

Magnetopause reconnection rate estimates for Jupiter's magnetosphere based on interplanetary measurements at ~ 5 AU

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Abstract. We make the first quantitative estimates of the magnetopause reconnection rate at Jupiter using extended *in situ* data sets, building on simple order of magnitude estimates made some thirty years ago by Brice and Ionannidis (1970) and Kennel and Coroniti (1975, 1977). The jovian low-latitude magnetopause (open flux production) reconnection voltage is estimated using the Jackman et al. (2004) algorithm, validated at Earth, previously applied to Saturn, and here adapted to Jupiter. The high-latitude (lobe) magnetopause reconnection voltage is similarly calculated using the related Gérard et al. (2005) algorithm, also previously used for Saturn. We employ data from the Ulysses spacecraft obtained during periods when it was located near 5 AU and within 5° of the ecliptic plane (January to June 1992, January to August 1998, and April to October 2004), along with data from the Cassini spacecraft obtained during the Jupiter flyby in 2000/2001. We include the effect of magnetospheric compression through dynamic pressure modulation, and also examine the effect of variations in the direction of Jupiter's magnetic axis throughout the jovian day and year. The intervals of data considered represent different phases in the solar cycle, such that we are also able to examine solar cycle dependency. The overall average low-latitude reconnection voltage is estimated to be ~ 230 kV, such that the average amount of open flux created over one solar rotation is ~ 500 GWb. We thus estimate the average time to replenish Jupiter's magnetotail, which contains ~ 300 – 500 GWb of open flux, to be ~ 15 – 25 days, corresponding to a tail length of ~ 3.8 – 6.5 AU. The average high-latitude reconnection voltage is estimated to be ~ 130 kV, associated with lobe "stirring". Within these averages, however, the estimated voltages undergo considerable variation. Generally, the low-latitude reconnection voltage exhibits a "background" of ~ 100 kV that is punctuated by one or two significant enhancement events during each solar rotation, in which the voltage is elevated to ~ 1 – 3 MV. The high-latitude voltages are estimated to be about a half of these values. We

note that the peak values of order a few MV are comparable to the potential drop due to sub-corotating plasma flows in the equatorial magnetosphere between $\sim 20 R_J$ and the magnetopause, such that during these periods magnetopause reconnection may have a significant effect on the otherwise rotationally dominated magnetosphere. Despite such variations during each solar rotation, however, the total amount of open flux produced during each solar rotation varies typically by less than $\sim 30\%$ on either side of the overall average for that epoch. The averages over individual data epochs vary over the solar cycle from ~ 600 GWb per solar rotation at solar maximum to ~ 400 GWb at solar minimum. In addition we show that the IMF sector with positive clock angle is favoured for reconnection when the jovian spin axis clock angle is also positive, and vice versa, although this effect represents a first order correction to the voltage, which is primarily modulated by IMF strength and direction.

Keywords. Magnetospheric physics (Electric fields; Planetary magnetospheres; Solar wind-magnetosphere interactions)

1 Introduction

It is well known that the dynamics of Jupiter's near-planet magnetosphere are dominated by the internal energy source of planetary rotation, rather than by the solar wind as is the case at the Earth, coupled with the production, transport, and loss of plasma from the highly productive Io source (e.g. Hill, 1979; Pontius, 1997; Vasyliunas 1983; Delamere and Bagenal, 2003). The tapping of the huge energy reservoir that is Jupiter's rotation is spectacularly manifest in the magnetosphere-ionosphere coupling current system that is associated with the breakdown of corotation of iogenic plasma and the formation of the main auroral oval (Cowley and Bunce, 2001; Hill 2001; Khurana, 2001; Southwood and Kivelson, 2001). This dominance of corotational flow over that of the Dungey cycle driven by the solar wind is traditionally illustrated by a comparison of the magnitudes

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of the equatorial electric fields associated with each flow, which for Jupiter are directed radially outward and dusk-dawn respectively (Brice and Ioannidis, 1970; Kennel and Coroniti, 1975, 1977). These authors used “typical” values of the solar wind parameters based on Pioneer-10 and -11 data to estimate a Dungey cycle voltage of ~ 1 MV, compared this with a corotation voltage of order ~ 400 MV, and concluded that the dominant regime was that of corotation.

To quote Kennel and Coroniti (1975) in this context, however, “one neglects reconnection at his peril”. The dominance of corotation does not preclude there being a significant solar wind interaction as well, as is manifest in the existence of a substantial magnetic tail that extends at least ~ 4 AU downstream of Jupiter, to the orbit of Saturn. For example, Lepping et al. (1983) presented evidence to suggest that Saturn was within Jupiter's tail during the Voyager-2 encounter. In addition to the existence of an extended magnetotail, HST images of Jupiter's aurora show features located at higher latitudes than the main auroral oval that have been interpreted as being due to magnetic reconnection occurring at the dayside magnetopause (e.g. Pallier and Prangé, 2001, 2003; Grodent et al., 2003, 2004; Bunce et al., 2004). In a recent study, Nichols and Cowley (2005) used a realistic middle magnetosphere current sheet field along with a theoretically computed equatorial plasma angular velocity profile, and showed that the vast majority of the above corotation potential falls within $\sim 20 R_J$ of the planet. They calculated the potential difference between $20 R_J$ and $60 R_J$ to be only ~ 4 MV. Cowley et al. (2005) recently expanded this work to include the outer magnetosphere region and computed the voltage between $\sim 20 R_J$ and the magnetopause to be ~ 10 MV. Thus, the magnetopause reconnection estimate of ~ 1 MV is clearly not negligible in this region. Reconnection-related flows may not therefore always be negligible within Jupiter's magnetosphere, though they will be confined to its outer regions. However, magnetopause reconnection is expected to be strongly modulated by the magnitude and orientation of the interplanetary magnetic field (IMF), such that one might expect associated modulation of these flows, as well as of the dynamics of the jovian cusp, and the tail regions (e.g. Acuña et al. 1983; Nishida and Maezawa, 1986). The rate of dayside reconnection may therefore have important implications for other tail processes, such as the loss of iogenic plasma down-tail in the Vasylunas cycle (Vasylunas, 1983).

Save for the simple order of magnitude estimates of Brice and Ioannidis (1970) and Kennel and Coroniti (1975, 1977), however, to date no systematic study of the magnitude of jovian magnetopause reconnection exists. In this paper we address this issue using the extended *in situ* data now available. Specifically, we examine Ulysses and Cassini data from intervals when the spacecraft were located near Jupiter's orbit near the ecliptic plane at ~ 5 AU, these intervals also allowing us to sample various phases of the solar cycle. We begin in Sect. 2 with a discussion of the data intervals used in this study, and follow in Sect. 3 with a description of the empirical formulas we employ to estimate the reconnection-

associated voltages. We then present our results in Sect. 4, followed by a discussion of the results and summary in Sect. 5.

