



# Rocket in situ observation of equatorial plasma irregularities in the region between E and F layers over Brazil

Siomel Savio Odriozola<sup>1,2</sup>, Francisco Carlos de Meneses Jr.<sup>1,3</sup>, Polinaya Muralikrishna<sup>1</sup>, Alexandre Alvares Pimenta<sup>1</sup>, and Esfhan Alam Kherani<sup>1</sup>

<sup>1</sup>National Institute for Space Research – INPE, Av. dos Astronautas, 1758, Jd. Granja-CEP: 12227-010 SJC, SP, Brazil

<sup>2</sup>Department of Space Geophysics, Institute of Geophysics and Astronomy – IGA, Calle 212, 2906, La Coronela, La Lisa, Havana, Cuba

<sup>3</sup>State Key Laboratory of Space Weather, National Space Science Center NSSC, Chinese Academy of Sciences (CAS), No. 1 Nanertiao, Zhongguancun, Haidian district, 100190, Beijing, China

*Correspondence to:* Siomel Savio Odriozola (siomel.savio@inpe.br)

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**Abstract.** A two-stage VS-30 Orion rocket was launched from the equatorial rocket launching station in Alcântara, Brazil, on 8 December 2012 soon after sunset (19:00LT), carrying a Langmuir probe operating alternately in swept and constant bias modes. At the time of launch, ground equipment operated at equatorial stations showed rapid rise in the base of the F layer, indicating the pre-reversal enhancement of the F region vertical drift and creating ionospheric conditions favorable for the generation of plasma bubbles. Vertical profiles of electron density estimated from Langmuir probe data showed wave patterns and small- and medium-scale plasma irregularities in the valley region (100–300 km) during the rocket upleg and downleg. These irregularities resemble those detected by the very high frequency (VHF) radar installed at Jicamarca and so-called equatorial quasi-periodic echoes. We present evidence suggesting that these observations could be the first detection of this type of irregularity made by instruments onboard a rocket.

**Keywords.** Ionosphere (equatorial ionosphere; plasma temperature and density; plasma waves and instabilities)

## 1 Introduction

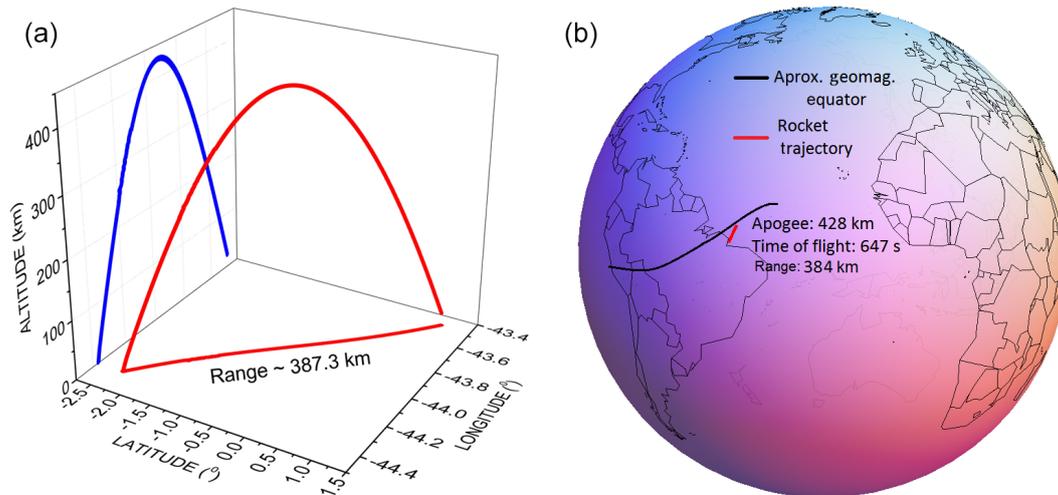
Many ionospheric studies using rocket-borne probes have focused their attention on nighttime equatorial plasma irregularities. A significant part of these studies has been devoted to the analysis of irregularities located at typical F region heights (e.g., Rino et al., 1981; LaBelle et al., 1986; Jahn and

LaBelle, 1998; Muralikrishna et al., 2004). Studies using in situ measurements of equatorial nighttime E region irregularities, however, are quite a few, as noted by Sinha et al. (2011). The lack of references is most noticeable when it comes to the intermediate region between the E and F layers ( $\sim 130$ – $300$  km).

Although data from other studies show irregular structures in the electron concentration in the F valley region (Abdu et al., 1991; Vickrey et al., 1984; Sinha et al., 1999; Raizada and Sinha, 2000; Hysell et al., 2005; Muralikrishna and Abdu, 2006), up to now, only Vickrey et al. (1984) and Raizada and Sinha (2000) have analyzed the irregularities detected in this region. Both studies used measurements obtained using a Langmuir probe onboard rockets that flew in low-latitude trajectories.

Plasma irregularities located at typical E and F valley heights have been studied before using VHF radar observations. Woodman and Chau (2001) were the first to characterize irregular structures in E and lower F regions using observations from the Jicamarca Radio Observatory (JRO). They used the term “Equatorial QuasiPeriodic” or EQP to label these irregular echoes.

More recently, Chau and Hysell (2004) reported large-scale plasma waves at high altitude in the E region over Jicamarca and around twilight. Using an aperture synthesis imaging technique, the authors were able to show a striking and complex set of striations evidencing different levels of signal-to-noise ratio (SNR), representing a range of different electron density regions.



**Figure 1.** Rocket trajectory (a) and flight details (b). The blue parabola in (b) is the projection of the trajectory onto a vertical plane.

