to affect our results, owing to the relatively contracted polar cap in this case; the flows of interest were located at $\sim 70^\circ$ magnetic latitude and thus well within the fields of view (FOV) of the auroral zone radars. Another factor is the placement of the equatorward boundary of the convection – the so-called Heppner–Maynard boundary (HMB; Heppner and Maynard, 1987). This can be influenced by irregular data coverage and also by the inclusion of slower E-region scatter, which can be found at near-ranges in the radar FOV. We therefore carefully inspected the placement of the HMB in our analysis and modified the automatically generated boundary to remove unphysical steps that occasionally occurred due to the inclusion of possible E-region contamination (near-range and slow flows). We further compared the boundary to the auroral images and found that it was generally located close to the equatorward edge of the auroral oval, as expected. Finally, we also considered the choice of background model, something that Walach et al. (2022) found could influence the map potential output in regions where the number of measurements is low. We decided upon the Ruohoniemi and Greenwald (1996) model over more recent options for two main reasons. First, this model was derived from auroral zone data for only a few years prior to our interval. This might be expected to make it more appropriate than more recent models, such as Thomas and Shepherd (2018), which are constructed from data based on a different solar cycle and include data from different geophysical regions. Second, the model proposed by Ruohoniemi and Greenwald (1996) explicitly focuses our interpretation on regions in which direct measurements exist, such that the choice of model has less influence on the resulting flow vectors anyway.

All SuperDARN convection maps shown in this study were produced based on the aforementioned processing steps. Each map includes streamlines of the electric equipotentials and their values with black solid and broken contours on the dusk and dawn sides, respectively. Convention is to assign the opposite signs to the ionospheric electrostatic potential with positive and negative vorticity; the electric potential at dawn is generally positive (maximum potential denoted by a plus sign) and is generally negative (minimum shown by a cross) at dusk. To estimate the two-dimensional flow velocities, we use the available radar line-of-sight measurements, with the transverse velocity component derived from the map potential solution. We choose to present these composite vectors, rather than the commonly used global electrostatic solutions, because we are mainly interested in the flows driven by dynamic magnetotail processes that are not well resolved in the global patterns.

3 Overview of nightside distorted TPAs

The first (left) panel in Fig. 1 shows an example of a regular TPA, with straight bar-shaped emissions, connecting the nightside and dayside auroral ovals. In nightside distorted TPAs, as shown in the second and third panels, their nightside ends become distorted toward the pre- or post-midnight sectors, respectively. Nowada et al. (2020) identified the TPAs with the nightside distortions as being J- and L-shaped TPAs, based on their resemblance to the letters J and L. Taking a look at the J- and L-shaped TPAs, the distorted directions of their nightside ends are, of course, opposite to each other, but otherwise no significant difference can be seen, particularly in the emissive parts that straightforwardly cross the polar cap to the dayside. Most cases of nightside distorted TPA were observed during the northward IMF intervals with a dominant IMF- B_y component (Nowada et al., 2020). However, some regular and nightside distorted TPAs can be observed, even under southward IMF but usually where the IMF orientation had previously been persistently northward (Fear and Milan, 2012a, b, or see the nightside distorted TPA event list in Nowada et al., 2020).

4 Observations of solar wind conditions and ionospheric flows associated with the nightside distorted TPA

4.1 Formation of the closed magnetic field within the J-shaped TPA

Figure 2 shows a case of a nightside distorted TPA (J-shaped TPA), which was observed on 28 January 2002, along with the corresponding geomagnetic activity, IMF conditions, and global ionospheric flows. We identified the onset time (09:25:27 UT) of this distorted TPA based on the IMAGE FUV-WIC data via visual inspection. In Fig. 2a, the geomagnetic activity and IMF conditions obtained from the OMNI solar wind database during a 4 h interval between 07:30 and 11:30 UT are shown. The time interval corresponding to the nightside distorted TPA observation is bracketed by two gold broken lines. The black dotted lines with the labels from (a) to (h) indicate eight key times of interest discussed later and shown in Fig. 3. During the interval of the nightside distorted TPA, the magnetosphere was geomagnetically quiet, with AL and AU index values smaller than -20 and 35 nT, respectively. The associated IMF- B_y component was dominantly negative (dawnward).

The dip in the AL index down to -150 nT, between 07:30 and 08:10 UT, suggests that an auroral substorm occurred more than 1 h prior to the TPA onset time (09:25 UT). The AL magnitude subsequently decreased, that is, the substorm entered the recovery phase from 08:10 to 08:50 UT. A larger substorm, whose AL peak was over -240 nT, occurred after the disappearance of the TPA (10:31 UT). Therefore, the

J-shaped TPA was seen during a geomagnetically quiet interval in the polar region between two auroral substorms. Such quiet magnetospheric conditions are favorable for the formation of nightside distorted TPAs.

During the interval from at least 1 h prior to the onset of the arc ($\sim 08:20$, $\sim 09:20$ UT), the IMF- B_z component was predominantly northward. However, from 09:20 UT, the IMF- B_z component turned to, and persisted in, a southward orientation. The J-shaped TPA was seen under almost entirely southward IMF conditions, although the IMF- B_z component transiently turned northward just after the onset of the nightside distorted TPA, and fluctuated between northward and southward directions until $\sim 09:42$ UT. After these fluctuations, the orientation of the IMF- B_z component persisted southward. During the TPA interval, the average clock angle was $\sim -120^\circ$, due to a strong IMF- B_y component. The solar wind plasma conditions were stable during the TPA interval (not shown here).

Figure 2b shows an example of the plasma flow velocity vectors between 09:54 and 09:56 UT from the nightside ionosphere covering the magnetic local time (MLT) range, from 18 to 6 h, overlaid on a greyscale image of main auroral oval and J-shaped TPA detected by IMAGE FUV-WIC in the Northern Hemisphere at 09:54:08 UT. The time of this image corresponds to the gold solid vertical line on the solar wind data in Fig. 2a. The J-shaped (nightside distorted) TPA is comprised of the bar part, which is growing toward the dayside, with a slight dawnward sense in the post-midnight sector and the nightside end, which is distorted toward the pre-midnight sector. The associated velocity vectors are projected onto a geomagnetic grid. The green curve shows the HMB, that is, the lower latitude limit for the ionospheric plasma convection pattern. The vector length and color code are assigned according to the intensity of ionospheric flow velocity in units of meters per second (m s^{-1}). Around the poleward edge of the main auroral oval, westward flows from the post-midnight to the pre-midnight sector, whose speed is between 400 and 750 m s^{-1} , were observed. These ionospheric flow signatures can be seen to have crossed from the poleward of the main aurora oval just post-midnight and adjacent to the TPA growth point. Assuming that the poleward edge of the auroral oval is a proxy for the boundary between the open and closed magnetic flux regimes, then these flows would thus appear to be associated with magnetotail reconnection. Although the flows here occurred during an interval of southward IMF, they have the same significant characteristics as the ionospheric flow signatures of Tail Reconnection during IMF Northward Non-substorm Intervals (TRINNI)s flows (e.g., Grocott et al., 2003, 2004). In this case, westward fast plasma flows at the poleward edge of the main auroral oval across the midnight meridian. According to the average statistical picture of the ionospheric return flow, given by Reistad et al. (2018), TRINNI-type return flows can be seen even under the southward IMF, as long as the IMF- B_y component is present (as is the case shown here). Further-