2 Data intervals

The monitoring (or the lack thereof) of the interplanetary medium upstream of Jupiter is a continuing problem for planetary magnetospheric physicists. Without knowledge of the condition of the solar wind and IMF as it impinges on Jupiter's magnetosphere it is almost impossible to infer the effect of the former on the latter. However, there exists significant data concerning the interplanetary medium in the vicinity of Jupiter's orbit. In this paper we use data from the Ulysses and Cassini spacecraft, both of which have made flybys of Jupiter, in 1992 and 2000/2001 respectively. However, since 1992, Ulysses has completed two orbits over the poles of the Sun and has hence passed through the region of Jupiter's orbit twice more, in 1998 and 2004, though Jupiter itself was not then in the vicinity of the spacecraft. Thus, if one considers “Jupiter space” to be the region near Jupiter's orbit in which one may sample solar wind conditions that are similar to those which Jupiter itself experiences, we can also use two further intervals of data. The semi-major axis, eccentricity, and inclination of Jupiter's orbit are ~ 5.2 AU, ~ 0.048 , and $\sim 1.3^\circ$, respectively (Fränz and Harper, 2002). Hence Jupiter's heliocentric distance varies between ~ 5.0 and 5.4 AU, and its orbit is modestly inclined to the J2000 ecliptic plane. Here we define “Jupiter space” to be the region within 5° heliocentric latitude of the J2000 ecliptic plane, and with a heliocentric range between 5.0 and 5.4 AU. Ulysses' orbit is such that it reaches aphelion of ~ 5.4 AU near to 0° heliocentric latitude, hence the latter restriction on heliocentric range is automatically satisfied for the spacecraft. We therefore consider Ulysses data over the following periods – from January to June 1992, January to August 1998, and April to October 2004. Cassini data from November and December of 2000 were also employed, limited to this period due to the somewhat incomplete data coverage in the relevant region (discussed further below). In order to place these intervals in the context of the solar cycle, we show in Fig. 1 a plot of 25-day averaged sunspot number, obtained from the Solar Influences Data Analysis Centre (SIDC), Belgium, against time from 1991 to mid-2005, along with grey boxes which denote the intervals considered. It can be seen that the intervals cover various phases in the solar cycle, such that we may expect the nature of the solar wind to vary accordingly (e.g. McComas et al., 2001). Studies at Earth (e.g. Hapgood et al., 1991) have shown that the strength of the IMF does vary systematically over the solar cycle, with the magnitude of the field being 50% higher at solar maximum than at solar minimum. These variations should therefore be reflected in the reconnection voltages. The Ulysses 1992 interval samples solar maximum/beginning declining phase, and would be expected to be characterised, therefore, by high field strengths. The

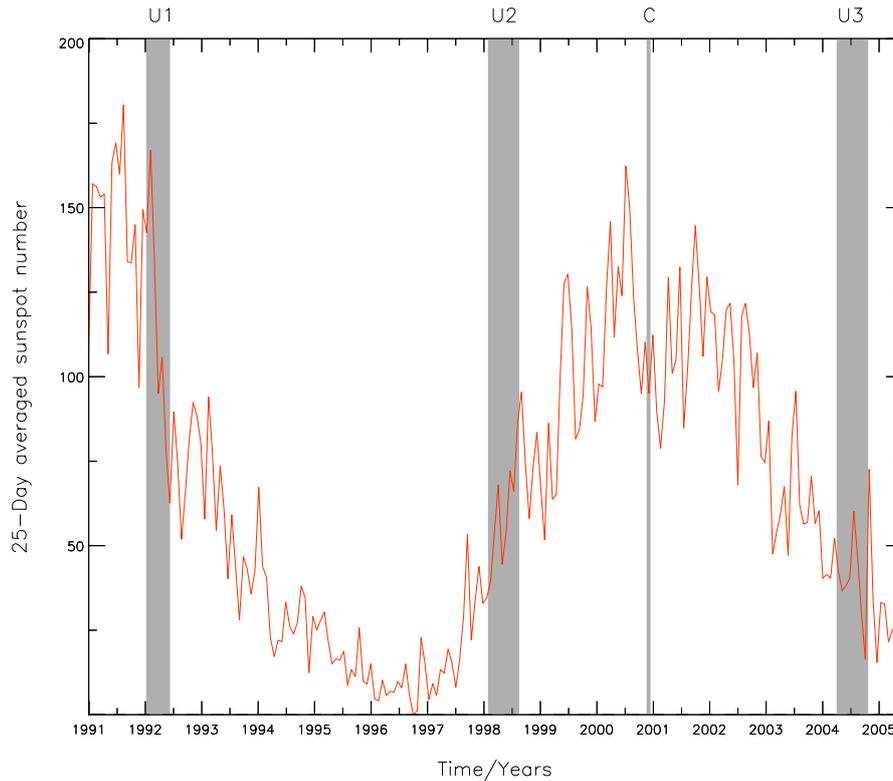


Fig. 1. Plot showing the 25-day averaged sunspot number versus time from 1991 to mid-2005. The grey boxes indicate the intervals of spacecraft data considered in this paper, marked U or C for Ulysses and Cassini respectively.

Ulysses 1998 interval lies in the middle of the rising phase, such that we would expect the field strengths to be moderately less than in the 1992 interval. The Cassini millennium flyby took place during solar maximum, such that, again, we may expect high field strengths. Finally, the Ulysses 2004 interval samples declining phase/solar minimum, such that the field strengths would be expected to be relatively low with respect to the other intervals. Overall, therefore, this study samples a range of solar wind conditions, such that we would expect to observe a corresponding variation in the reconnection voltage deduced.

3 Calculation of the reconnection voltage

In this section we will develop empirical formulas, based on terrestrial experience, which allow us to estimate the reconnection voltages at Jupiter. We concentrate mainly on the “low-latitude” reconnection voltage, ϕ_L , equal to the rate of open flux production, which is significant for the flux content and dynamics of the tail. We also consider, however, the “high-latitude” or lobe reconnection voltage, ϕ_H , which does not change the amount of open flux in the system if reconnection occurs at one lobe only for a given interplanetary field line, as will generally be the case. However, lobe reconnection does result in the circulation of flux within the open field region, and may influence high-latitude precipitation and auroral patterns.

The general formula for calculating the magnetopause reconnection voltage ϕ for any planet, equal to the rate of open flux production, is $\phi = v_{sw} B_{\perp} L$. Here v_{sw} is the velocity of the solar wind, B_{\perp} is the magnitude of the perpendicular component of the IMF (such that $v_{sw} B_{\perp}$ is the interplanetary electric field), and L is the width of the solar wind channel, perpendicular to B_{\perp} , for which reconnection occurs with the planetary field. As stated in the introduction, presently available estimates of the low-latitude jovian magnetopause reconnection voltage stem from the studies of Brice and Ioannidis (1970) and Kennel and Coroniti (1975, 1977), who assumed that $L = \beta L_M$, where L_M is the “width” of the magnetosphere, and β is a “reconnection efficiency parameter” assumed on the basis of terrestrial experience to be ~ 0.1 . For example, Kennel and Coroniti (1977) employed “canonical” values of $v_{sw} = 400 \text{ km s}^{-1}$ and $B_{\perp} = 1 \text{ nT}$ based on Pioneer-10 and -11 data, along with a “width” of the magnetosphere of $L_M = 100\pi R_J$, to obtain a value of $\sim 1 \text{ MV}$.

Today, thirty years later, the theory of reconnection is not yet developed to a stage where *ab initio* calculations of the jovian magnetopause reconnection voltage can be made. However, numerous studies at Earth using spacecraft and radar data have established clear empirical relationships between interplanetary parameters and reconnection voltage data that we can reasonably scale and apply to other planets. Here we employ an empirical formula for the voltage that was first developed and validated for the Earth (Perrault and