In the equatorial region over Brazil, a statistical study with observations made during 581 nights (from 2001 to 2008) was conducted by Kherani et al. (2012) using a low-power coherent scatter radar located at the São Luís equatorial station. Similar to the results of Woodman and Chau (2001), the structures were observed during occurrences of an ascending irregular bottom-type or bottom-side F layer and at the same time or prior to an ascending and well-developed plume. Furthermore, and unlike Jicamarca, the research showed that the frequency of occurrence of this type of irregularity was extremely low and the few observations of such events were concentrated in summer months (Southern Hemisphere) and years of high solar activity.

In this work, the vertical distribution of electron density was obtained using a Langmuir probe as a payload of a two-stage rocket, which flew toward the magnetic equator from the rocket launching station in Alcântara, Brazil. Both the upleg and downleg profiles, showed small- and medium-scale plasma irregularities in the intermediate region between the E and F layers. Given the characteristics of these profiles, in our results we propose that these irregularities are the EQP irregularities only detected, so far, by VHF coherent scatter radars. In the next section, the rocket flight and payload details, ionosonde observations and Langmuir probe data will be presented. Later, our observations will be compared with a range time intensity (RTI) map previously obtained at Jicamarca Radio Observatory. We will conclude by summing up our results and considerations.

## 2 Observations

### 2.1 Rocket flight

As part of the Iguaiaba operation, a two-stage VS-30 Orion rocket was launched from the equatorial rocket launching

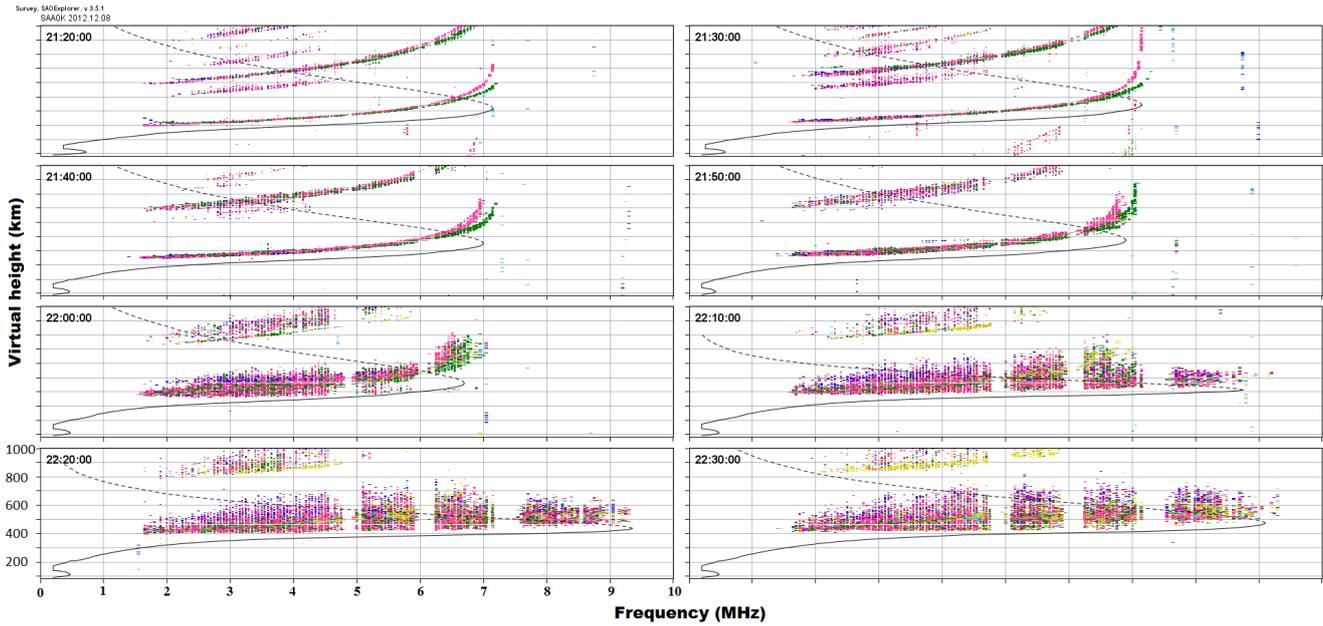
station in Alcântara (CLA) located at 2.4° S, 44.4° W (dip latitude 5.5° S), Brazil, on 8 December 2012 (LT = UT – 3). The launch time (19:00 LT) was carefully selected in order to intercept irregular structures that might arise soon after sunset. Flight parameters and the rocket trajectory are shown in Fig. 1.

As can be seen, the trajectory was toward the magnetic equator in the northeast. The horizontal range was  $\sim 384$  km, and the apogee of the rocket reached typical F region altitudes ( $\sim 428$  km). The flight was tracked by three radars over 647 s. The rocket carried three instruments: a two-axis magnetometer, a GPS receiver prototype and a conical Langmuir probe (CLP). Data from a GPS receiver confirm the vehicle's trajectory provided by the radars. From magnetometer data, it was possible to estimate the rotation ( $\sim 5$  rps) and precession ( $\sim 0.09$  rps) of the rocket along the flight.