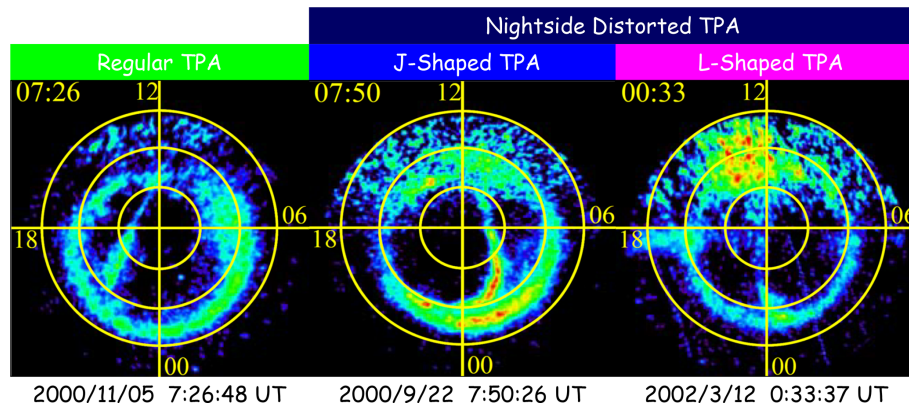


Figure 1. Representative examples of the IMAGE FUV-WIC observations for three types of the transpolar arc (TPA) morphologies. Regular TPA (left) and nightside distorted TPAs, including J- (center) and L-shaped (left) TPAs, as observed on 5 November 2000, 22 September 2000, and 12 March 2002, are shown, respectively. These nightside distorted TPAs are referenced from the nightside distorted TPA database, as provided in Nowada et al. (2020). The top and bottom sides in each panel show the noon and midnight sides, for which the magnetic local times (MLTs) are 12 and 24 h, respectively. The dawn (6 h MLT) and dusk (18 h MLT) meridians correspond to the right and left sides, respectively. The concentric circles in each panel show the magnetic latitude (MLat) at 60, 70, and 80°, respectively. The color code is expressed in analog-to-digital units (ADUs), which are proportional to the observed auroral brightness (see the details in Mende et al., 2000b). The dayglow and background optical noises from the IMAGE FUV-WIC data are removed by the techniques described in Nowada et al. (2020).

more, the TRINNIs occur within global ionospheric convection flow, excited by the ongoing and modest dayside reconnection, and under an absence of substorm activity, which are all true of the present interval. The presence of TRINNI-type flows explicitly suggests that nightside magnetic reconnection has occurred and, thus, plays a role in the formation of this J-shaped TPA, providing a source of closed flux, as previously proposed by Nowada et al. (2020). Considering that the dynamic properties of a TPA within the polar cap, such as its drift motion, are governed by global ionospheric flows (e.g., Milan et al., 2005; Fear et al., 2015), our observation of plasma flows near the nightside end of a TPA may be the key phenomenon to understanding the cause of the distortion. The ionospheric flow velocity on the nightside distorted part of the TPA was faster than that on the poleward edge of the main auroral oval. Therefore, the nightside distorted part appears to be distinct from the main auroral oval in this case.

Figure 3 shows a time series of the overlaid plots of IMAGE FUV-WIC and SuperDARN radar data from 70 min prior to the onset time of the J-shaped TPA at 09:25:27 UT. A number of features can be seen in the nightside flow during this time. In Fig. 3a and b, 70 to 60 min prior to the J-shaped TPA onset time, any significant flows seem to be restricted to the pre-midnight sector, associated with structure in the brighter parts of the pre-midnight auroral oval. These intensifications of the poleward boundary, that is, the poleward boundary intensification (PBI), are likely to be features of the recovery phase of the preceding substorm (e.g., Lyons et al., 1999). The flows in the post-midnight region have become enhanced by Fig. 3c, about 50 min prior to the TPA onset. These flows seem to cross the polar cap bound-

ary at around 01:00 ~ 02:00 MLT. In Fig. 3d, about 40 min prior to onset, the flows can be seen to have moved to earlier local times, in concert with the PBI also having moved to an earlier MLT (~ 00:00 MLT). Now, flows across the poleward edge of the main auroral oval around the midnight sector can be seen; these are highlighted by the zoomed-up flow profiles in the orange-framed region. These flow signatures were observed until the onset of the TPA, as also shown in the orange-framed magnifications of Fig. 3e–h. From 30 min prior to the TPA onset (Fig. 3e–h), these flows evolve into the classical signature of TRINNIs – oval-aligned return flows – that act to remove and redistribute the closed flux into the auroral zone at earlier local times.

The ongoing observation of TRINNIs from –30 min up to the onset time of the TPA suggests that the nightside reconnection persisted around the TPA growth point. However, of particular significance is that, at the times shown in Fig. 3g and h, there also exists a region downward of the TRINNI return flows that consists of slow (and stagnant) flows lower than 250 m s^{-1} , as indicated with blue-colored vectors. Subsequently, the TPA began to grow within these stagnant flow regions, which is consistent with the conventional TPA formation model based on nightside magnetic reconnection, as proposed by Milan et al. (2005). Indeed, this stagnant flow region can still be seen in the flow pattern at 09:47:59 UT in Fig. 5b, when the TPA was still growing toward the dayside. Based on these TRINNI flow profiles detected by the SuperDARN radar measurements, we can infer that nightside reconnection occurred before and at the TPA onset time and is thus likely to be part of the mechanism for the formation of the J-shaped TPA. This is consistent with the fundamental