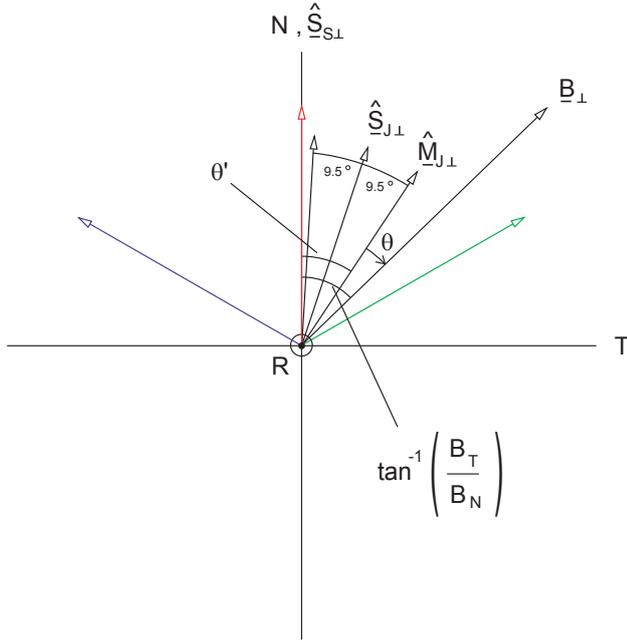


Fig. 2. Schematic diagram showing the vectors relevant to the calculation of the reconnection voltage, projected into the plane perpendicular to the Sun-Jupiter direction (i.e. the T-N plane). The Sun's spin vector $\hat{S}_{S\perp}$, also representing the N-axis, points vertically upwards, such that the T-axis points to the right, and the outward radial axis R points out of the page. Angles about the R-axis are measured clockwise from the N-axis. Jupiter's spin and magnetic axes are shown by the vectors labelled $\hat{S}_{J\perp}$ and $\hat{M}_{J\perp}$. As indicated, each jovian day $\hat{M}_{J\perp}$ undergoes an excursion about $\hat{S}_{J\perp}$ through an angle essentially equal to the $\sim 9.5^\circ$ offset between the magnetic axis and the spin axis. The maximum positive and negative excursions of $\hat{S}_{J\perp}$ as Jupiter orbits the Sun are shown (exaggerated) by the green and blue vectors respectively (the amplitude is actually $\sim 7.9^\circ$), while the red vector shows the case when $\hat{S}_{J\perp}$ is aligned with $\hat{S}_{S\perp}$, i.e. with the N-axis. The IMF direction is shown by the long vector labelled B_{\perp} . The clock angle of the field θ , i.e. the angle between $\hat{M}_{J\perp}$ and B_{\perp} , is taken to be positive clockwise as shown.

Akasofu, 1978; Milan et al., 2004), and was then adapted for use at Saturn by Jackman et al. (2004). This expression for the low-latitude dayside reconnection voltage, equal to the rate of open flux production, is given by

$$\phi_L = v_{sw} B_{\perp} L_o \cos^4(\theta/2), \quad (1)$$

such that the width of the solar wind channel in the above discussions is given by $L = L_o \cos^4(\theta/2)$, where θ is the “clock” angle between the IMF vector and Jupiter's magnetic axis projected onto a plane perpendicular to the Sun-Jupiter direction. Here we do not attempt to map the observed solar wind conditions to the planet, which would clearly be inappropriate for the 1998 and 2004 epochs, but simply calculate the voltage as though Jupiter was present at the spacecraft, as will be discussed further below. The calculation is, in principal, similar to that of Kennel and Coroniti (1975, 1977).

However, here we employ an extended time series of data rather than just “spot” values, and can thus examine the variations that occur over each solar rotation, from one rotation to another, and over the solar cycle.

With regard to the parameter L_o in Eq. (1), Milan et al. (2004) concluded that an appropriate empirical value for the Earth is $L_o \sim 5 R_E$, i.e. approximately half the subsolar radius of the planet's magnetosphere. This result was used by Jackman et al. (2004) to scale the formula for use at Saturn, and, in the absence of a model tailored specifically for the jovian case, the same logic is now applied here as a reasonable estimate for Jupiter. Here we therefore take L_o to be half the subsolar radius of Jupiter's magnetosphere R_{mp} , a value which is dependent on the dynamic pressure of the solar wind. Huddleston et al. (1998) give an empirical formula for this dependence, that is

$$R_{mp} = \frac{35.5}{p_{dyn}(\text{nPa})^{0.22}} R_J, \quad (2)$$

where $p_{dyn}(\text{nPa})$ is the dynamic pressure of the solar wind in nPa, given by

$$p_{dyn} = \rho v_{sw}^2, \quad (3)$$

where ρ is the mass density of the solar wind plasma, derived from the number density n_{sw} by multiplication by the mean solar wind plasma particle mass.

The function $\cos^4(\theta/2)$ represents the modulation of the rate of reconnection due to the direction of the IMF. It should be noted that the function is such that it is equal to 1 when $\theta=0^\circ$, 0.25 when $\theta=\pm 90^\circ$, and 0 when $\theta=180^\circ$, such that reconnection occurs most rapidly when the IMF is directed northward (recall that the polarity of Jupiter's magnetic field is opposite to that of the Earth's), and shuts off when the IMF is southward. Most commonly, however, the IMF lies near the equatorial plane, such that the prevailing clock angles are $\pm 90^\circ$, depending on the IMF sector. As with the L_o parameter, this function of θ represents a reasonable estimate for the jovian case, given experience at Earth and the absence of knowledge to the contrary at Jupiter.

Here we also employ a related formula for the high-latitude reconnection voltage. Specifically we calculate the high-latitude reconnection voltage using Gérard et al.'s (2005) adaptation of the Jackman et al. (2004) expression for low-latitude reconnection, given by

$$\phi_H = \frac{1}{2} v_{sw} B_{\perp} L_o \sin^4(\theta/2). \quad (4)$$

In this expression the sin function has replaced the cos dependence in Eq. (1), indicating its association with southward IMF, and the magnitude is estimated to be approximately half that of the low-latitude reconnection voltage (see Gérard et al. (2005) and Bunce et al. (2005) for further information).

We now consider the calculation of the “clock” angle θ from the magnetic data. The coordinate system in which the magnetometer data for both Ulysses and Cassini are rendered is the RTN system. This is defined such that R points radially

away from the Sun, N is perpendicular to this axis and is such that the R-N plane contains the Sun's spin axis vector, and T then forms the right-handed set, pointing generally in the direction of planetary motion. Figure 2 shows a schematic of the relevant vectors and angles projected onto the T-N plane (hence the “ \perp ” labels on the vectors), looking towards the Sun. Note that clock angles about the R-axis are measured clockwise from the N-axis. The Sun's projected spin axis $\hat{S}_{S\perp}$, which also represents the N-axis, points vertically upwards, such that the T-axis points to the right. Jupiter's spin axis projected into the T-N plane, $\hat{S}_{J\perp}$, then lies at some angle to $\hat{S}_{S\perp}$ dependent on Jupiter's orbital longitude. Over the course of Jupiter's orbit around the Sun, $\hat{S}_{J\perp}$ moves between angles of $\sim -7.9^\circ$ and $\sim +7.9^\circ$ relative to the N-axis, represented in Fig. 2 by the blue and green vectors respectively (the effect being considerably exaggerated in the diagram for sake of clarity). The red vector then represents the case where $\hat{S}_{S\perp}$ and $\hat{S}_{J\perp}$ are co-aligned along the N-axis in the T-N plane. Jupiter's magnetic axis is offset from the spin axis by $\sim 9.5^\circ$, such that it oscillates around the spin axis by this amount over the course of a jovian day. When projected onto the T-N plane, this offset angle is only modified in the second decimal place for all spin axis angles, such that, as indicated in the figure, we just take the angle θ' between the Sun's projected spin vector $\hat{S}_{S\perp}$ and Jupiter's magnetic axis vector $\hat{M}_{J\perp}$ to be the spin axis clock angle plus and minus 9.5° . The clock angle of the field, i.e. the angle between B_\perp and $\hat{M}_{J\perp}$, as shown in Fig. 2, is thus given by

$$\theta = \tan^{-1} \left(\frac{B_T}{B_N} \right) - \theta', \quad (5)$$

where B_T and B_N are the T and N components of the IMF. We also have in terms of these components

$$B_\perp = \sqrt{B_T^2 + B_N^2}. \quad (6)$$

From any given IMF data point we thus derive a range of voltages due both to the effect of Jupiter's orbital motion, shown in Fig. 2 by the red, green, and blue vectors, and its rotation, shown in the figure by the range of $\hat{M}_{J\perp}$ around $\hat{S}_{J\perp}$. The effect of the direction of $\hat{M}_{J\perp}$ on the derived reconnection rates can be readily understood from Fig. 2. Suppose, for example, the angle of $\hat{M}_{J\perp}$ relative to the N-axis is negative, towards the direction of the blue vector. Then if B_\perp points in the $-T$ direction it will have a component northward relative to $\hat{M}_{J\perp}$ and reconnection will be favoured. If, on the other hand, B_\perp points in the $+T$ direction it will have a southward component relative to $\hat{M}_{J\perp}$, and reconnection will be disfavoured. The opposite applies when $\hat{M}_{J\perp}$ points toward the green vector, of course. We should note, however, that these spin and magnetic axis effects represent only first order corrections to the voltage, which, as will be seen, is dominated by the large variations in the IMF magnitude and direction at Jupiter.