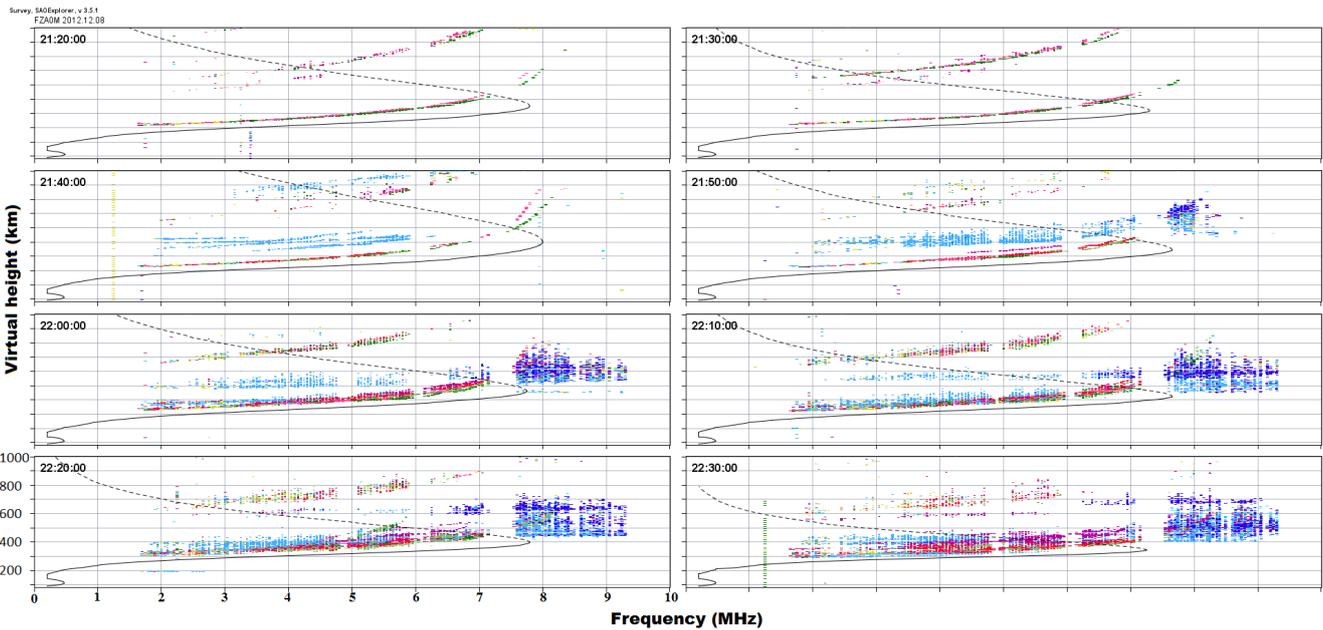
### 2.2 Digisonde observations

Ground stations located at São Luís (2.6° S, 44.2° W, dip latitude 6° S) and Fortaleza (3.9° S, 38.4° W, dip latitude 15° S) regularly probe the ionosphere using digisondes at the same time of rocket launch. São Luís station is located only 35 km south of CLA; therefore, it was used to monitor the ideal conditions (pre-launch rocket) that precede the bubble's onset, while Fortaleza station is positioned  $\sim 684$  km southeast of the launch center. Evidence of the presence of plasma irregularities was observed in the ionograms of both stations.

A sequence of eight ionograms taken at São Luís around the time of launching are shown in Fig. 2. Until 19:00 LT, ionograms over this station do not exhibit any spread echoes. Minimum virtual height of the F layer ( $h'_{\min}F$ ), however, increases suddenly until a first manifestation of spreading in the echo received by the digisondes becomes visible around 19:00 LT, indicating a fast upward movement of the F layer



**Figure 2.** Sequence of ionograms taken on 8 December 2012 (18:20–19:30 LT) at São Luís. Spread F begins around 19:00 LT.



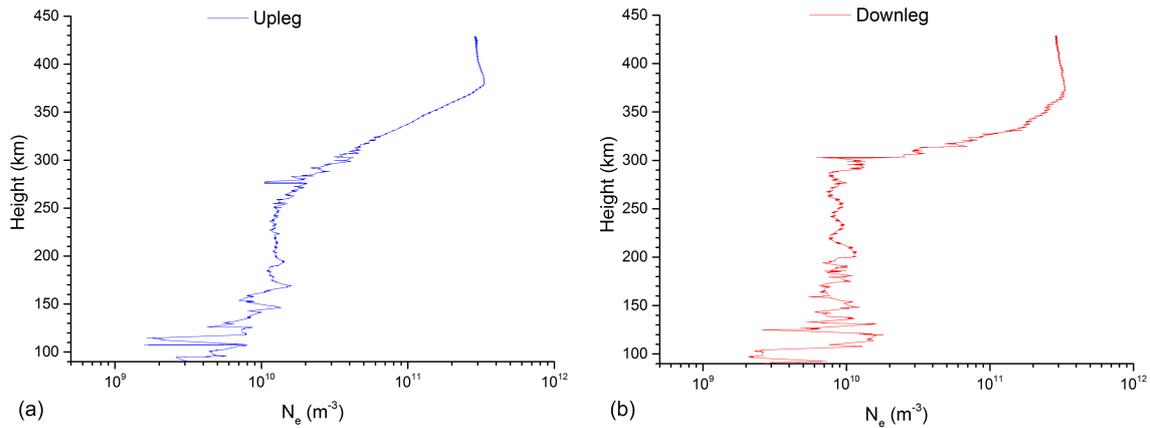
**Figure 3.** Sequence of ionograms taken on 8 December 2012 (18:20–19:30 LT) at Fortaleza. Spread F begins around 19:30 LT.

prior to the sunset. The spread F condition is evident in the ionograms for the rest of the night.