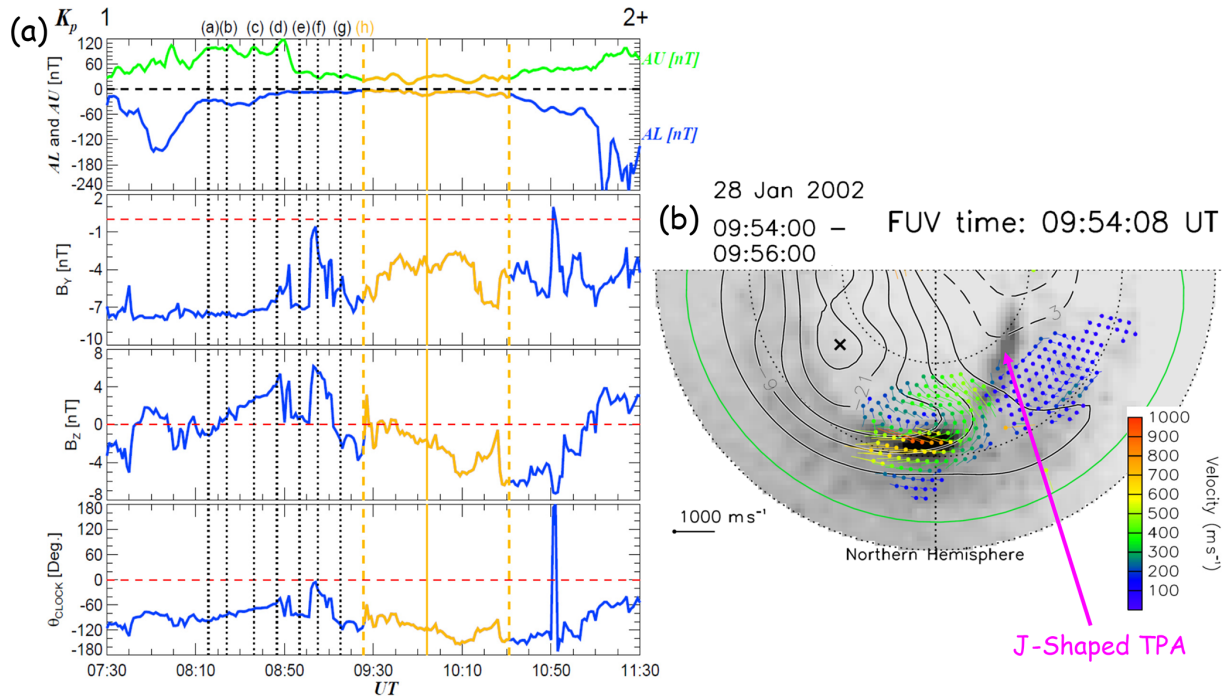


Figure 2. Panel (a) shows the plots of the geomagnetic activity indices and OMNI solar wind data, during 4 h between 07:30 and 11:30 UT on 28 January 2002, are displayed. The AL and AU indices, representing the geomagnetic activity at high latitudes, the y (dawn–dusk) and z (north–south) components of the IMF in geocentric solar magnetospheric (GSM) coordinates, and the IMF clock angle, as calculated with a formula of $\arctan(\text{IMF-}B_y / \text{IMF-}B_z)$. The K_p index to represent the average global geomagnetic condition during the interval of interest is also shown in the AL and AU plot panel. The TPA interval is bracketed by two gold broken lines and the solar wind condition, corresponding to 09:54:08 UT (b), is marked by a gold solid vertical line. The black dotted lines with the labels, from (a) to (h), in the top panel show the solar wind conditions corresponding to the times of the eight IMAGE FUV-WIC/SuperDARN plots in Fig. 3. Panel (b) shows the nightside ionospheric flow velocity vectors measured by SuperDARN, overlaid onto the snapshot of the IMAGE FUV-WIC data, including the aurora oval and nightside distorted (J-shaped) TPA seen at 09:54:08 UT in the Northern Hemisphere in geomagnetic coordinates. The concentric dotted circles show the magnetic latitude (MLat) at 60, 70, and 80°, respectively. The left, bottom, and right sides in each panel show 18, 24, and 6 h in magnetic local time (MLT), respectively. The green curves show the Heppner–Maynard boundary, which is the lower latitude limit for the ionospheric plasma convection pattern. The vector length and color code are assigned according to the flow orientation and intensity of ionospheric plasma velocity in units of meters per second (m s^{-1}).

formation scenario of the nightside distorted TPA, as proposed by Nowada et al. (2020). It still remains unclear, however, how the nightside distortion of the TPA was formed.

We consider that the TRINNI return flows themselves may be involved in the formation of the nightside end distortion of a TPA. In Fig. 4a, which shows an overlaid plot of the IMAGE FUV-WIC and SuperDARN radar data at 08:23:57 UT, we find no significant TRINNI return flows across the post-midnight sector of main auroral oval, suggesting that nightside reconnection was not occurring at this location, where the TPA subsequently formed. However, there were weak plasma flows which crossed the open/closed field line boundary, which is inferred to be collocated with the poleward edge of the auroral oval, in the pre-midnight sector. In Fig. 4b (09:15:12 UT), prior to the subsequent J-shaped TPA onset (09:25:27 UT; Fig. 4c), fast ionospheric TRINNI-type flows ($\sim 1000 \text{ m s}^{-1}$) were seen in the midnight sector. These flows explicitly support the idea that nightside reconnection was

occurring at this time. Furthermore, the flows were rotating their orientation toward the west (duskside) across the open/closed polar cap boundary, thus also providing evidence for the return of newly closed magnetic flux toward the dayside. At the J-shaped onset time in Fig. 4c, the TRINNI flows were still observed around the midnight sector, approximately 1 h MLT earlier than the TPA growth location. This indicates that nightside reconnection occurred across the midnight sector at the TPA onset time but not at the precise location of the TPA growth. The presence of the extremely slow flows (blue-colored vectors) around the growth point of the TPA, however, suggests that newly closed flux here is not returning but stagnating. This is consistent with the idea that any flux closed in this sector would contribute to form the TPA.

This transition from the equatorward flow across a stationary reconnection line (the regular pre-TPA case) to stagnant flows and a poleward-moving reconnection line (TPA growth