The quantities thus required from the spacecraft are the B_T and B_N components of the IMF obtained from the Ulysses FGM/VHM (Balogh et al., 1992) and Cassini MAG

(Dougherty et al., 2004) magnetometers, and n_{sw} and v_{sw} , obtained from the Ulysses SWOOPS (Bame et al., 1992) and Cassini CAPS (Young et al., 2004) plasma particle detectors. We have averaged these data over ten minute intervals, and linearly interpolated between data gaps. This latter comment is most pertinent to the Cassini CAPS data, as the instrument was limited in its ability to make observations during the Jupiter flyby due to spacecraft pointing restrictions.

4 Results

Each interval of interplanetary data analysed here covers a number of months, such that it is not feasible to reproduce the data set in its entirety. Instead we give in Table 1 the values of the parameters occurring in Eq. (1) averaged over each solar rotation and each epoch, i.e. B_\perp , $\cos^4(\theta/2)$, v_{sw} , $R_{mp}(=2L)$, and ϕ_L , along with ϕ_H and the total open flux produced during each solar rotation Φ . For the Ulysses data one solar rotation is chosen to be the interval between two successive zeroes of Carrington longitude (approximately 25 Earth days), while for Cassini we chose the 25-day interval that had the most continuous data coverage from the magnetic field and plasma instruments. In Figs. 3–6 we also show one solar rotation of typical data from 1992, 1998, 2000, and 2004, respectively. The upper panels show the “data ingredients” of the calculated parameters, while the lower panels show the results. Thus, from top to bottom we have B_\perp in nT from Eq. (6), θ in degrees from Eq. (5), $\cos^4(\theta/2)$, v_{sw} in km s^{-1} , n_{sw} in cm^{-3} , p_{dyn} in nPa from Eq. (3) (note the log scale in the panels for these last two quantities), R_{mp} in R_J from Eq. (2), ϕ_L in kV from Eq. (1), ϕ_H in kV from Eq. (4), and finally Φ in GWb obtained by integrating ϕ_L over time from zero at the start of each interval. Parameters that depend on the IMF clock angle θ , i.e. θ itself, $\cos^4(\theta/2)$, ϕ_L , ϕ_H and Φ , are shown with the jovian spin and magnetic axis variations described above taken into account. That is, each “data point” is stretched into a vertical bar whose extent represents the range of voltage values due to the $\pm 9.5^\circ$ diurnal variation of Jupiter's magnetic axis vector. In addition, the orbital effect is included using the same colour scheme as used in Fig. 2. Thus the blue trace represents the voltage for the maximum negative angle of Jupiter's spin axis relative to the N-axis, green represents the voltage for the maximum positive angle, and red, which is plotted over the top of the blue and green traces where they overlap, represents the case where Jupiter's spin axis is aligned with the N-axis (and thus with the Sun's spin axis in the T-N plane). We note that the averaged values quoted below and in Table 1 correspond to the middle of the red trace.

We first consider the 1992 Ulysses data which represents solar maximum/beginning declining phase conditions. The interval we examine spans days 008–161, although we omit one solar rotation of data containing the magnetosphere flyby, such that we consider five solar rotations of data. An example of this data is shown in Fig. 3 covering the solar rotation during days 085–110. The features in the magnetic

Table 1. Table summarising the parameters occurring in Eqs. (1) and (4) for the low- and high-latitude magnetopause reconnection voltages over the intervals examined. Each value is the average for the particular solar rotation, except Φ which is the total amount of open flux produced over the rotation period (i.e. the time integral of ϕ_L). The averages over each epoch are shown in bold. The intervals in italics are those used as illustrative examples in this paper. Note that one solar rotation is omitted from the 1992 Ulysses data, corresponding to the jovian magnetosphere flyby.

Epoch	Days	B_{\perp} / nT	$\cos^4(\theta/2)$	v_{sw} / km s ⁻¹	$R_{mp}(=2L) / R_J$	ϕ_L / kV	ϕ_H / kV	Φ / GWb
1992 Ulysses	008-033	1.21	0.38	502	58	372	248	801
	059-085	1.04	0.34	474	62	293	173	643
	<i>085-110</i>	<i>0.68</i>	<i>0.35</i>	<i>436</i>	<i>63</i>	<i>265</i>	<i>93</i>	<i>582</i>
	110-135	0.99	0.33	416	63	260	156	571
	135-161	1.03	0.28	429	59	236	176	519
	Average	1.00	0.30	461	60	279	166	611
1998 Ulysses	024-050	0.53	0.34	375	65	144	68	316
	050-075	0.72	0.36	373	69	240	90	528
	<i>075-101</i>	<i>0.82</i>	<i>0.30</i>	<i>373</i>	<i>66</i>	<i>249</i>	<i>99</i>	<i>546</i>
	101-126	0.71	0.32	411	68	224	113	493
	126-151	0.70	0.32	493	69	262	125	575
	151-177	0.95	0.32	435	63	306	154	672
	177-202	0.61	0.37	467	76	228	98	501
	202-228	0.83	0.31	415	68	215	167	471
	Average	0.73	0.33	418	68	234	114	513
2000 Cassini	323-348	0.83	0.30	489	59	223	128	483
2004 Ulysses	092-118	0.64	0.30	468	74	212	113	466
	118-143	0.63	0.30	422	73	180	105	396
	143-168	0.57	0.30	426	72	161	92	353
	<i>168-194</i>	<i>0.58</i>	<i>0.32</i>	<i>435</i>	<i>71</i>	<i>197</i>	<i>84</i>	<i>433</i>
	194-219	0.61	0.27	445	72	162	130	355
	219-245	0.66	0.33	471	72	151	166	331
	245-270	0.74	0.38	458	74	265	137	583
	270-295	0.48	0.38	386	72	174	62	382
	Average	0.61	0.32	439	73	188	111	412

field and plasma data are generally consistent with the occurrence of corotating interaction regions (CIRs), which are known to be a dominant feature of the solar wind in the vicinity of Jupiter's orbit (e.g. Gosling and Pizzo, 1999; Gazis, 2000; Kunow, 2001), and are associated with heliospheric current sheet (HCS) crossings. It can be seen in the clock angle panel that in this interval the IMF exhibits essentially a two-sector structure (the sectors being identified by their respective $\pm 90^\circ$ prevailing clock angles), with a crossing of the HCS occurring somewhere between day 90 and day 97, during which interval of generally elevated field strengths there is seen to be considerable structure in the magnetic field and plasma parameters. The clock angle of the IMF is a major factor in the modulation of the rate of magnetopause reconnection. This dependency is illustrated in the third panel, which shows the factor $\cos^4(\theta/2)$ that appears in Eq. (1) and varies between one and zero depending on the IMF direction, as discussed above. The relative values of the different colours (spin axis orientations) will be discussed below. As indicated above, the HCS crossing is accompanied by a corresponding enhancement in the field magnitude to ~ 2 nT between days 90–97. During this period the plasma number density shown in the fifth panel is highly variable, ranging between ~ 0.05 and ~ 2 cm⁻³, indicating significant embedded structure within the compression region. The re-

gion of enhanced field is then followed by a rarefaction region, in which the field strength is low (a few tenths of a nT), the plasma velocity steadily decreases over time from ~ 500 km s⁻¹ to ~ 400 km s⁻¹, and the number density is relatively stable at ~ 0.2 cm⁻³. A second HCS crossing appears to occur on day 108, accompanied by a second interval of enhanced field strengths that begins approximately one day later, the beginning of which can be just seen at the end of the interval shown. The dynamic pressure, and hence the estimated subsolar magnetopause distance, shown in the sixth and seventh panels, reflect the large changes in the number density, such that the subsolar magnetopause distance ranges between ~ 40 and $\sim 100 R_J$ in the vicinity of the first HCS crossing during days 90–97, and is reasonably constant at $\sim 60 R_J$ during the rarefaction region.