Figure 3 shows the ionograms obtained at Fortaleza simultaneously with those appearing in Fig. 2. In this case the manifestation of a spread echo first appeared in a signal coming from the north-northeast (echoes in light blue color). Signals vertical above Fortaleza station show a range spread F at 19:30 LT or 30 min after the appearance of the same condition at São Luís and the rocket’s launch. The rise of  $h'_{\min} F$

over Fortaleza is not as pronounced as that observed in São Luís.

From ionograms of both stations, it can be said that equatorial and low-latitude irregularities developed at the time of launch and after the rocket’s flight.



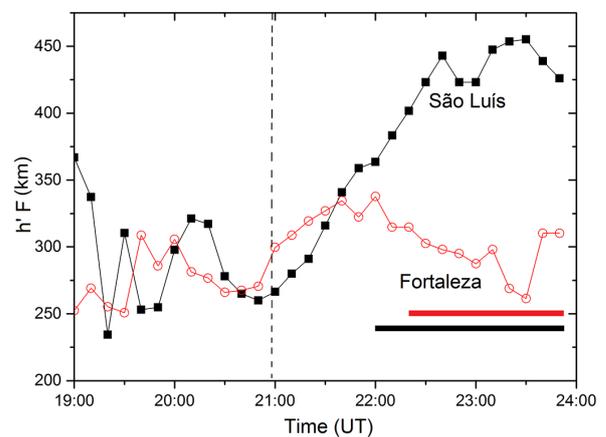
**Figure 4.** Electron density profiles obtained with a Langmuir probe data onboard a VS-30 Orion rocket on 8 December 2012, 19:00 LT.

### 2.3 Langmuir probe measurements

The Langmuir probe used in this experiment was mounted on the nose tip of the rocket and was able to operate alternately in swept and constant bias modes. During the swept mode, the bias applied to the probe was linearly increased from  $-1$  to  $2.5$  V in  $1.5$  s and then maintained (constant bias mode) at  $2.5$  V for  $1$  s. The current collected by the probe during the constant bias mode is directly proportional to the local electron density ( $N_e$ ) and can be used to estimate the  $N_e$  height profile. This current is filtered onboard in two bands: the low-pass filter, which limits the sample rate to  $\sim 250$  Hz (referred to as DC), and the high-pass filter, which allows sample rates as high as  $\sim 2.2$  kHz (referred to as AC). From the AC signal, the probe could measure electron density irregularities in a wide range of scale sizes, depending on the velocity of rocket. For instance, at typical rocket velocities, the probe can detect fluctuations of up to  $1$  m at  $150$  km of altitude as well as up to  $0.6$  m at  $316$  km. The  $N_e$  height profiles shown in Fig. 4 were obtained from the DC signal.

Between  $100$  and  $300$  km of altitude, several irregularities are visible in both the upleg and downleg. This type of irregularity is rarely found or discussed in the literature using in situ measurements (e.g., Hysell et al., 2005; Vickrey et al., 1984; Raizada and Sinha, 2000). During previous rocket experiments over the Brazilian territory, only two experiments carried out in Natal ( $5.8^\circ$  S,  $35.2^\circ$  W, dip latitude  $12^\circ$ ) showed a similar electron density vertical profile: one was a Javelin sounding rocket launched by NASA on 18 November 1973 at 21:22 UT (Vickrey et al., 1984); the second was a Sonda III rocket launched on 11 December 1985 at 23:30 UT (Abdu et al., 1991). It is important to note that the trajectories of these rockets were off the geomagnetic equator by more than  $8^\circ$ .

Above  $320$  km the profiles do not show any depletion that can be identified as a plasma bubble, even though ionograms indicate the presence of large-scale irregularities. The most interesting features appear between  $200$  and  $300$  km during

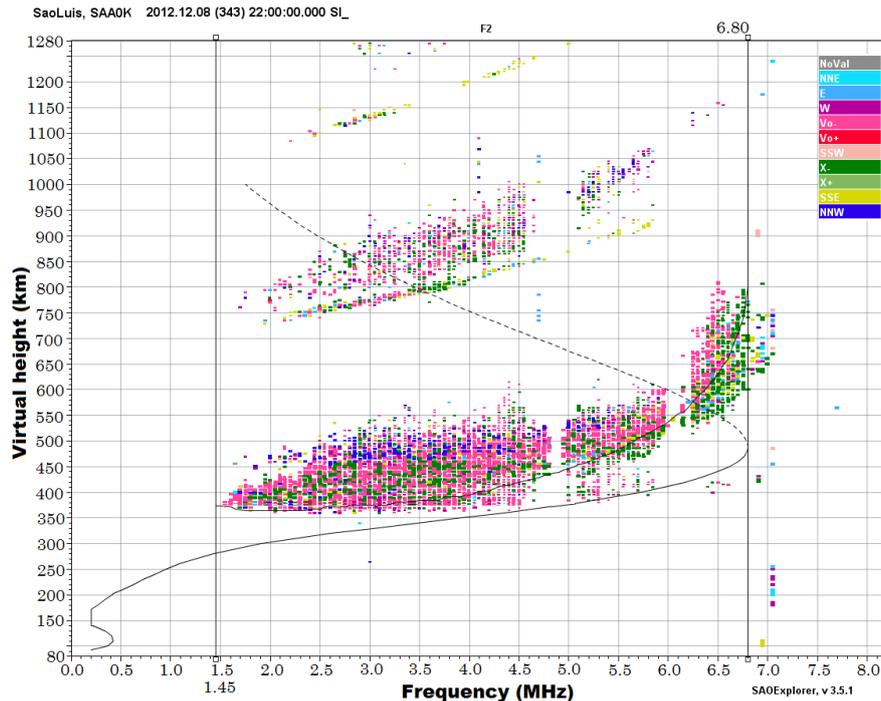


**Figure 5.** Upward and subsequent downward movement of the F layer observed at São Luís and Fortaleza during the flight of the rocket on 8 December 2012. The local sunset at São Luís represented by the vertical dashed line. The local sunset at Fortaleza occurred  $\sim 24$  min before São Luís. The red and black bars represent the range of time during which a spread F condition was observed in the ionograms taken at the Fortaleza and São Luís stations, respectively.

the downleg (Fig. 4b). The  $N_e$  height profiles exhibit a wave vertical pattern with an apparent vertical wavelength of  $20$ – $28$  km. The altitude in which these irregularities appeared and the quasi-periodic nature of the irregularities in a height range typical of the valley region were the key features that motivated the comparative study presented in the next section.