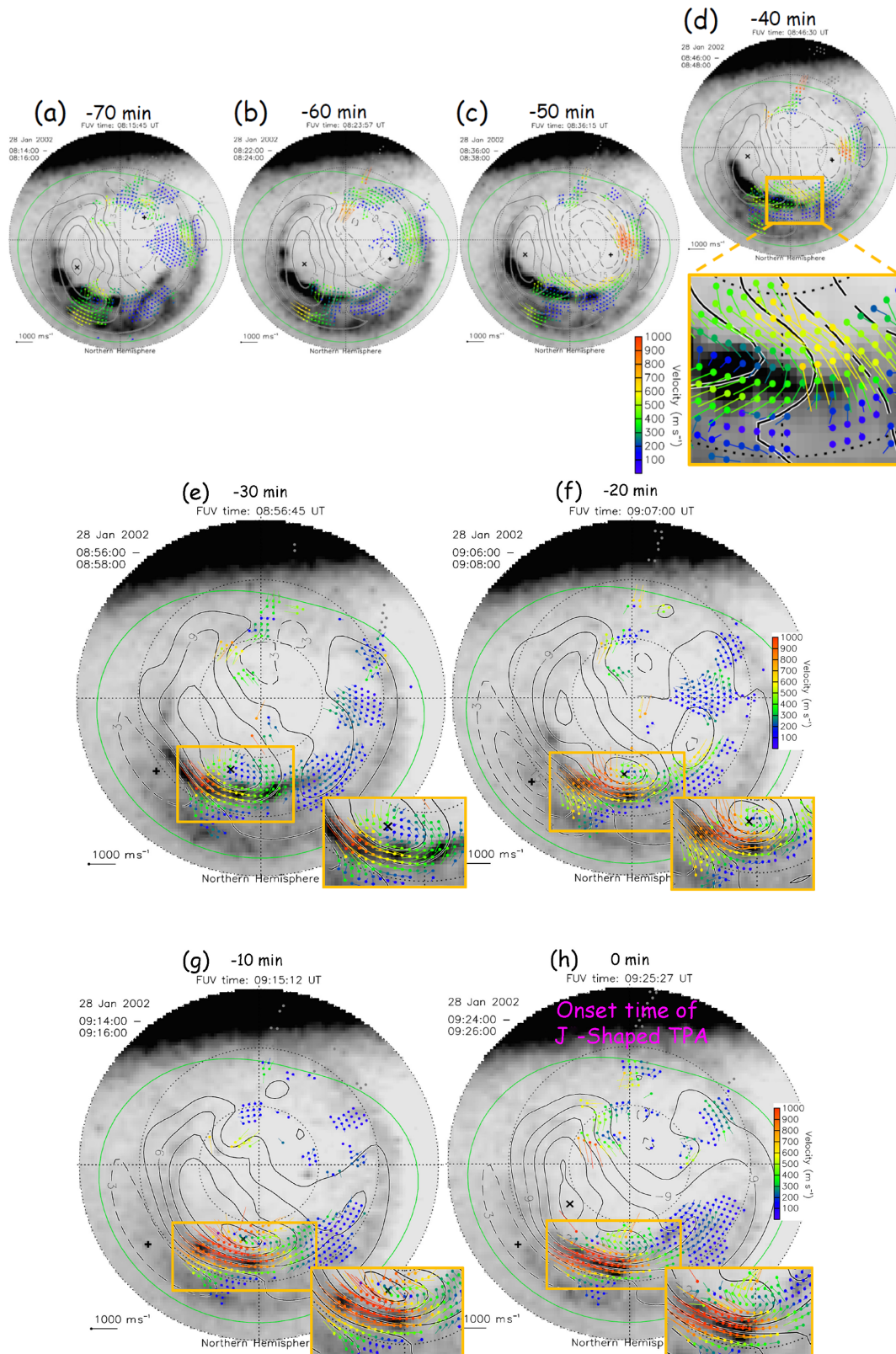


Figure 3. The eight selected overlaid plots of the SuperDARN radar and IMAGE FUV-WIC data are displayed. From panels (a) to (g), the overlaid plots, between 70 and 10 min prior to the onset time of J-shaped TPA (09:25:27 UT), as displayed in panel (h), are shown. Zoomed-up ionospheric TRINNI flow profiles in the regions including the midnight sector are shown within orange-framed boxes. The dayside auroral images and SuperDARN data are included, although the plot format in panels shown is basically the same as in Fig. 2b.

case; see Fig. 7 in Nowada et al., 2020) can be seen when comparing the flows around the location where the TPA protrudes into the polar cap between Fig. 4b and c. Of particular interest is that the equatorward plasma flow region has not ceased but has clearly moved from the TPA location (post-midnight) to being adjacent to the base of the TPA in the midnight sector. This implies that the stagnant flux of the TPA is azimuthally restricted and that the newly closed flux adjacent to the TPA is still contributing to a TRINNI return flow channel. Furthermore, plasma flows at the eastward end of the nightside distorted part of the TPA (i.e., the part adjacent to the main oval) also appear to be flowing out of the TPA and along the main auroral oval. This suggests that the field lines mapping to the equatorward end of the TPA, which corresponds to the closed lobe flux crossing the equatorial plane nearest to the Earth, have started to flow out of the TPA where they also contribute to the TRINNI-type return flow channel. We suggest that the nightside end of a TPA is distorted by these duskward plasma flows, which are flowing along the distorted nightside end of the TPA. The presence of two potential sources of return flux – the nightside end of the TPA and the TPA-adjacent flows – may explain the presence of the distinct auroral feature that forms the base of the J-part of the TPA adjacent to the main oval.

4.2 Global ionospheric plasma flows driven by dayside and nightside magnetic reconnection

Dayside magnetic reconnection globally drives the plasma flows within the polar cap (e.g., Dungey, 1961; Cowley and Lookwood, 1992). The TRINNI mechanism for generation of the flows at the distorted nightside end of the TPA requires an ongoing dayside reconnection, although Fear et al. (2015) suggest that TPAs are associated with a suppression of open flux production at the dayside. Here, we briefly consider the global convection during the TPA interval. At the J-shaped TPA onset time (Fig. 5a), the TPA just started to grow into the polar cap from the post-midnight auroral oval, which is marked with a yellow star. As discussed above, fast ionospheric TRINNI flows with a velocity of $\sim 1000 \text{ m s}^{-1}$ were seen adjacent to this, providing clear evidence of nightside reconnection occurrence. During the growth of the nightside distorted TPA (Fig. 5b), the ionospheric TRINNI return flows were still found in the vicinity of the poleward edge of the main auroral oval across the midnight sector (in particular, see the zoomed-up flow profiles in the orange-framed boxes). These flows suggest that magnetotail reconnection persisted even during the growth of the TPA and formed the closed field lines of the distorted part of the TPA. Further ionospheric flows along the distorted TPA nightside end were also observed.

Turning now to the dayside region, there were anti-sunward and anti-sunward/duskward plasma flows which entered the polar cap across the open/closed field boundary (poleward edge of the dayside main auroral oval), as high-

lighted by the cyan boxes in Fig. 5b. These flow signatures provide key evidence for the occurrence of the dayside reconnection (e.g., Cowley and Lookwood, 1992; Neudegg et al., 2000; Milan et al., 2000, and references therein). At the onset time of the J-shaped TPA (Fig. 5a), the dayside reconnection may have not yet been occurring because anti-sunward plasma flows across the dayside open/closed field line boundary were absent at this time, despite the IMF being oriented weakly southward (refer to interval (h) in Fig. 2a). However, the ionospheric flows did subsequently begin to enter the dayside polar cap, and continued to do so while the nightside distorted TPA was growing from the nightside main auroral oval at 09:47:59 UT, completely reaching the dayside oval by 10:28:58 UT, as shown in the two plots in Fig. 5b. It may be that ongoing dayside magnetic reconnection and subsequent excitation of the TRINNI flows are required elements of the mechanism by which the nightside distorted (J-shaped) TPAs are produced.