The bottom three panels of Fig. 3, which are the most important in terms of this paper, then show the reconnection voltages calculated from these parameters using Eqs. (1) and (4), and the amount of open flux produced since the start of the interval. The low-latitude reconnection voltage that occurs in the rarefaction region has a “background” value of the order ~ 100 kV, consistent with a field strength of ~ 0.2 nT, a plasma velocity of ~ 450 km s⁻¹, a subsolar magnetopause distance of $\sim 60 R_J$, and a prevailing IMF clock angle of 90° . During this solar rotation, however, there was one

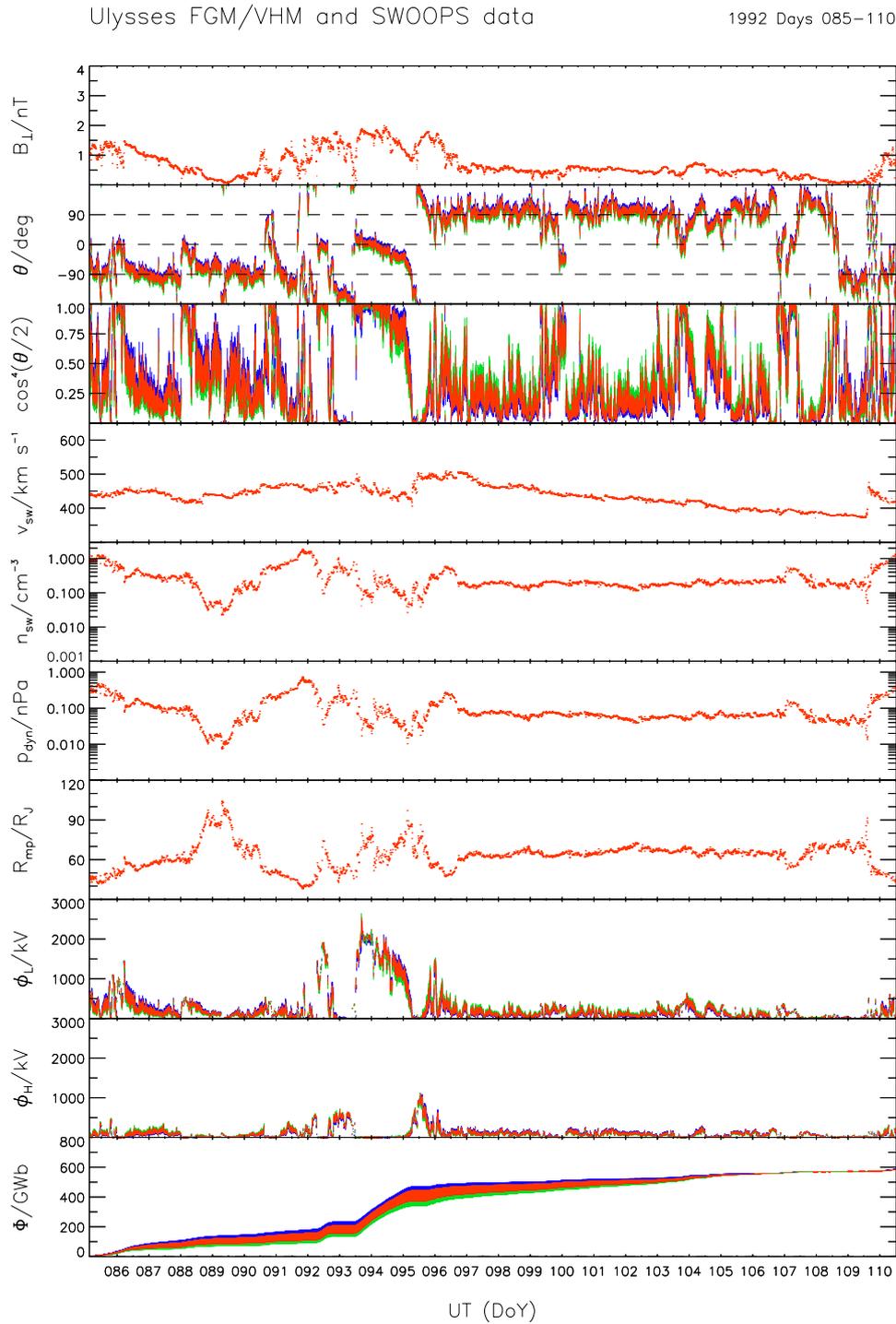


Fig. 3. Plot showing the “ingredients” of the models for the reconnection voltages from Ulysses data over days 85–110 in 1992, together with the estimated voltages and integrated open flux production. From top to bottom the panels show B_{\perp} in nT given by Eq. (6), θ in degrees given by Eq. (5), v_{sw} in km s^{-1} , n_{sw} in cm^{-3} on a log scale, p_{dyn} given by Eq. (3) in nPa on a log scale, R_{mp} in R_J given by Eq. (2), ϕ_L in kV given by Eq. (1), ϕ_H in kV given by Eq. (4), and finally Φ in GWb calculated by integrating ϕ_L over time, starting from zero at the initial time. In the plots for θ , ϕ_L , ϕ_H , and Φ each “data point” is stretched into a vertical line representing the effect produced by the $\sim\pm 9.5^\circ$ diurnal variation of Jupiter’s dipole axis offset. In addition, the blue, green, and red traces represent differing orientations of Jupiter’s spin axis relative to the Sun’s spin axis during the jovian orbit, as shown in Fig. 2.