### 3 Results and discussions

It is now well established that vertical rise in the F layer base, as observed by ionosondes, is associated with a pre-reversal enhancement (PRE) in the vertical plasma drift. This rapid increase in  $h'_{\min}F$  may also be an effect of the rapid rise of



**Figure 6.** Ionogram obtained at São Luís at 19:00 LT. The F peak height of the  $N_e$  vertical profile (black dashed line) was 474 km.

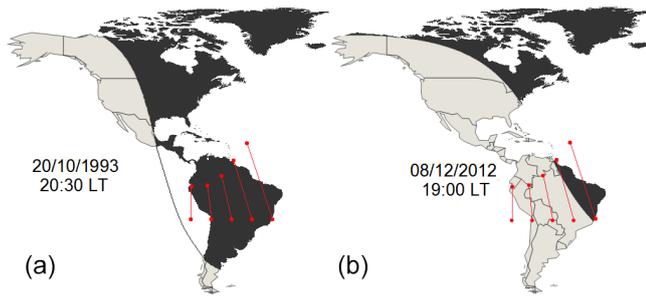
the solar terminator close to sunset hours. However, Bittencourt and Abdu (1981) showed that when the F layer height is above a threshold of 300 km, the apparent F layer vertical displacement velocity, inferred from ionosonde measurements, is the same as the vertical plasma drift velocity. The variation of the virtual height for both stations, depicted in Fig. 5, shows a rapid upward movement of the F layer at São Luís starting around 18:50 LT (21:50 UT) and during the local sunset. This movement continued well beyond the time of the launch and the  $h'_{\min}F$  reach a value of near 450 km. In Fortaleza, the  $h'_{\min}F$  was neither very long nor abrupt. This is expected due to the separation from the magnetic equator of the two locations. Over São Luís, the geomagnetic field lines are more horizontal than over Fortaleza, and PRE in the vertical drift is more pronounced over locations with lower dip latitudes.

We can affirm that large-scale irregularities were settled at F region heights over Alcântara by the time of launch based on two facts: firstly, the PRE is related to the variation in conductivity along geomagnetic field lines between conjugate regions and this has a direct relation with the development of equatorial irregularities. Secondly, from Figs. 2 and 3 we see that the spread F conditions first appeared over São Luís and then over Fortaleza, recalling that there is  $\sim 24$  min time difference between the two locations.

In Fig. 6 we show a single ionogram over São Luís taken at the same time of the launch. The rocket could not reach the F peak at  $\sim 480$  km, and therefore the plasma depletion eventually found above the F layer base was not detected by

the probe. Another aspect to be mentioned is that the irregular structures detected by the probe and below 300 km do not appear in the ionograms. As noted by Woodman (2009), valley-type irregularities do not produce significant spread traces in ionograms. Additionally, after a careful inspection of ionograms in Fig. 2, it can be noted that the sounding frequencies start at 1.5 MHz in all of the cases. According to the  $N_e$  height profile, represented in a black curve over the echoes in Fig. 6, the true height that corresponds with this minimum sounding frequency is around 280 km. The majority of irregularities detected by the Langmuir probe were below this altitude and could have produced echoes at frequencies below 1.5 MHz.

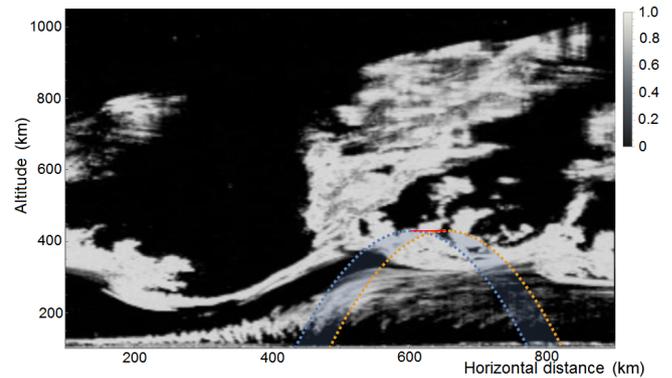
The quasi-periodic irregularities observed during the downleg part of the flight led us to compare our in situ measurements with the other and only type of data showing such irregularities: an RTI map elaborated from VHF radar observations. Unfortunately, the low-power VHF radar located very near CLA and used in other studies (e.g., de Paula and Hysell, 2004; de Paula et al., 2011; Kherani et al., 2012) was inoperative between October 2012 and February 2013; thus, it was decided to use the RTI map found in Woodman and Chau (2001) to gain insight about our in situ measurements. This RTI map obtained on 20 October 1993 over Jicamarca is a great example of E and F valley region irregularities and was good enough for our purposes: to make a comparison of the spatial distribution of irregularities as observed by our in situ measurements and the VHF radar in Jicamarca.



**Figure 7.** Solar terminator and geomagnetic field lines projected onto the American sector map corresponding to the date and time of the observations made in Jicamarca (a) and Alcântara (b).