When considering global ionospheric convection patterns during an interval of dominant dawnward IMF- B_y , we expect duskward flows in the dayside polar cap and dawnward flows in the nightside polar cap (e.g., Cowley and Lookwood, 1992). In this case, with a TPA growing in the post-midnight sector, dawnward plasma flows will thus be expected on the duskside of the TPA. The SuperDARN radars detected some evidence of these flows, just poleward of $\sim 80^\circ$ in MLat, as indicated by the red boxes in Fig. 5b. Although the radar scatter within the polar cap is limited, the indicative flow pattern, as seen in the two plots in Fig. 5b, is consistent with Dungey cycle driving during a period of the dawnward IMF- B_y component under the southward IMF conditions in the Northern Hemisphere. The observed development of the J-shaped TPA from the nightside main auroral oval to the dayside was clearly not impeded by these Dungey-cycle-driven flows.

5 Discussions

5.1 Global ionospheric flow patterns associated with the J-shaped TPA growth

In this paper, we have tried to unravel the formation of the nightside distorted part of the J-shaped TPA using ionospheric flow observations by the SuperDARN radars. When the nightside distorted TPA was observed, the ionospheric TRINNI return flows were seen on the main auroral oval across the midnight sector, suggesting a formation mechanism associated with a nightside reconnection (e.g., Nowada et al., 2020). Interestingly, TRINNI flows were also observed in the region of the TPA growth, albeit at a reduced rate ($\sim 400 \text{ m s}^{-1}$) relative to the adjacent region ($700 \text{ m s}^{-1} \sim 900 \text{ m s}^{-1}$; see Fig. 4). This suggests that the flows at the TPA base did not fully stagnate, but that their reduced rate might explain the build-up of closed flux. Their presence might also explain the nightside distorted part of

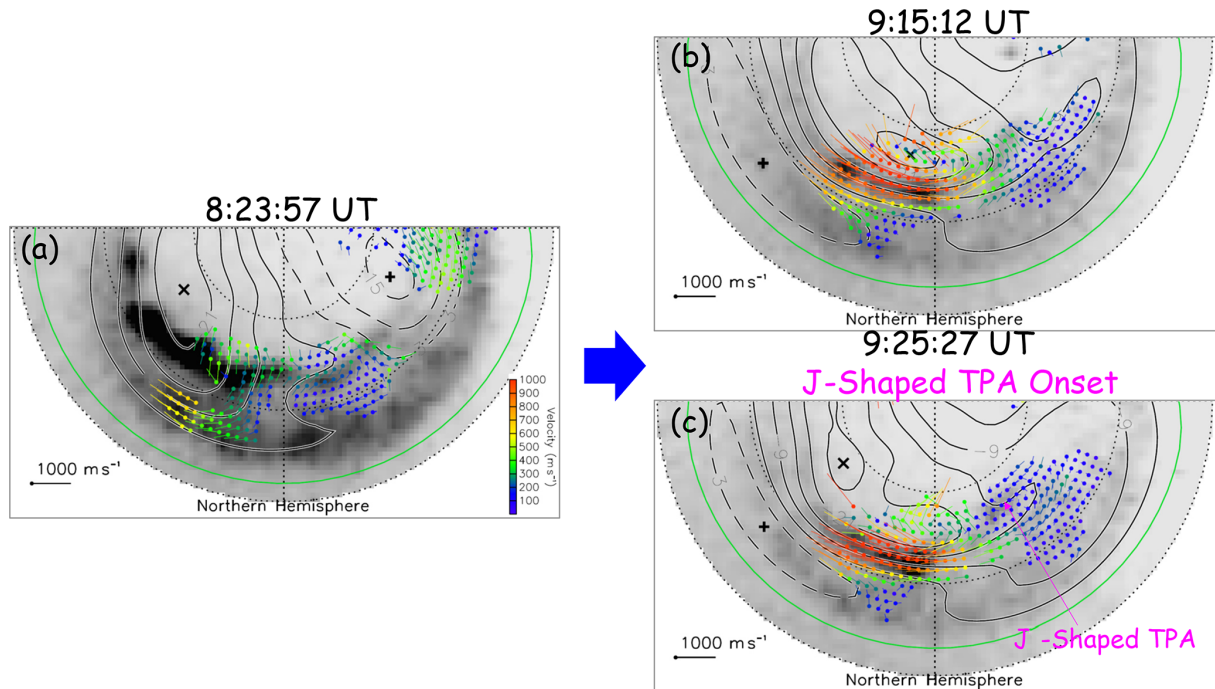


Figure 4. The three selected snapshots of the overlaid SuperDARN radar and IMAGE FUV-WIC data, to investigate the TRINNI flows and polar cap convection patterns, are shown. These overlaid plots are zoomed up to the nightside sector from 18 to 6 h in MLT. The overlaid plots, at about 1 h and at 10 min prior to the onset of the J-shaped TPA (09:25:27 UT), are shown in panels (a) and (b), respectively. Panel (c) displays the plot at the onset time of the J-shaped TPA. All figure formats are the same as those in Figs. 2b and 3.

the TPA, maintained by the newly closed flux being returned in the dusk convection cell by the TRINNI flow. This is consistent with the idea that the reduced-rate return flows are related to the nightside distortion of the TPA.

Considering the global ionospheric flow patterns which are estimated based on the SuperDARN radar observations and illustrated in Fig. 6, the dayside reconnection-driven flows enter the dayside polar cap in the dusk sector, highlighted with the thick cyan curved arrows and box, and are consistent with the plasma flows shown in cyan boxes in Fig. 5b. Furthermore, these flows subsequently feed the dawnward flows at higher latitudes, as highlighted with thick red curved arrows in the red box, which are being primarily driven by nightside magnetic reconnection. This is then also consistent with the post-midnight origin of the nightside reconnection flows close to the growth point of the TPA. However, it remains unclear whether or not the TRINNI originate at the same downtail location as the closed field lines of the J-shaped TPA.

Despite the persistence of Dungey cycle plasma flow patterns being driven by ongoing low-latitude dayside reconnection, which might be expected to inhibit the growth of a TPA, the J-shaped TPA observed here ultimately grew across the polar cap to the dayside. The J-shaped TPA's formation process is consistent with the model proposed by Nowada et al. (2020), that is, the closed field lines associated with the TPA are formed by nightside magnetic reconnection,

which is demonstrated by the presence of TRINNI return flows (thick black curve). Furthermore, the reconnection points should retreat to further downtail as the nightside distorted TPA grows, as shown with green stars in the duskside (surrounded by a blue box). Nowada et al. (2020) showed that reconnection-associated upward field-aligned currents (FACs), plausibly triggered by the nightside reconnection as shown with purple solid arrows, can be a source of the J-shaped TPA (magenta shading). In this case, however, conclusive signatures on the existence of upward FACs around the TPA cannot be shown because, unfortunately, the magnetic field data obtained by low-altitude orbiters were absent during the interval of interest. Instead, we are able to show some indication of the TPA-associated FAC flowing sense, which is provided by equivalent ionospheric current (EIC), which is estimated based on the geomagnetic field data from the SuperMAG ground observatory network (Gjerloev, 2012).