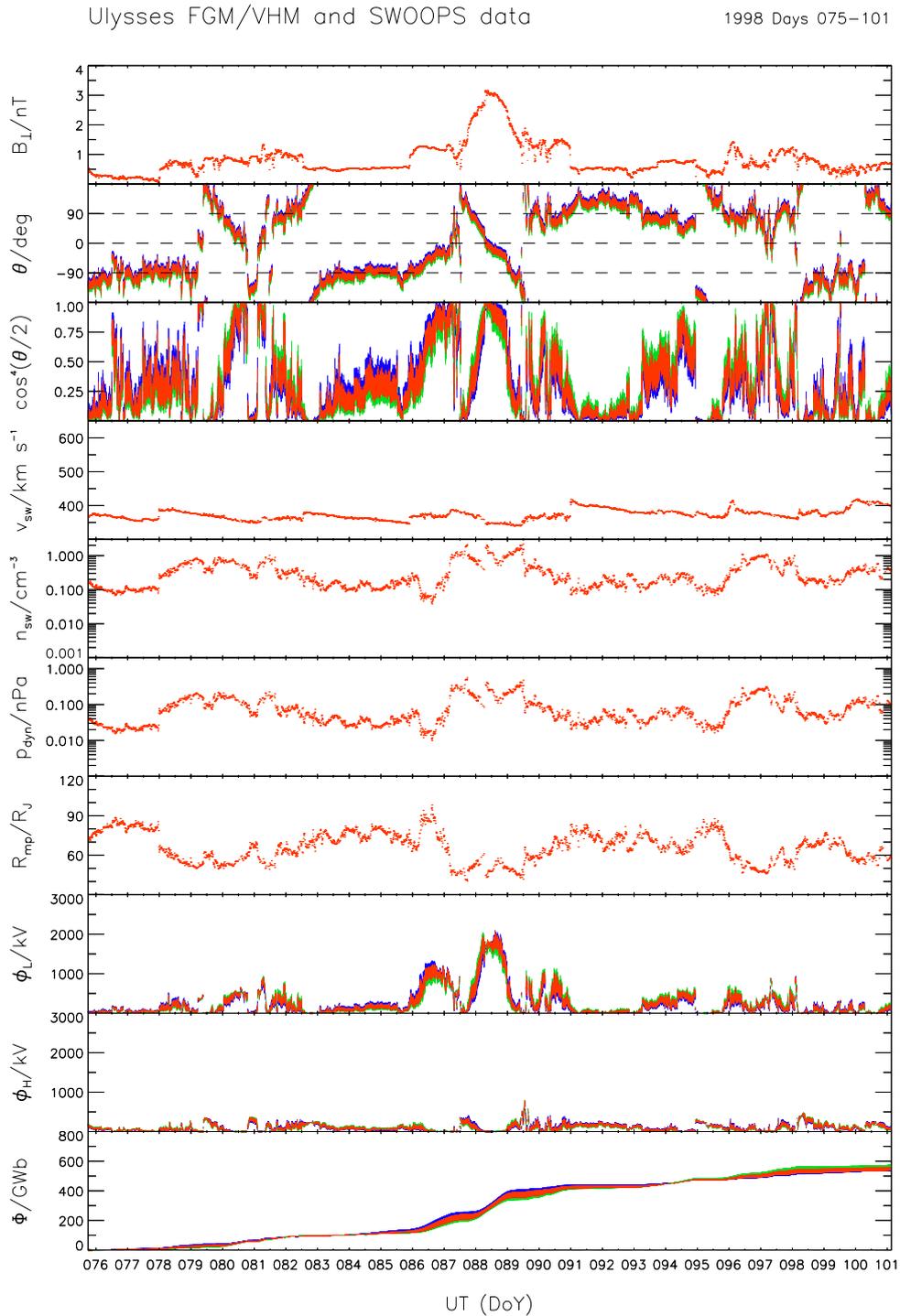


Fig. 4. As Fig. 3 for Ulysses data from the interval over days 75–101 in 1998.

significant period of enhanced voltage values associated with the region of high strength, often northward, IMF between days 90 and 97, in which the estimated low-latitude voltage peaked at ~ 2.4 MV. In this period the high-latitude voltage also peaked at ~ 1 MV where the IMF briefly turned southward. It can be seen in the bottom panel that the total open flux estimated to have been produced over this solar rotation

is ~ 580 GWb, corresponding to an average low-latitude voltage of ~ 265 kV over the interval. However, it should be noted that ~ 350 GWb, over half of the total estimated to have been produced over the solar rotation, is produced between days 90 and 97, ~ 200 GWb of which is created during days 93 and 94 alone.

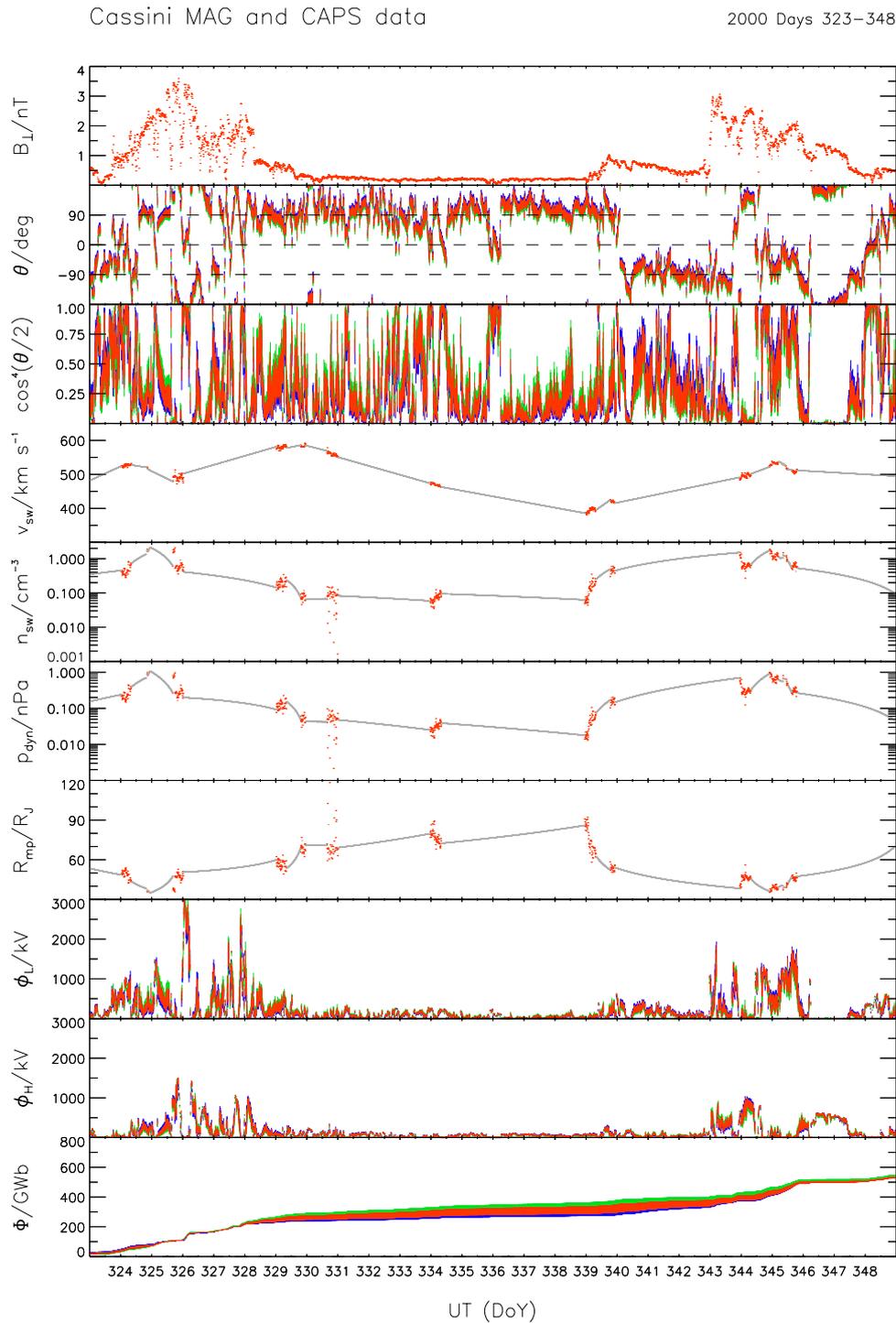


Fig. 5. As Fig. 3 for Cassini data from the interval over days 323–348 in 2000.

With regard to the spin and magnetic axis effects on the rate of reconnection, the former can be seen in the difference between the variously coloured data in the θ , $\cos^4(\theta/2)$, ϕ_L , ϕ_H and Φ panels, while the latter can be seen in the vertical extent of the data bars (although as the red bars are plotted over the blue and green where they overlap, only the full width of the red trace can be seen). Thus during

the predominantly negative B_T sector ($\theta \simeq -90^\circ$) before day ~ 90 , the blue trace, representing the case where the spin axis is at its maximum negative angle to the N-axis (see Fig. 2), exhibits higher $\cos^4(\theta/2)$, and thus low-latitude voltage, than the red and green traces. Conversely, during the predominantly positive B_T sector ($\theta \simeq +90^\circ$) between days ~ 97 and ~ 107 , the green trace, representing the maximum positive