Before proceeding with the description of this comparison, it is important to discuss the differences and similarities between the locations from which the data compared in this work were obtained. Both the Alcântara launch station and Jicamarca are located very close to the geomagnetic equator, differing by magnetic declination and the geomagnetic secular variation, which are greater over Alcântara. This difference in the angle of magnetic declination ultimately controls the seasonal occurrence of the conditions related with plasma bubble onset over the geomagnetic equator (Abdu et al., 1981). An alignment between the magnetic field line and the solar terminator favors the conditions for the growth of instabilities close to the geomagnetic equator. In Fig. 7, it can be seen that this alignment is similar in both Jicamarca (left panel) and Alcântara (right panel). This feature allowed us to assume that the conditions encountered by the rocket during its flight were similar to those present during the VHF radar observation. Although the RTI map and the  $N_e$  height profile were obtained at longitudinally separated locations and different times, the comparison can be carried out. As discussed at the end of this section and noted by Kherani et al. (2012) and Woodman and Chau (2001), the irregularities that appear in the E and valley regions are closely related to bottom-type and top-side (plumes or bubbles) irregularities.

Because we did not have the original RTI map, it was decided to work with gray levels of the color-inverted image from the paper of Woodman and Chau (2001) and calculate the levels of grayscale integrated at fixed heights inside the shaded area in Fig. 8. The original RTI map presents the irregularities as a SNR, being the background noise the random signals generated by radio stars, i.e., the noise is not altitude dependent. The signal came from irregularities backscattering the waves emitted by the radar following the Bragg condition. In the case of Jicamarca, the scales of these irregularities are 3 m and can be assumed as tracers for intermediate- and large-scale structures in the plasma, whose boundaries are demarcated by the backscatter (Hysell, 2000). The SNR is represented in decibels, given the wide range of values that can have this ratio. Our representation



**Figure 8.** Rocket trajectory of the experiment conducted in December 2012 superimposed over the color-inverted RTI map taken at Jicamarca in October 1993 (reproduced from Woodman and Chau, 2001). We assume that the shadowed region in faint blue represents the possible trajectories that may have had the rocket in a physical scenario similar to that measured at Jicamarca.

is a linearization of the original SNR and can be considered as an approximate picture of the physical scenario present at the time of the observations studied in Woodman and Chau (2001). The aim of this procedure is to catch the spatial distribution (horizontal range or time in the RTI map; altitude) of the echoes (irregularities).

In our redefined grayscale of the original RTI map, a black color (gray level = 0) represents the noise or a region without any plasma irregularities that might eventually produce a backscatter signal. Conversely, white regions (gray level = 1) correspond to plasma irregularities and turbulent regions that provoke echoes measured by the radar. The area that was selected for the calculation of the integrated grayscale levels (faint blue in Fig. 8) represents all possible trajectories that may have shown the rocket in an ionosphere with a scenario similar to that reported by Woodman and Chau (2001) over Jicamarca.

The horizontal scale showed in Fig. 8 is the same used in Woodman and Chau (2001). In that work, the authors assumed an eastward plasma drift velocity of  $\sim 110 \text{ m s}^{-1}$ . By overlaying the trajectory of the rocket over the RTI map, we assume the same velocity (value and direction). In magnetically quiet conditions and at night, the plasma drifts eastward. Dst (disturbance storm time) geomagnetic index as well as all-sky images (Savio et al., 2016) from a station located southeast of Alcântara confirm that during the launch the magnetic activity was quiet. In relation to the value of the zonal drift velocity, different studies (Pimenta et al., 2003; de Paula and Hysell, 2004; Rodrigues et al., 2004) reveal that over the Brazilian sector the zonal drift velocity is higher ( $\sim 150 \text{ m s}^{-1}$ ) than that reported in Jicamarca. Taking this drift velocity difference into account would be equivalent to contracting the branches of the parabola on a fixed horizontal axis. This can surely contribute to the differences that may arise in our comparison.

The results of the integrated grayscale levels calculated at intervals of 5 km are shown in Fig. 9. It is worth noting that the rocket's trajectory was east-northeast and, on an RTI map, the east direction is normally located on the left side of the image. That is why the integrated grayscale profile corresponding to the rocket's upleg (left panel of Fig. 9) is calculated from the shadow region located on the right side of Fig. 8.

In order to compare Langmuir probe data with the RTI map generated from the echoes registered over Jicamarca, we calculate the  $N_e$  irregularity fluctuations ( $\delta N_e = \frac{|N_e - \langle N_e \rangle|}{\langle N_e \rangle}$ ). The  $N_e$  irregularity fluctuation profile defined in this way represents the ratio of the measured  $N_e$  profile to the smoothed  $N_e$  profile using an exponential moving average, or  $\langle N_e \rangle$ . We consider that the  $N_e$  profile measured by the probe is a superposition of the average ionosphere and the fluctuations that eventually the presence of different irregularities along the rocket's trajectory may cause.

As a purpose for our comparison,  $\delta N_e$  was considered as the in situ analogue of the SNR obtained by the radar observations in the sense that both data represent the spatial distribution of irregularities. In our in situ measurements we have such a distribution along the path of the rocket. This spatial distribution is observed in a narrow band as it passes above the radar, with a zonal plasma drift of  $\sim 110 \text{ m s}^{-1}$  in an eastward direction in the case of the VHF radar installed at Jicamarca. The main difference between these observational methods is that in the first case the instrument moves with such a velocity that it allows us to consider the plasma as being at rest while the probe intersects different irregularities along the trajectory of the rocket. In the second case, the irregularities appear in the radar's field of view as it moves in an eastward direction under quiet geomagnetic conditions.