In Fig. 7, EIC distributions at 09:26 (Fig. 7a), 10:00 (Fig. 7b), and 10:10 UT (Fig. 7c), projected onto the IMAGE FUV-WIC imager data in geomagnetic coordinates, are shown to estimate the orientation and scale of the FAC system around the growing J-shaped TPA. The EIC vectors (red bars) are derived by rotating the horizontal magnetic field components (local magnetic north–south and east–west components) 90° clockwise, using the same calculation techniques proposed by Glassmeier et al. (1989), Moretto et

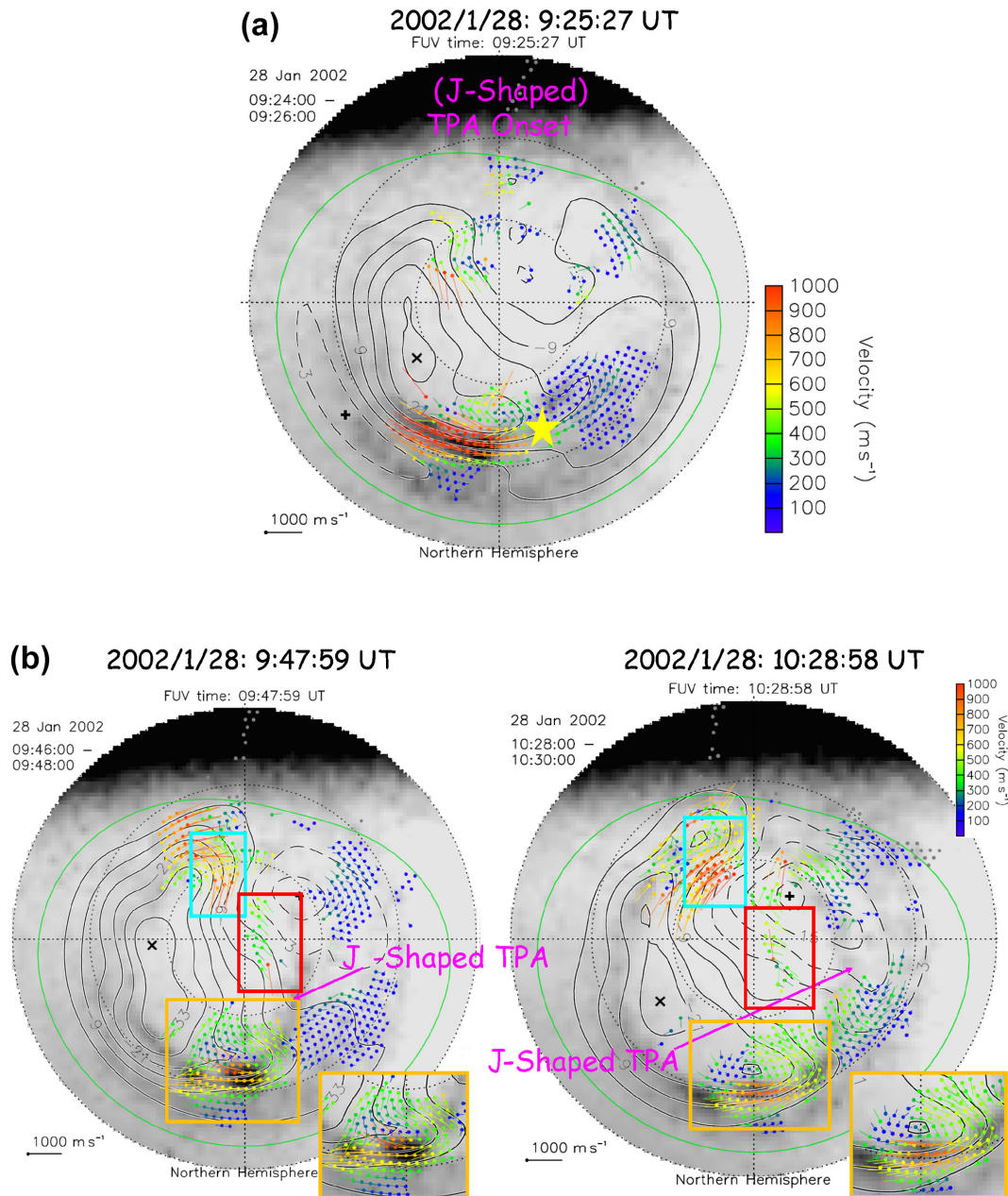


Figure 5. (a) An overlaid plot of the ionospheric plasma flow pattern and IMAGE FUV-WIC imager data at the onset time (09:25:27 UT) of the J-shaped TPA is shown. The TPA growth point toward the dayside is marked with a yellow star. (b) The two panels show the overlaid plots at 09:47:59 and 10:28:58 UT, when the fast flows entered the dayside polar cap across the poleward edge of the dayside main auroral oval, which is a proxy of the open/closed field line boundary, during the TPA growth, as highlighted by the cyan boxes. The dawnward ionospheric flows, which are consistent with the flow patterns, as expected from Dungey cycle driving during a period of downward IMF- B_y component under the southward IMF (IMF- $B_z < 0$), are surrounded by red boxes. The profiles of TRINNI-type westward fast flows around the midnight sector (highlighted with by orange boxes) are shown within the orange-framed box. All figure formats are the same as Fig. 3.

al. (1997), Motoba et al. (2003), and references therein. The major trends of the EIC vectors in close proximity to the regions of growth of the J-shaped TPA (magenta thick and curved arrows) exhibit a significant counterclockwise rotation, implying that upward FACs might be generated around the J-shaped TPA. Significant distortion-aligned duskward

plasma flows from around the J-shaped TPA growth point (orange curved arrow in Fig. 6) can be explained within a framework of Dungey-cycle-driven plasma flow patterns in the polar region. The nightside ionospheric flow patterns during the J-shaped TPA, as seen in Figs. 3, 4, and 5, clearly have dawn–dusk asymmetry, suggesting that the IMF- B_y com-

Table 1. The event number of the nightside distorted TPA, which was categorized by three types of the IMF- B_z polarity (northward IMF – IMF- $B_z > 0$; southward IMF – IMF- $B_z < 0$; turning from northward to southward IMF – from IMF- $B_z > 0$ to IMF- $B_z < 0$) is shown. This table is made based on the 17 nightside distorted TPA events which were selected from the IMAGE FUV-WIC observations from 2000–2005 (see Table S1).