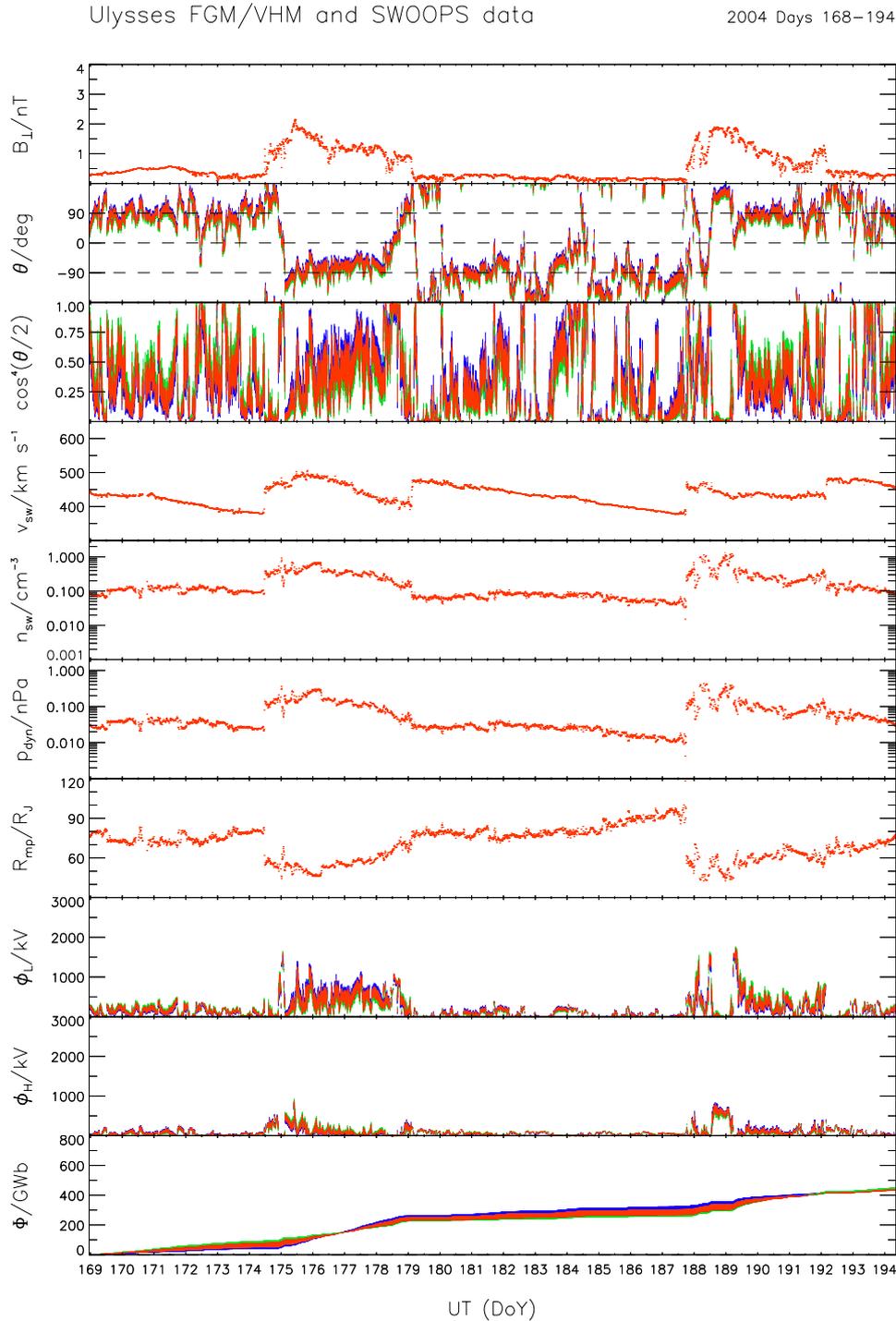


Fig. 6. As Fig. 3 for Ulysses data from the interval over days 168–194 in 2004.

angle of the spin axis to the N-axis, exhibits the higher low-latitude voltage, although as stated previously these variations are relatively small with respect to the overall modulation due to the changes in IMF strength and direction. However, whilst the above variations may be relatively small, the cumulative effect of the initial sustained period of negative IMF clock angle can be seen in the Φ panel by the gradual

separation of the coloured data and the widening of the trace until day 97, with the blue line higher than the green. However, the reverse is the case in the following sector, such that the coloured lines then re-converge. Thus, if over the solar rotation as a whole there occurs approximately equal intervals of negative and positive B_T sectors the total amount of open flux created is not greatly affected by the direction

of Jupiter's spin axis. We should note that during the period shown in Fig. 3, Jupiter was in fact near to the spacecraft following the Jupiter flyby, and had an orbital longitude of $\sim 160^\circ$. Thus the spin axis had an angle of $\sim +3^\circ$ relative to the N-axis, such that the actual values appropriate to the period lie roughly mid-way between the red and green traces. As a summary we show in Table 1 the values of the parameters that occur in Eq. (1) averaged over each solar rotation during this data epoch. The average amount of flux estimated to be created per solar rotation during the 1992 interval is ~ 610 GWb, which corresponds to an average low-latitude voltage of ~ 280 kV, with the values ranging between ~ 520 GWb and ~ 800 GWb per rotation, as shown. The averaged high-latitude voltages similarly vary between ~ 90 and ~ 250 kV, with an overall average of ~ 170 kV.

Turning now to the 1998 Ulysses data, representing the rising phase of the solar cycle, we have examined eight solar rotations of data spanning days 024–228. In Fig. 4 we show an example for the solar rotation during days 75–101. These data clearly suggest rather different solar wind conditions to those seen in 1992. The clock angle of the IMF switches often between $\sim \pm 90^\circ$, such that six “sectors” can be distinguished, and no large-amplitude stream interactions are evident. Instead the key event in terms of low-latitude magnetopause reconnection during this interval is the flux rope that passes over the spacecraft between $\sim 12:00$ UT on day 87 and $\sim 12:00$ UT on day 90, characterised by an enhancement in the magnetic field to ~ 3 nT and a smooth 360° rotation in the field direction. The average plasma velocity is ~ 370 km s $^{-1}$, lower than the ~ 435 km s $^{-1}$ average during the 1992 interval shown in Fig. 3, and the range of velocities of $\sim \pm 25$ km s $^{-1}$ is also less than the ~ 140 km s $^{-1}$ range in Fig. 3. The average number density, ~ 0.3 cm $^{-3}$, is similar to that in Fig. 3, and the range of variability is also similar. The low-latitude dayside reconnection voltage again exhibits a “background” of the order of a hundred kV, which is punctuated by an enhancement event associated with the flux rope and surrounding structure. During this event the estimated low-latitude voltage is elevated to a peak of ~ 2 MV, and ~ 150 GWb of open flux is estimated to be created during the two day period from 12:00 UT on day 87 to 12:00 UT on day 90. Over the whole ~ 5 day period between days 86 and 91, ~ 300 GWb is estimated to be created, i.e. over half of the ~ 550 GWb estimated to have been produced over the solar rotation as a whole. The high-latitude reconnection voltage does not exhibit any significant enhancement since, although the IMF briefly turned southward twice as the flux rope passed over the spacecraft, the field magnitude is not as high during these times as it was during the period of northward IMF. The parameter values occurring in Eq. (1), averaged throughout each of the eight solar rotations examined and the whole 1998 epoch, are again shown in Table 1. The average amount of open flux that is created over the whole 1998 interval is ~ 510 GWb per solar rotation (corresponding to an average low-latitude voltage of ~ 230 kV), with values ranging between ~ 320 GWb and ~ 670 GWb. The averaged high-latitude voltage is ~ 110 kV.

Figure 5 shows Cassini data between days 323 and 348 in 2000, which represent solar maximum conditions. Both magnetic field and plasma data during the Jupiter flyby period are, unfortunately, not continuous, such that the interval selected represents the most complete 25-day interval in terms of coverage from both magnetic field and plasma instruments together. The grey lines in the plasma data show the linear interpolation between the gaps in the plasma values. Again, the colours represent the variations due to Jupiter's spin and magnetic axis variations, as above, but in this case Jupiter was in the vicinity of the spacecraft when these data were taken and had a spin axis angle of $\sim +7^\circ$ relative to the N-axis. Therefore the green trace is the one appropriate to actual conditions at Jupiter in this case. A clear compression-rarefaction structure can be now seen in the field magnitude and plasma velocity data, although the clock angle data are highly variable around the HCS crossings. The first compression region, which lasts from near the start of the interval shown to day ~ 330 , contains large field magnitudes, up to ~ 3.5 nT, and is accompanied by a corresponding large enhancement in the estimated voltages, up to ~ 3 MV and ~ 1.5 MV for low- and high-latitudes, respectively. There follows a rarefaction region of low field strength, of order a few tenths of a nT, and decreasing plasma velocity from ~ 580 km s $^{-1}$ to ~ 400 km s $^{-1}$, associated with low voltages of order a hundred kV. A second, minor compression region begins on day 339 associated with the HCS crossing a day later. A region of enhanced field follows, beginning on day ~ 343 , in which the estimated low- and high-latitude voltages are elevated to ~ 2 MV and ~ 1 MV, respectively. Over this interval approximately ~ 480 GWb of open flux is estimated to have been produced, corresponding to an average low-latitude voltage of ~ 220 kV. The average high-latitude voltage is ~ 130 kV.