The results appearing in Fig. 10 show a moving average of the  $N_e$  irregularity fluctuation profile ( $\delta N_e$ ) obtained in the manner explained above.

In order to make the comparison more straightforward, it was decided to overlap the Figs. 9 and 10. The result is shown in Fig. 11.

In spite of comparing data obtained with different instruments and in different locations and time periods, the similarity of the two profiles is quite prominent. During the upleg part of the flight (Fig. 11, left panel), the integrated SNR from radar data decreases between 100 and 120 km because the region measured by radar at Jicamarca is furthest from the equatorial electrojet. After this decrease, an absence of irregular structures appears in a range between 120 and 190 km in the original RTI map and as a consequence the integrated grayscale levels reach their minimum values. Meanwhile, successive depletions at  $\sim 130$ ,  $\sim 160$  and  $\sim 190$  km (Fig. 4a), each being less pronounced than the previous one, are represented as three successive peaks in the left panels of Figs. 10 and 11. The last pick or increase in the  $\delta N_e$  profile around 190 km has almost the same altitude as the first increase in the integrated grayscale levels, responding to the

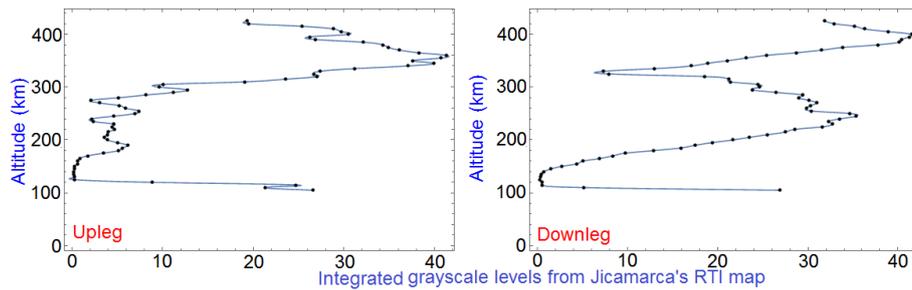
presence of the first visible striations in the original RTI map inside the region of grayscale level integration (Fig. 8). From this height and up to  $\sim 380$  km, the behavior of both profiles is very similar, with the integrated SNR profile shifted to higher altitudes by approximately 20 km in respect to the in situ profile. The lack of irregularities during the last 90 km in the rocket's upleg was well captured in the  $\delta N_e$  profile represented in Fig. 10 (left panel). Above 400 km this profile reached its minimum values.

The downleg part of the flight (Fig. 11, right panel) shows similarity in the height dependence of both profiles starting at 330 km until the end of the comparison. Above that altitude the Langmuir probe did not detect any irregularity. In the altitude interval from 200 to 300 km, a set of complex striations is visible on the RTI map (Fig. 8). The grayscale changes from clear to dark regions alternately and the small fluctuations on the integrated grayscale levels over these regions observed in Figs. 9 and 11 are a direct consequence of this feature. These fluctuations resemble the wave-like part of the  $N_e$  height profile detected by in situ measurements during the downleg; depicted in Fig. 4b. The  $\delta N_e$  height profile calculated from rocket data does not clearly show this behavior due to the smoothing process carried out to obtain the curves shown in Fig. 10. In this part of the profile, differing from the upleg part of the flight, the integrated grayscale levels are shifted to lower altitudes of about 20 to 60 km (depending on the altitude) in respect to the irregularity fluctuation profile from our in situ measurements.

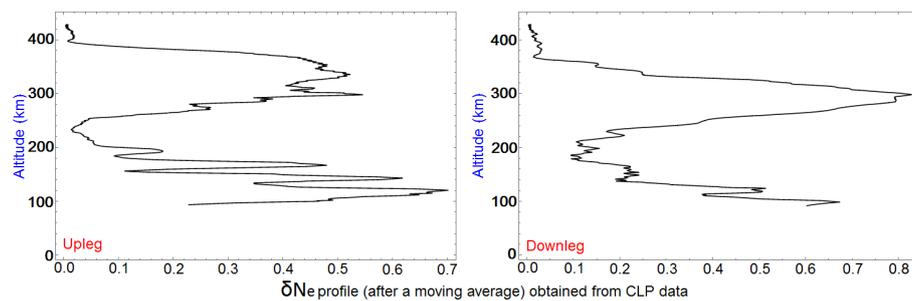
Although the Langmuir probe onboard the rocket did not intercept any large-scale irregularities that could be associated with a plasma bubble or a plume, the ionograms shown in Fig. 2 indicate that the irregularities found at altitudes below the base of the F layer appear simultaneously with irregular F structures of larger scale (bottom-type and/or bottom-side F layer and plumes). This feature was observed in previous studies carried out in Brazil (Kherani et al., 2012). In that work, it was noted that the slope of the valley-type irregularities, as observed on an RTI map, was proportional to the ascension rate of bottom-type and/or bottom-side F layer irregularities. This is easily recognizable from Fig. 8 for the case of Jicamarca.

According to Abdu et al. (2009) and Takahashi et al. (2010), the ascension rate of the bottom-type and/or bottom-side F layer is representative of the zonal electric field of F irregular structures under the modulation of wave-like disturbance. This leads to numerous studies (e.g., Alam Kherani et al., 2004; Mukherjee and Patra, 2014; Kherani and Patra, 2015) aiming to understand the physical origin of this valley-type plasma irregularity through the possible link with the fringe field (polarization field) of F irregularities.