	IMF- $B_z > 0$	IMF- $B_z < 0$	From IMF- $B_z > 0$ to IMF- $B_z < 0$
J-shaped TPA	8	2	1
L-shaped TPA	5	0	1
Total	13	2	2

ponent influenced the nightside magnetosphere; that is, the nightside plasma sheet deformation and magnetic field line twisting in the magnetotail were caused by the IMF- B_y penetration (e.g., Cowley, 1981, 1994; Milan et al., 2005; Fear and Milan, 2012a). Nowada et al. (2020) demonstrated that the nightside distorted TPAs also grow to the dayside under dawn–dusk asymmetric ionospheric flow patterns. In this study, we are able to reveal the details of the ionospheric flow patterns that cause the nightside distorted part of the TPA using ground-based radar observations.

The scenario for the formation of nightside distortion of a TPA and possible whole TPA growth model, as illustrated in Fig. 6 and Nowada et al. (2020), can be applied to J- and L-shaped TPAs formed during northward IMF intervals. Table 1 shows the event number of the nightside distorted TPA, categorized by three types of the IMF- B_z polarity based on the 17 nightside distorted TPA events, which were selected from the IMAGE FUV-WIC observations from 2000–2005 and include the nine TPA events used in Nowada et al. (2020; see Table S1). In most TPA cases during the northward IMF intervals, TRINNI flow signatures were detected with the SuperDARN radar arrays, suggesting that magnetotail reconnection occurrences and the explanations with reconnection-based nightside distorted TPA formation process can be expected. On the contrary, the nightside distorted TPA under purely southward conditions is only two events. In the case discussed in this study, we show the TPA nightside distortion formation and explain a possible whole J-shaped TPA growth, adopting a nightside distorted TPA formation model under northward IMF conditions (Nowada et al., 2020), while neither satellite nor SuperDARN radar profiles can be obtained in another case. Therefore, it is required to discuss enough space- and ground-observation data when considering general nightside distorted TPA formation processes under southward IMF conditions.

5.2 Significant differences between nightside distorted TPA, bending arcs, and double oval auroral forms

Given the variety of different TPA morphologies discussed previously in the literature, we here briefly outline the main

differences between the distorted arc discussed in this paper, and bending arcs and double oval auroral forms.

5.2.1 Difference from bending arcs

The most significant difference between nightside distorted TPAs and bending arcs is whether or not nightside magnetic reconnection is closely related to its formation (e.g., Kullen et al., 2015). During our J-shaped TPA interval, TRINNI return flows were clearly observed in the poleward edge of the nightside main auroral oval (near the growth point of the J-shaped TPA), suggesting that nightside reconnection persisted even during the TPA development and, according to our proposed mechanism, plays an essential role in the J-shaped TPA formation. Bending arcs, on the other hand, tend to occur in association with dayside magnetic reconnection (Kullen et al., 2015) and, thus, have a different formation process to the nightside distorted TPA presented here. In this study, the J-shaped TPA was observed during a southward IMF interval, and ionospheric plasma flow patterns indicating the occurrence of low-latitude dayside magnetic reconnection were also observed. These IMF conditions and associated ionospheric flow profiles are actually consistent with those found when bending arcs are likely to be formed (Kullen et al., 2015; Carter et al., 2015). However, in this case, even in the presence of dayside magnetic reconnection, which drove Dungey-cycle-driven plasma flow patterns, this TPA displayed no characteristics of a bending arc, with its growth toward the dayside being rather straightforward. Because bending arcs develop toward the dawn or dusk sector in a region more poleward of the dayside magnetic reconnection line (merging gap; Carter et al., 2015), the development profile of the J-shaped TPA is also different from that of bending arc. Therefore, we can be certain that the nightside distorted TPA discussed in our study is quite distinct from bending arcs.

5.2.2 Difference from double oval auroral forms

In association with our J-shaped TPA westward (duskside) flow, both equatorward and poleward regions of the TPA nightside distortion were observed in the SuperDARN data, as seen in the plasma flow patterns after the TPA onset (several panels of Figs. 2, 3, 4, and 5). These suggest that upward currents (FACs) should flow in the regions equatorward and poleward of the nightside distortion of the TPA, although a detailed FAC profile cannot be estimated based on the EIC vectors poleward of the TPA nightside distortions because of an absence of sufficient geomagnetic field data (see Fig. 7). In the region just poleward of the main auroral oval, including the distorted part of TPA (Fig. 7a, b), we can suggest the presence of upward FACs around the distortion at the TPA nightside end. These FAC profiles around the nightside distortion of the TPA are inconsistent with the double auroral oval structures elucidated by Ohtani et al. (2012; see their

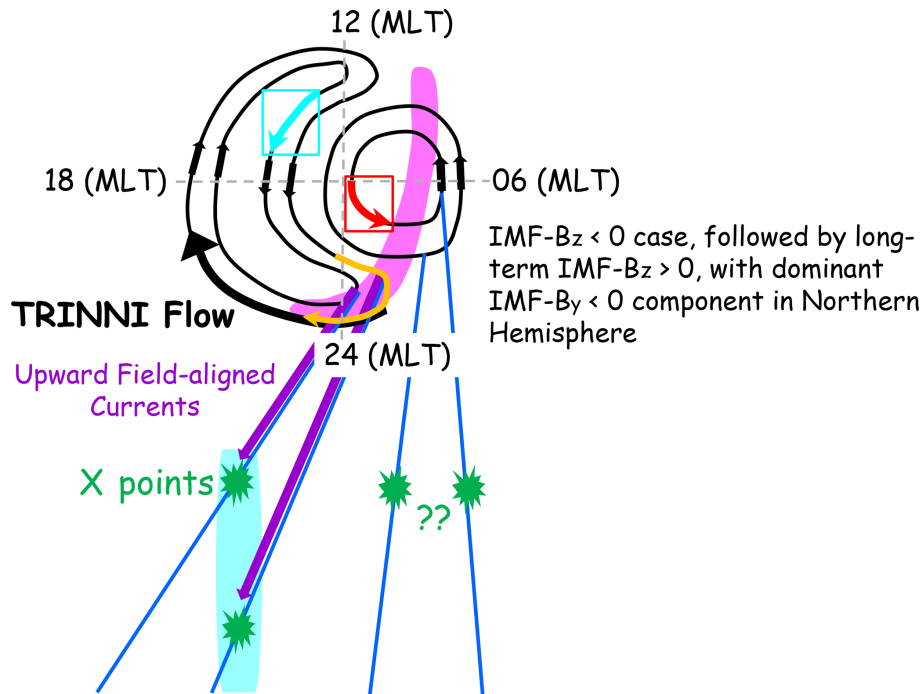


Figure 6. A schematic illustration of the global ionospheric convection flow patterns in the Northern Hemisphere driven by dayside and nightside magnetic reconnections during the growth interval of the J-shaped TPA (magenta) is shown. This illustration is the modified Fig. 3 of Grocott et al. (2005). These ionospheric flow patterns are expected to be seen when the IMF- B_z are southward with a dominant dawnward component. The ionospheric flows, highlighted by thick cyan and red curved arrows and surrounded with cyan and red squares, are corresponding to those highlighted with the same colored squares, as shown in Fig. 5b. The TRINNI return flows and plasma flows out of the TPA growth point, which may lead to the formations of the observed J-shaped TPA and its nightside end distortion, are shown with black thick and orange curved arrows, respectively. The reconnection-generated upward field-aligned currents (FACs), as a source of nightside distorted TPA, are shown with purple solid arrows. The blue solid lines indicate magnetospheric closed field lines. The reconnection points, retreating as the TPA grows to the dayside, are shown with green stars, and their retreat line is highlighted with a blue box. The day–night and dawn–dusk meridian lines are shown with gray broken lines.