We finally consider the 2004 Ulysses interval, representing declining phase/solar minimum conditions, in which we examine eight solar rotations of data spanning days 092–295. In Fig. 6 we show an example for the solar rotation during days 168–194. The IMF clock angle indicates a two sector structure, with associated compression and rarefaction regions. Again the rarefaction regions are associated with a low “background” voltage, while the compressions are associated with enhancements. The maximum estimated voltages in the compression regions are relatively modest in this example, ~ 1.6 MV and ~ 0.8 MV for low- and high-latitudes respectively, reflecting the lower field strength than in the other intervals examined. The amount of open flux estimated to have been produced during this solar rotation is ~ 430 GWb, corresponding to an average low-latitude voltage of ~ 200 kV. The results for the other solar rotations in the 2004 interval are shown in Table 1, where it can be seen the average amount of open flux created per rotation is ~ 410 GWb, corresponding to an average low-latitude voltage of ~ 190 kV. The average high-latitude voltage is ~ 110 kV.

With regard to the distribution of voltage values, we show in Fig. 7 a normalised distribution plot of the low-latitude reconnection voltage (specifically the middle values in the red

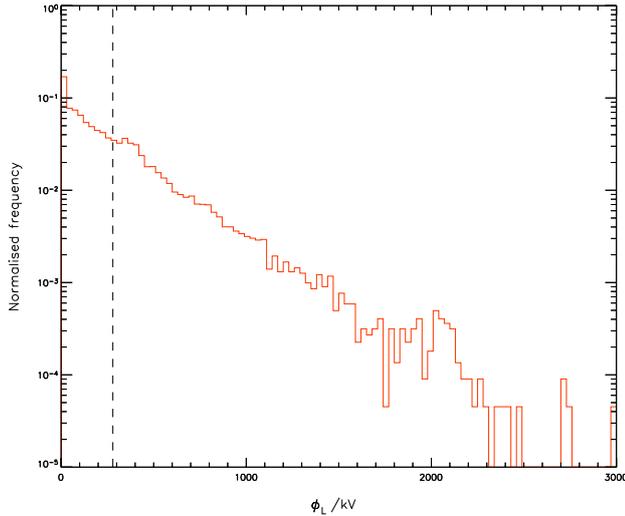


Fig. 7. Plot of the normalised distribution of low-latitude reconnection voltages derived from the entire 1992 Ulysses data set over days 008–161, shown on a linear-log scale. The values used are those from the middle of the red trace, i.e. the case when $\hat{S}_{S\perp}$ (i.e. the N-axis) and $\hat{M}_{J\perp}$ are co-aligned in the T-N plane. The vertical dashed line shows the location of the distribution mean.

trace) computed using the entire 1992 Ulysses data set over days 008–161 (omitting days 33–59), shown on a linear-log scale. The dashed vertical line shows the distribution mean of 279 kV. It can be seen that, despite the above discussions that might imply a “bimodal” distribution of voltages due to the prevalent background-enhancement region structure, the distribution is almost linear up to a voltage of ~ 2 MV, implying that the probability simply depends exponentially on the reconnection voltage. The equivalent plots for the 1998 and 2004 data sets are essentially similar, except that the slopes are modestly steeper, reflecting the lower average voltage values during these intervals.

5 Discussion and summary

In this paper we have estimated the low- and high-latitude jovian magnetopause reconnection voltages using an extended interplanetary data set at ~ 5 AU and empirical formulas that have been validated at Earth and previously applied to Saturn. We have employed Ulysses and Cassini data from 22 solar rotations when the spacecraft were in the vicinity of Jupiter’s orbit, and obtained time series of the reconnection voltages. We found that the average low- and high-latitude magnetopause reconnection voltages at Jupiter are of order ~ 230 kV and ~ 130 kV, respectively. This value of the low-latitude reconnection voltage is approximately ten times that derived for Saturn by Jackman et al. (2004), yet is a factor of ~ 4 less than the 1 MV estimate for Jupiter of Kennel and Coroniti (1977). However, a considerable amount of variation occurs within this average value. The “background” voltage of order ~ 100 kV, associated with

rarefaction regions, is generally punctuated by one or two events per solar rotation in which the magnetopause reconnection voltage is significantly elevated above the background. These events, in which the low-latitude voltage is enhanced to ~ 1 – 3 MV, are due to periods of high strength and/or northward IMF in compression regions and flux ropes. The high-latitude voltage is similarly found to be enhanced, to ~ 0.5 – 1.5 MV, when the IMF turns southward during these periods. Although the voltage across the magnetosphere as a whole associated with plasma corotation is ~ 400 MV, as previously estimated by Brice and Ionannidis (1970) and Kennel and Coroniti (1975, 1977), the voltage across the middle magnetosphere current sheet region is much less than this. As stated in the introduction, Nichols and Cowley (2005) and Cowley et al. (2005) recently estimated the voltage due to sub-corotation between $\sim 20 R_J$ in the equatorial plane and the magnetopause to be ~ 10 MV. Thus the Dungey cycle voltage, typically of order a few MV during enhancement events, is clearly not negligible compared with the corotation voltage across the middle and outer magnetospheres. The implication of these large voltages is that during enhancement events magnetopause reconnection may have a significant effect on an otherwise rotationally dominated magnetosphere. In addition, the probability distribution of reconnection voltage values appears to be exponentially dependant on the voltage, despite the apparent “bimodal” behaviour due to the rarefaction-enhancement region structure. We also considered the effect of the angle of Jupiter’s spin and magnetic axes on the voltage values, as this varies on the time scales of both a jovian day and a jovian year. We found that when Jupiter’s magnetic axis lies at a positive angle to the N-axis (see Fig. 2), the IMF sector with positive B_T exhibits higher estimated reconnection voltages than the sector with negative B_T , and vice versa. This effect, however, represents only a first order correction to the reconnection voltage, which is primarily modulated by the strength and direction of the IMF.

We also found that the average amount of open flux produced per solar rotation is ~ 500 GWb, with individual intervals during a given epoch varying by factors of less than $\sim \pm 30\%$, despite the interplanetary conditions varying considerably from interval to interval. However, the average value is found to depend on the solar cycle, ranging from an average of ~ 600 GWb per solar rotation in 1992 near solar maximum to ~ 400 GWb per solar rotation in 2004 near solar minimum. This result is consistent with that of Hapgood et al. (1991), who showed that the magnitude of the IMF is 50% higher at solar maximum than solar minimum at distances of ~ 1 AU. The amount of open flux in each lobe of Jupiter’s tail is estimated to be ~ 300 – 500 GWb, corresponding to a field of ~ 1 – 1.5 nT in a tail of radius $200 R_J$ (e.g. Acuña et al., 1983). Thus, we estimate that it takes ~ 15 – 25 days to produce enough open flux to replenish the tail. Following Dungey (1965), we multiply this time by a representative value of the solar wind speed, e.g. 450 km s^{-1} , to obtain an estimate of the length of the magnetotail of ~ 3.8 – 6.5 AU. This result is larger than that calculated by Kennel and Coroniti (1977) due to our smaller estimate of the

average reconnection voltage, but our value does allow Saturn to pass through Jupiter's tail, and, like our voltage estimate itself, is expected to contain considerable variation on the time scale of a solar rotation.

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