The main result of these numerical simulations was that the polarization electric field can penetrate up to 120 km and find and transport the plasma lying there, provided that suitable conductivity profiles under strongly driven conditions allow this fringe field penetration (Alam Kherani et al.,



**Figure 9.** Integrated grayscale levels of the modified RTI map shown in Fig. 8. Irregular structures are represented by higher integrated grayscale levels.



**Figure 10.** Irregularity fluctuation profile ( $\delta N_e$ ) obtained from the data collected by the Langmuir probe onboard the rocket. Irregular structures are represented by higher values of  $\delta N_e$ .

2004). Frequently, type-II E region irregularities are located at this altitude in the evening. Therefore, these structures could be the plasma source that is transported via eastward electric field, giving rise to the valley echoes detected in this study.

More recently, Kherani and Patra (2015) showed that the fringe field associated with large-scale irregularities penetrates only within the bell of  $\pm 5^\circ$  from the magnetic equator. These features can explain why valley-type irregularities are not observed in the RTI maps from VHF radar observations at low latitudes (Patra and Venkateswara Rao, 2007; Mukherjee and Patra, 2014). Also, Kherani and Patra (2015) explained the electrodynamic origin of the asymmetries noted in the raising structures (Fig. 8) in the E and F valley regions. Previous observations (Woodman and Chau, 2001; Kherani et al., 2012) show that the eastern rise of valley irregularities is more prominent compared with its western counterpart. Our results (see the  $N_e$  height profiles in Fig. 4) confirm this fact.

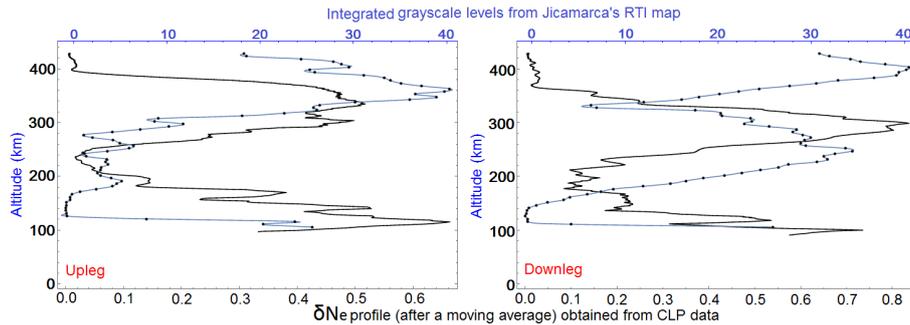
#### 4 Conclusions

The in situ detection of valley-type irregularities has been presented and discussed. Using observations from digisondes installed at equatorial (São Luís, dip latitude =  $6^\circ$  S) and low-latitude (Fortaleza, dip latitude =  $15^\circ$  S) stations we were able to affirm that large-scale irregularities were present at F region heights during the launch of a two-stage rocket that carried a Langmuir probe as a payload. During its trajectory,

the rocket was able to intercept irregular plasma structures of small and medium scales located in intermediate regions between the E and F layers. We needed an observational reference frame that would be able to explain our in situ observations in the best possible way considering the scarcity of data. The key feature in our experiment (different from other observations using instruments onboard a rocket) was the altitude at which these irregularities appeared and the quasi-periodic nature of the irregularities in a height range that has been previously detected in Jicamarca.

Knowing this, we compare the spatial distribution of irregularities as observed by two different instruments: a Langmuir probe installed as the payload of a rocket and a VHF radar operating at a different location but under the geomagnetic equator, similar to the rocket's trajectory. Therefore, we do not attempt a *pari passu* comparison between VHF radar observation and Langmuir probe measurements. In this work, we constructed altitudinal profiles of plasma irregularities with the data obtained by these two instruments and then compared these profiles.

The behavior of the irregularity fluctuation profile elaborated from data obtained by the aforementioned instruments was acceptably similar. The main difference comes from the fact that the probe onboard the rocket could not detect any plasma bubble at F region heights and the RTI map of Jicamarca shows the typical F irregular structures associated with a spread F event, e.g., plumes, bottom-type and bottom-side irregularities.



**Figure 11.** Irregularity fluctuation profile obtained from in situ measurements (black curve) together with the integrated SNR of the RTI maps from VHF radar measurements (blue curve and black dots).

The comparison leads us to conclude that the irregularities detected in this work are signatures of irregularities in the valley region previously labeled as Equatorial Quasi-Periodic. The nature of the irregularities cannot be determined only from the rocket measurements; thus, simultaneous radar measurements are needed in order to enhance the conclusion validity about the nature of the electron density structures observed in our rocket measurements.

Attempts made so far using numerical simulations have been able to explain many properties of the irregular structures detected in the present rocket experiment and also observed with VHF measurements. However, further numerical experiments need to be done to reproduce the quasi-periodic distribution of valley-type irregularities.

**Data availability.** The rocket data used in the present work are part of a very short database of in situ measurements. In order to be suitable for analysis, it demands an exhaustive step of data processing based on documentation as well as oral communications not available to the general public. Moreover, these data are currently in use by INPE's students in their PhD projects focused on other aspects of the experiment. An alternative for public accessibility of these data would also depend on a convenient scientific database that could receive them in the future. The ionograms shown in Figs. 2, 3 and 6 can be visualized in <http://www2.inpe.br/climaespacial/portal/ionosondes-home/> and requested by filling out the form below: <http://www2.inpe.br/climaespacial/SWMonitorUser/faces/register.xhtml>, from where it is also possible to request the parameters shown in Fig. 5.

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