Fig. 11). Therefore, the nightside distortion of the J-shaped TPA is also independent of the double auroral oval.

6 Conclusions

Nightside magnetic reconnection and associated FACs are integral processes in the formation of nightside distorted TPAs, such as the J-shaped TPA presented in this study. In particular, a migration of the equatorward plasma flows, which rotated to align with the main auroral oval at the point where the TPA starts to protrude into the polar cap toward the dayside, plays a significant role in the formation of the distorted nightside end of the TPA. These plasma flows may be interpreted within a framework of the dawn–dusk asymmetric polar cap plasma flow patterns, produced by ongoing Dungey cycle activity in the presence of a dominant IMF- B_y component. From the global ionospheric plasma flow patterns determined from SuperDARN radar observations, we can conclude that nightside distorted TPAs are formed by a juxtaposition of localized flow stagnation (as required for regular TPAs) in the presence of ongoing TRINNI-type tail-reconnection-driven

flows consistent with the distortion of the TPA nightside end. It may also be the case that this process is facilitated by a southward IMF, and an associated ongoing dayside reconnection, that is required to feed the TRINNI flows. As such, nightside distorted TPAs, including the J-shaped TPA in this study, are quite different from both double oval and bending arcs in terms of their formation process.

In the near future, the SMILE (Solar wind Magnetosphere Ionosphere Link Explorer) and STORM (Solar–Terrestrial Observer for the Response of the Magnetosphere) satellites, which include auroral ultraviolet imager (UVI) data with higher spatial and temporal resolutions than those on the Polar and IMAGE missions, will be launched. If the UVI auroral imager data can be safely acquired after a successful launch of these new satellites, then we can expect to collect more nightside distorted TPA events and to study the detailed features and formation mechanism of these J- and L-shaped TPAs more closely.

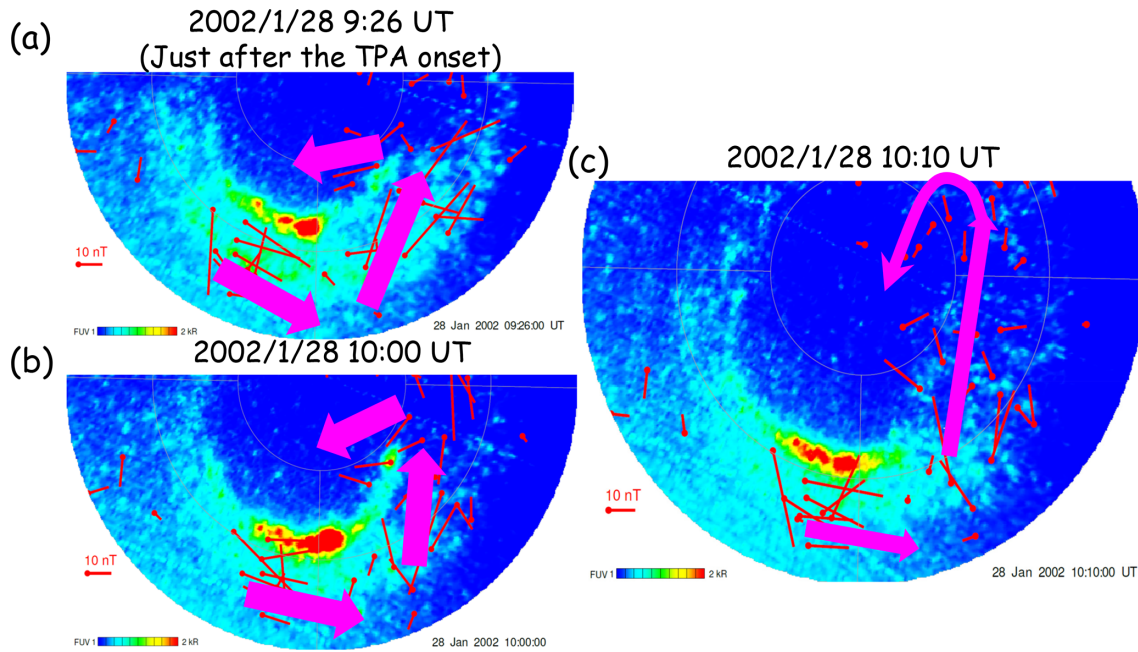


Figure 7. Equivalent ionospheric current (EIC) distributions at 09:26 (a), 10:00 (b), and 10:10 UT (c), projected onto the IMAGE FUV-WIC data in geomagnetic coordinates, are shown. The EIC vectors (red bars) are derived by rotating the horizontal magnetic field components (local magnetic north–south and east–west components) 90° clockwise. The geomagnetic field was measured at the ground magnetic observatories from the SuperMAG ground observatory network (Gjerloev, 2012). Each panel is oriented such that the right, bottom, and left sides are corresponding to dawn (6 h), midnight (24 h), and dusk (18 h) in MLT, respectively. The major directional trends of the EIC vectors in close proximity to the regions of growth of the J-shaped TPA are shown with magenta thick and curved arrows. The white circles show the MLat values from 60 to 80° as they go inward. The color codes are assigned according to unit of Rayleigh.

Code availability. The software codes to visualize the SuperDARN radar data in this study were downloaded and used from <https://doi.org/10.5281/zenodo.6473603> (SuperDARN Data Analysis Working Group et al., 2022).

Data availability. All SuperDARN radar data are processed by the software of `fitacf v1.2` and `make_grid v1.14.er` and can be obtained from <https://www.bas.ac.uk/project/superdarn> (British Antarctic Survey Super Dual Auroral Radar Network team, 2019). IMAGE FUV-WIC data were accessed from <http://image.gsfc.nasa.gov> (Imager for magnetopause-aurora global exploration science center team, 2019). Solar wind OMNI (ACE MFI and SWE) data were obtained from the Coordinated Data Analysis Web (https://cdaweb.gsfc.nasa.gov/cdaweb/istp_public/, NASA Goddard Space Flight Center Coordinated Data Analysis Web team, 2019) provided by GSFC/NASA. The ground magnetometer data were obtained from the website of the SuperMAG ground observatory network (<https://supermag.jhuapl.edu/>, Jesper W Gjerloev and SuperMAG team, 2020).

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